Symmetry energy, neutron skin and neutron star radius constraints from chiral EFT interactions

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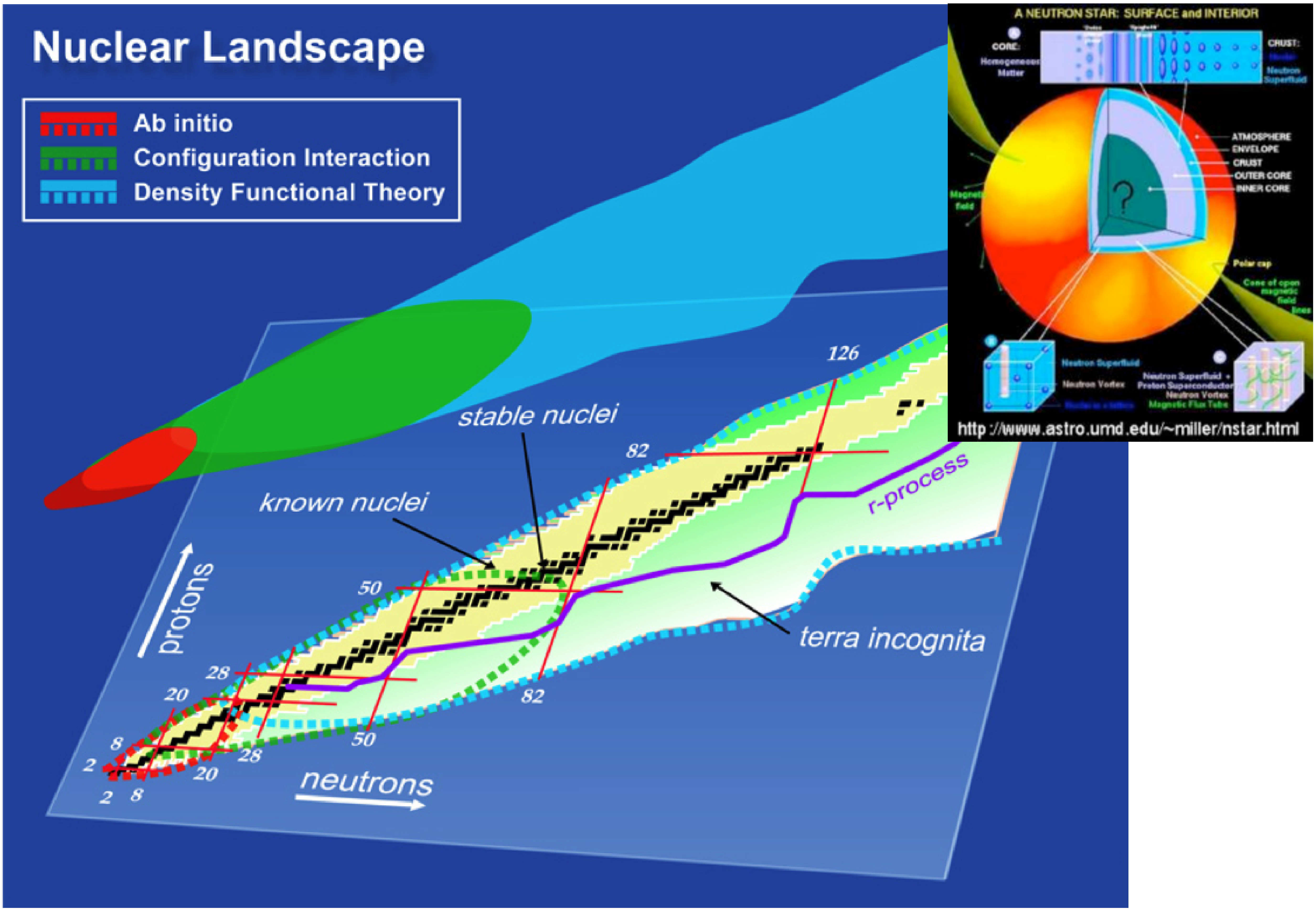
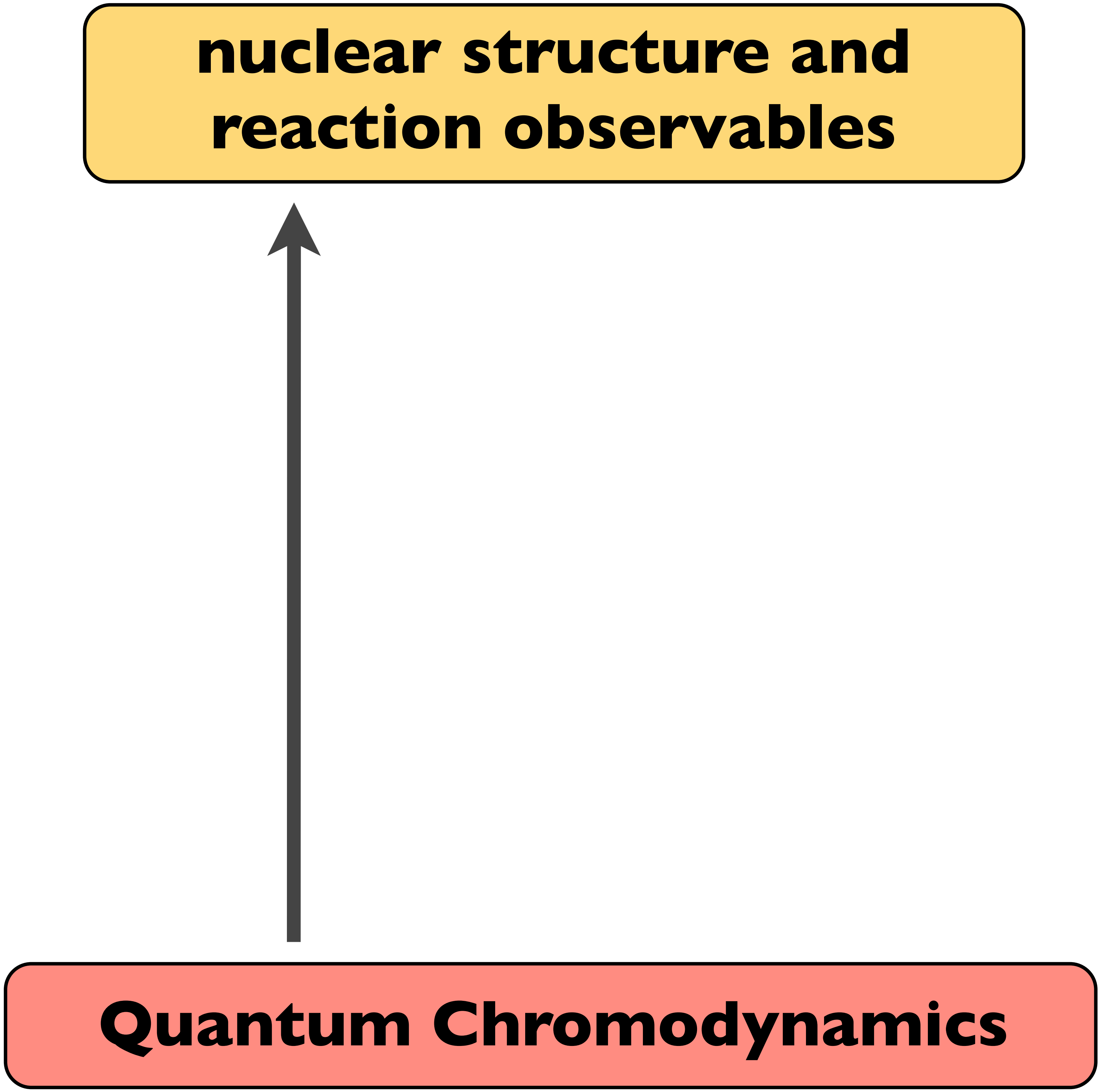
The theoretical nuclear landscape several years ago...



Ab initio nuclear structure theory



Ab initio nuclear structure theory

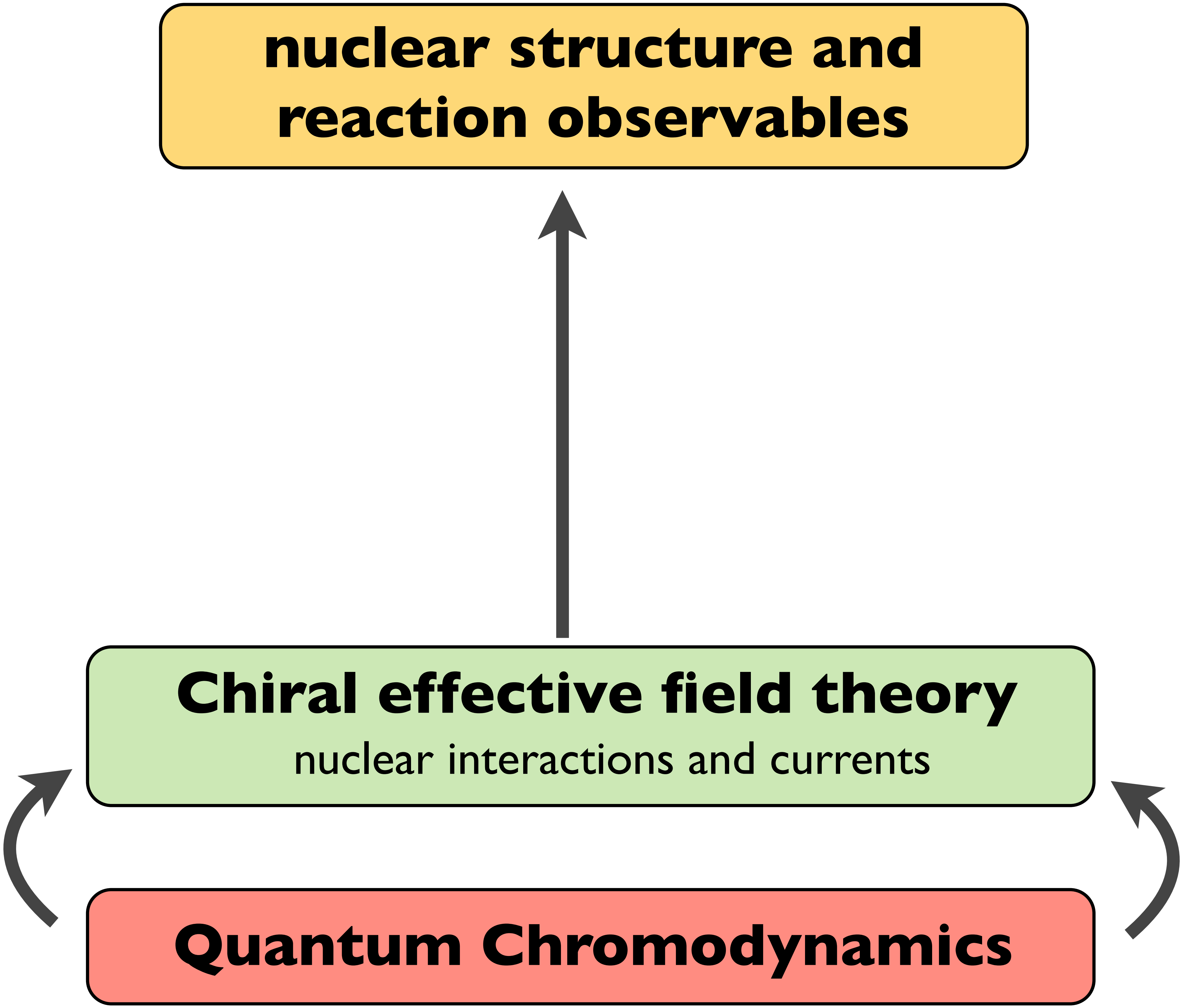
nuclear structure and reaction observables

Lattice QCD

- requires extreme amounts of computational resources
- currently limited to 1- or 2-nucleon systems
- current accuracy insufficient for precision nuclear structure

Quantum Chromodynamics

Ab initio nuclear structure theory



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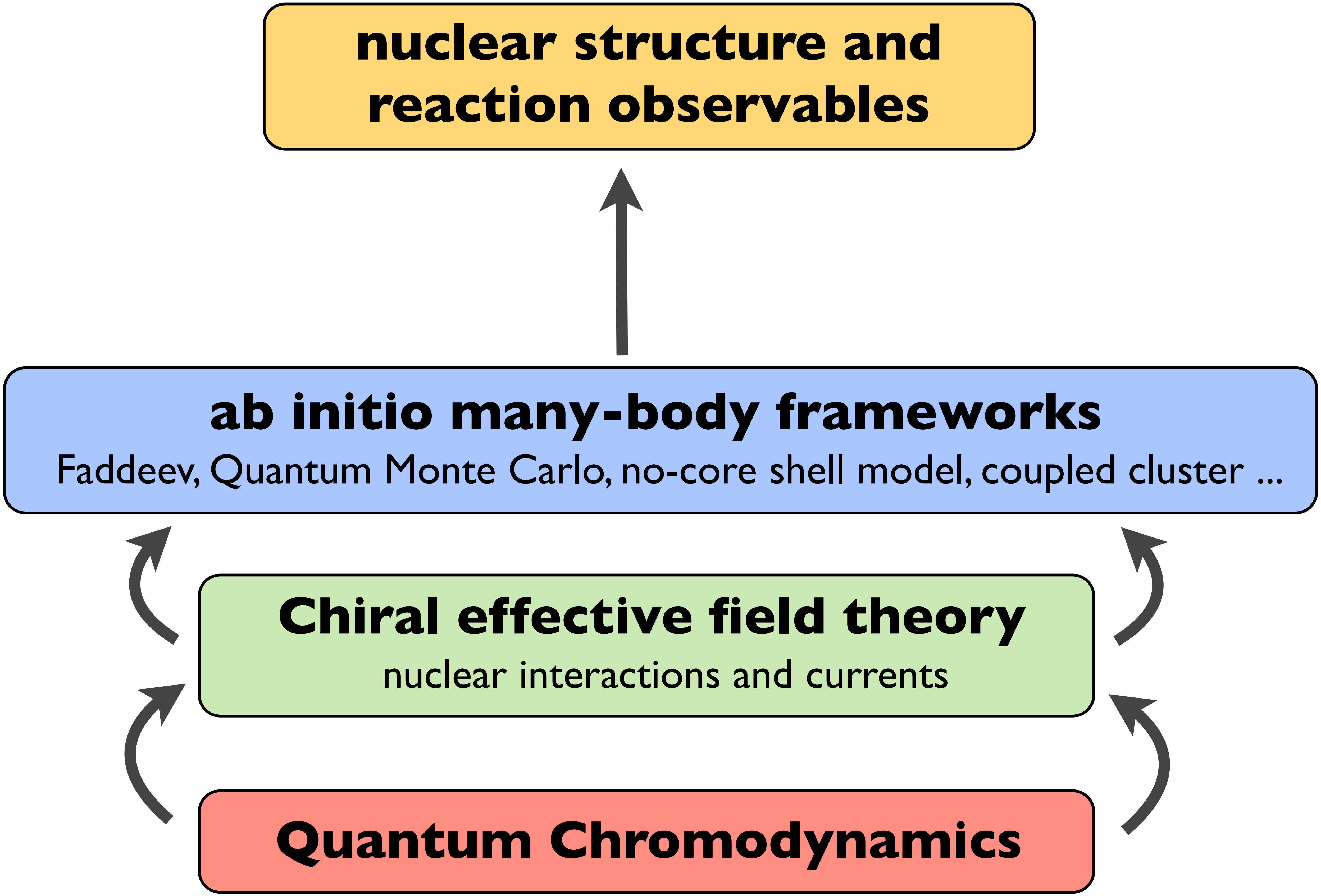

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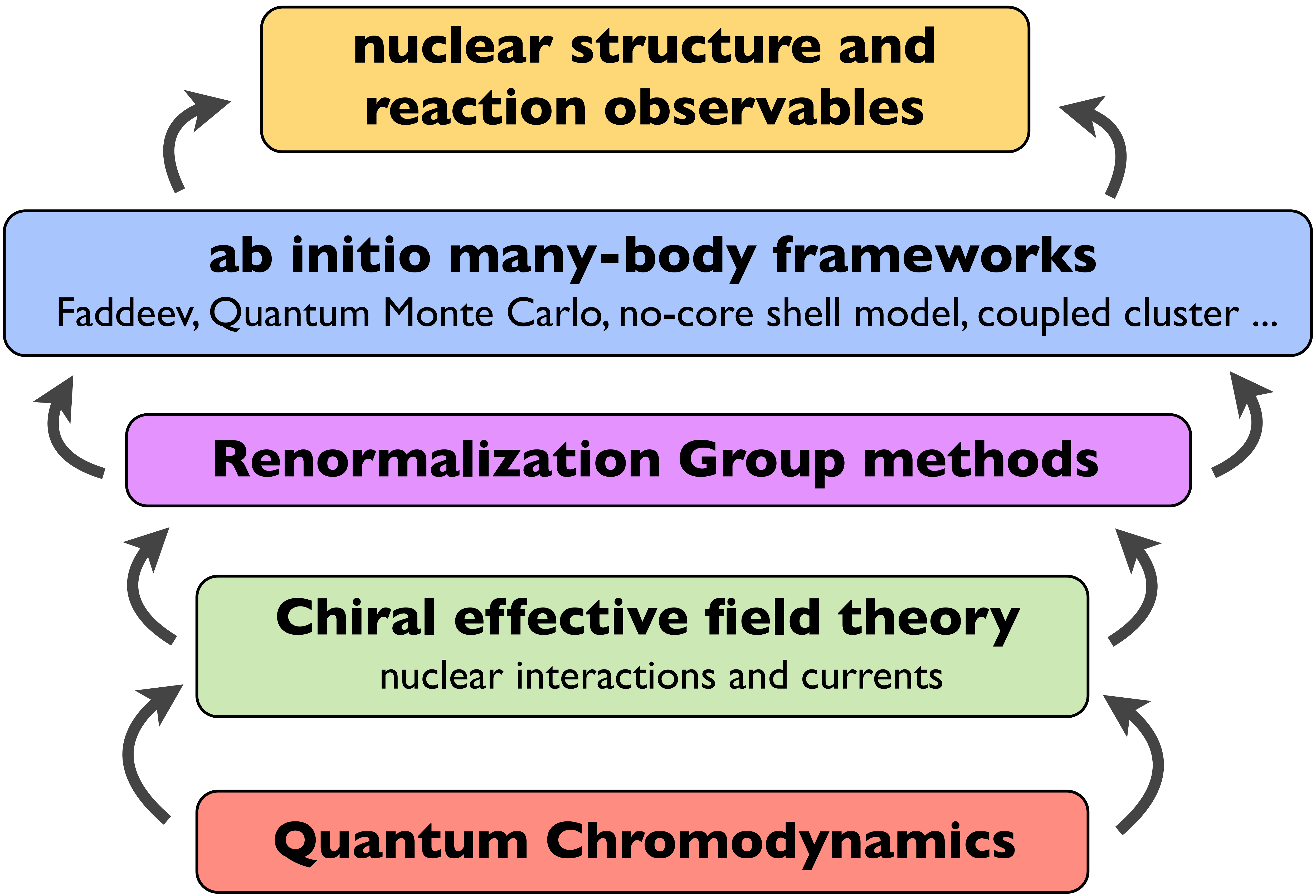
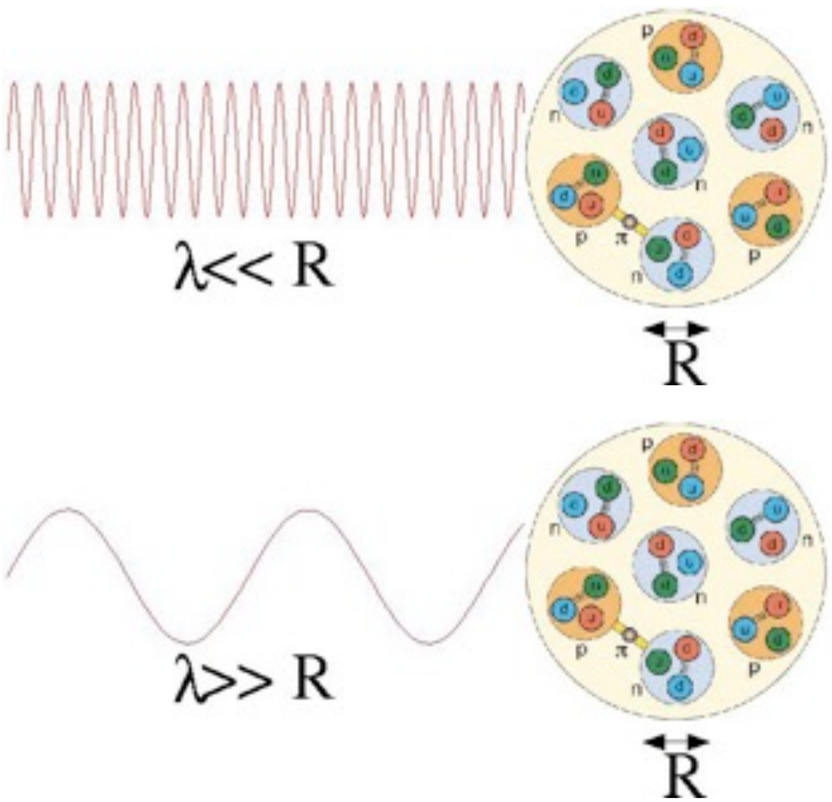


Nuclear effective degrees of freedom



- if a nucleus is probed at high energies, nucleon substructure is resolved
- at low energies, details are not resolved

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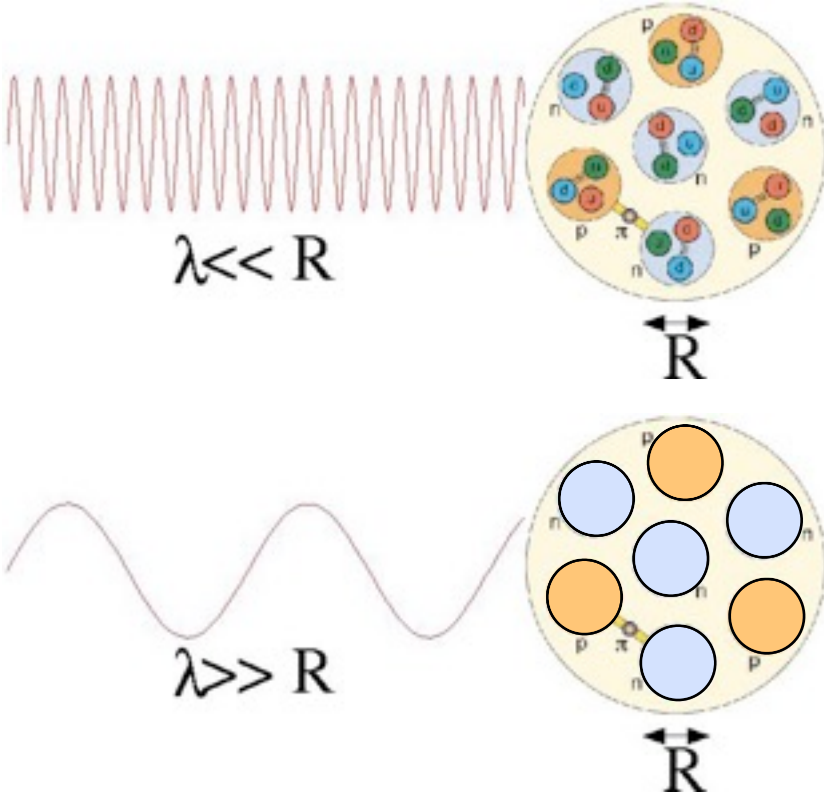
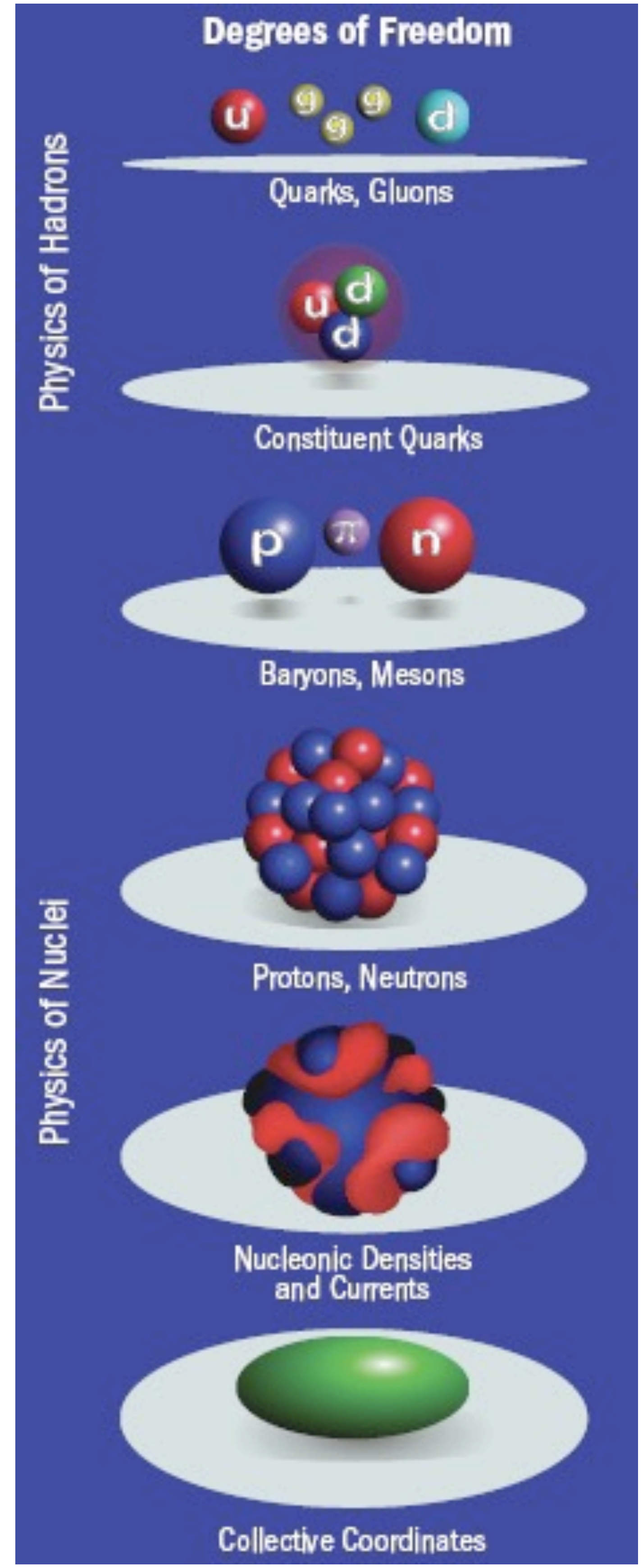


- if a nucleus is probed at high energies, nucleon substructure is resolved
- at low energies, details are not resolved
- replace fine structure by something simpler (like multipole expansion), low-energy observables unchanged

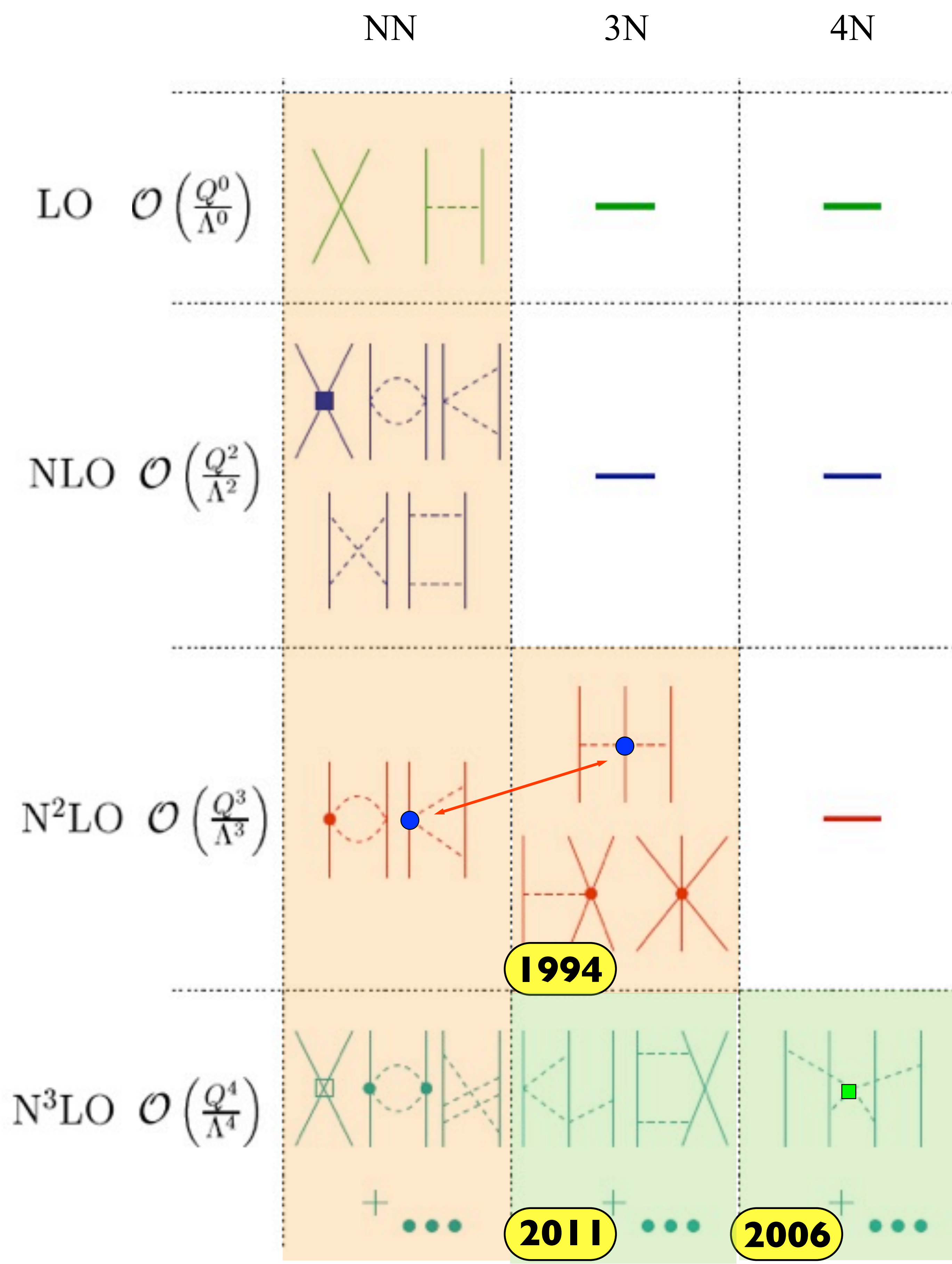
⟶ effective field theory



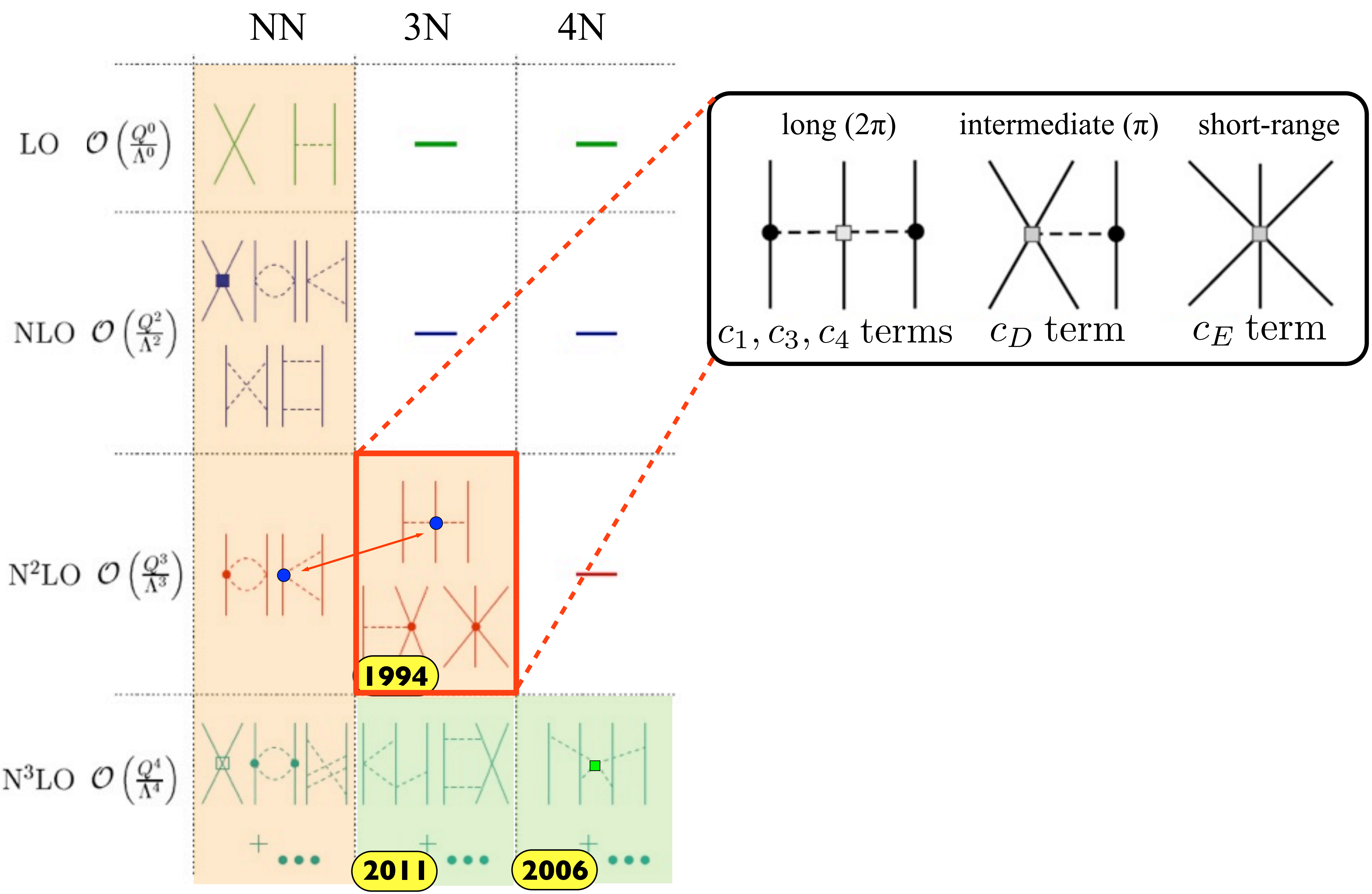
Resolution ↓

Chiral effective field theory for nuclear forces

- choose relevant degrees of freedom: here nucleons and pions
- operators constrained by symmetries of QCD
- short-range physics captured in few short-range couplings
- separation of scales: Q << Λ_b, breakdown scale Λ_b~500 MeV
- power-counting: expand in powers Q/Λ_b
- systematic: work to desired accuracy, obtain error estimates



Many-body forces in chiral EFT



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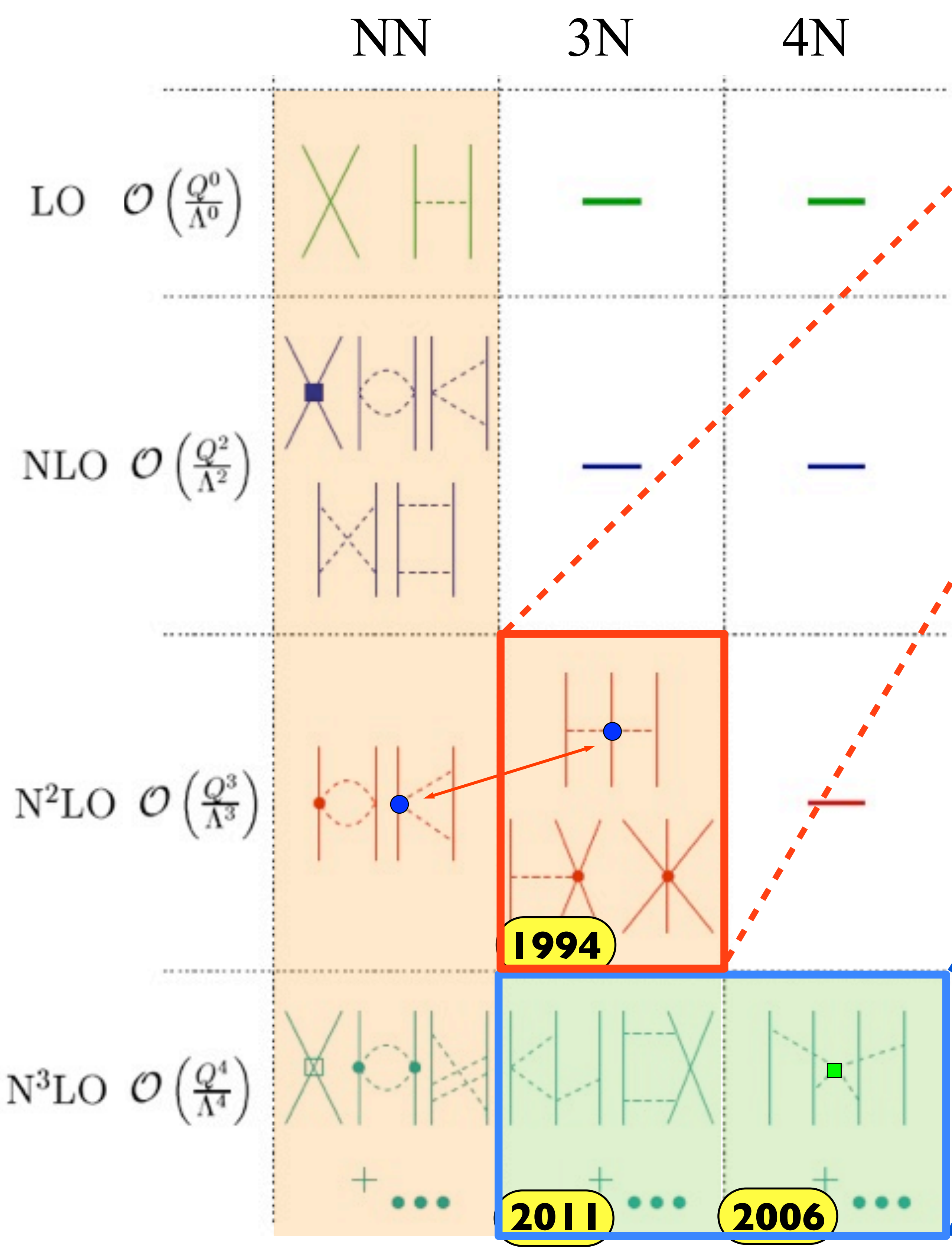
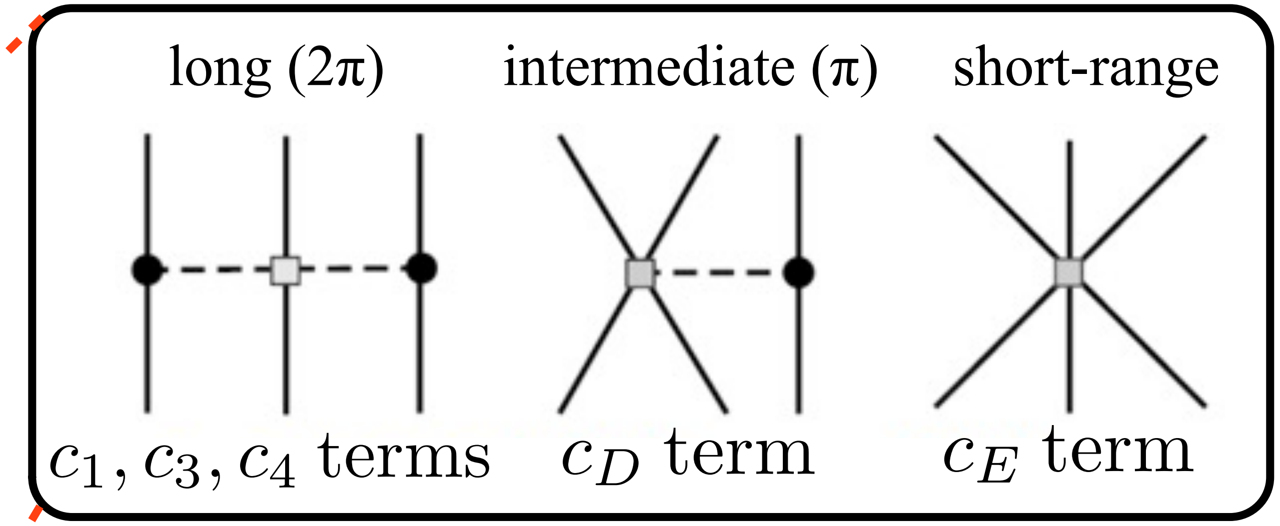


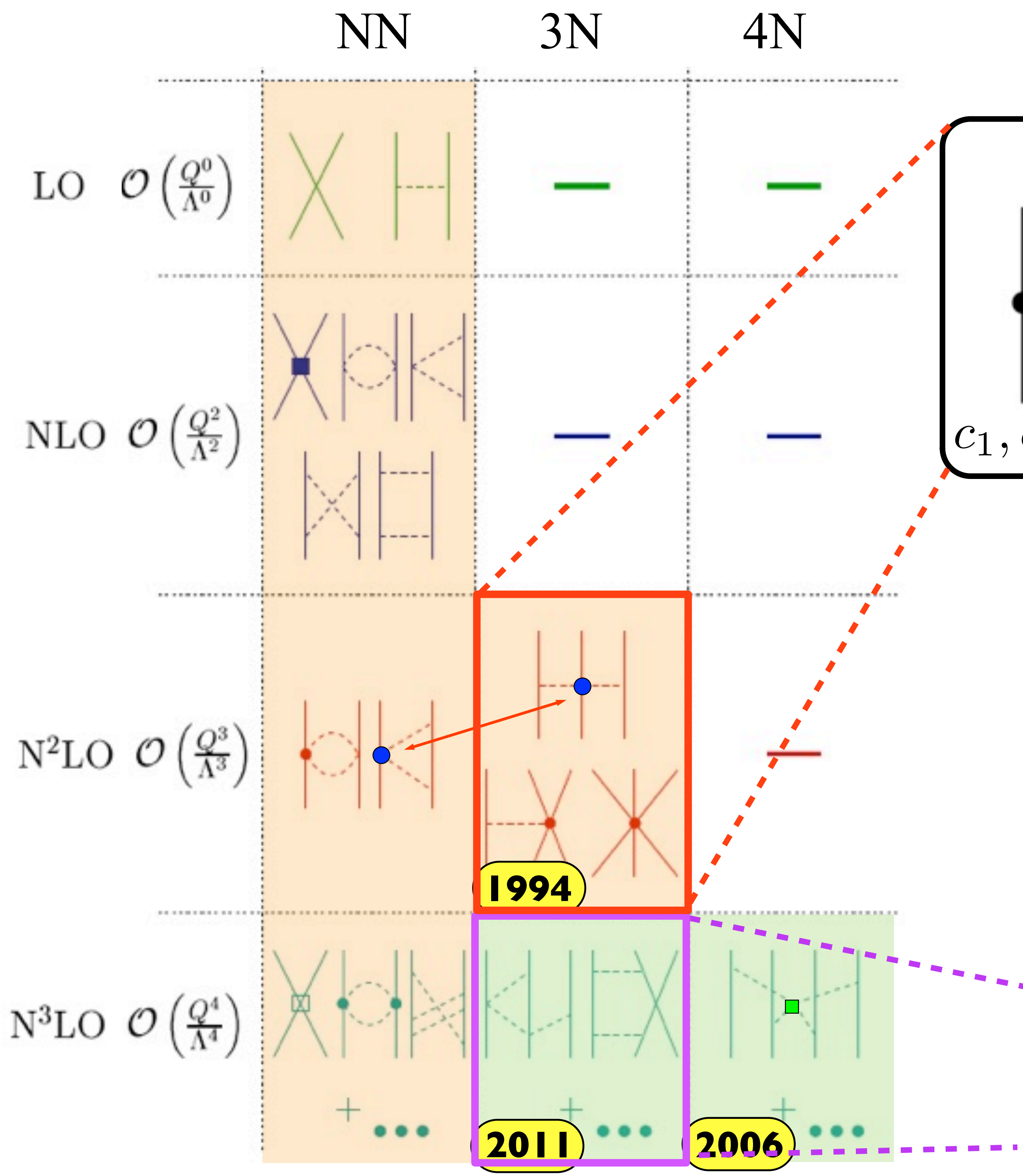


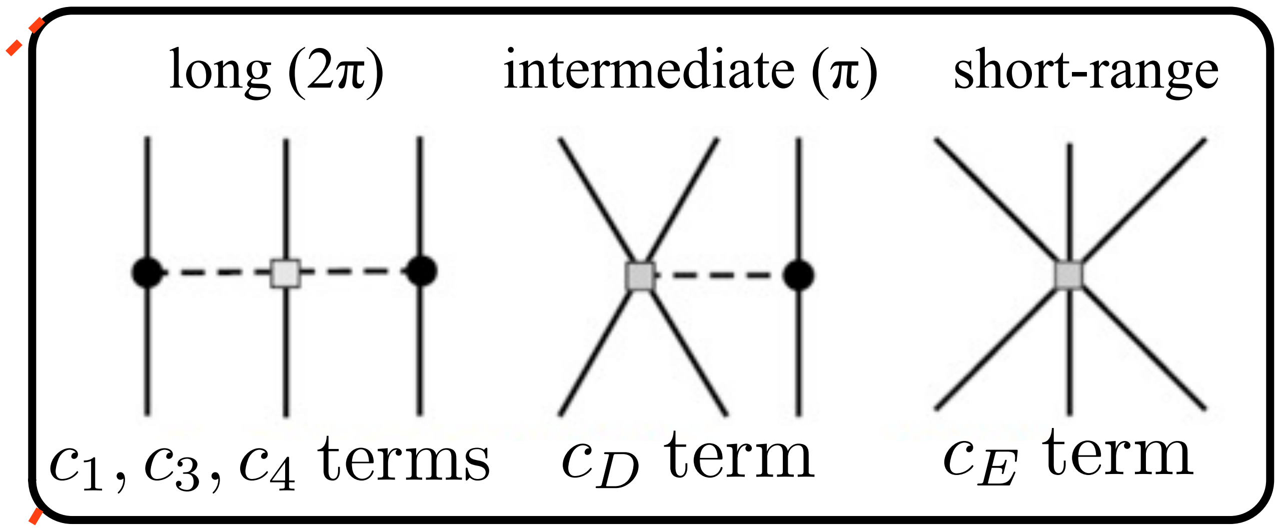
first incorporation in calculations of neutron and nuclear matter

Tews, Krueger, KH, Schwenk, PRL 110, 032504 (2013)
Krueger, Tews, KH, Schwenk, PRC 88, 025802 (2013)

all terms predicted

Many-body forces in chiral EFT





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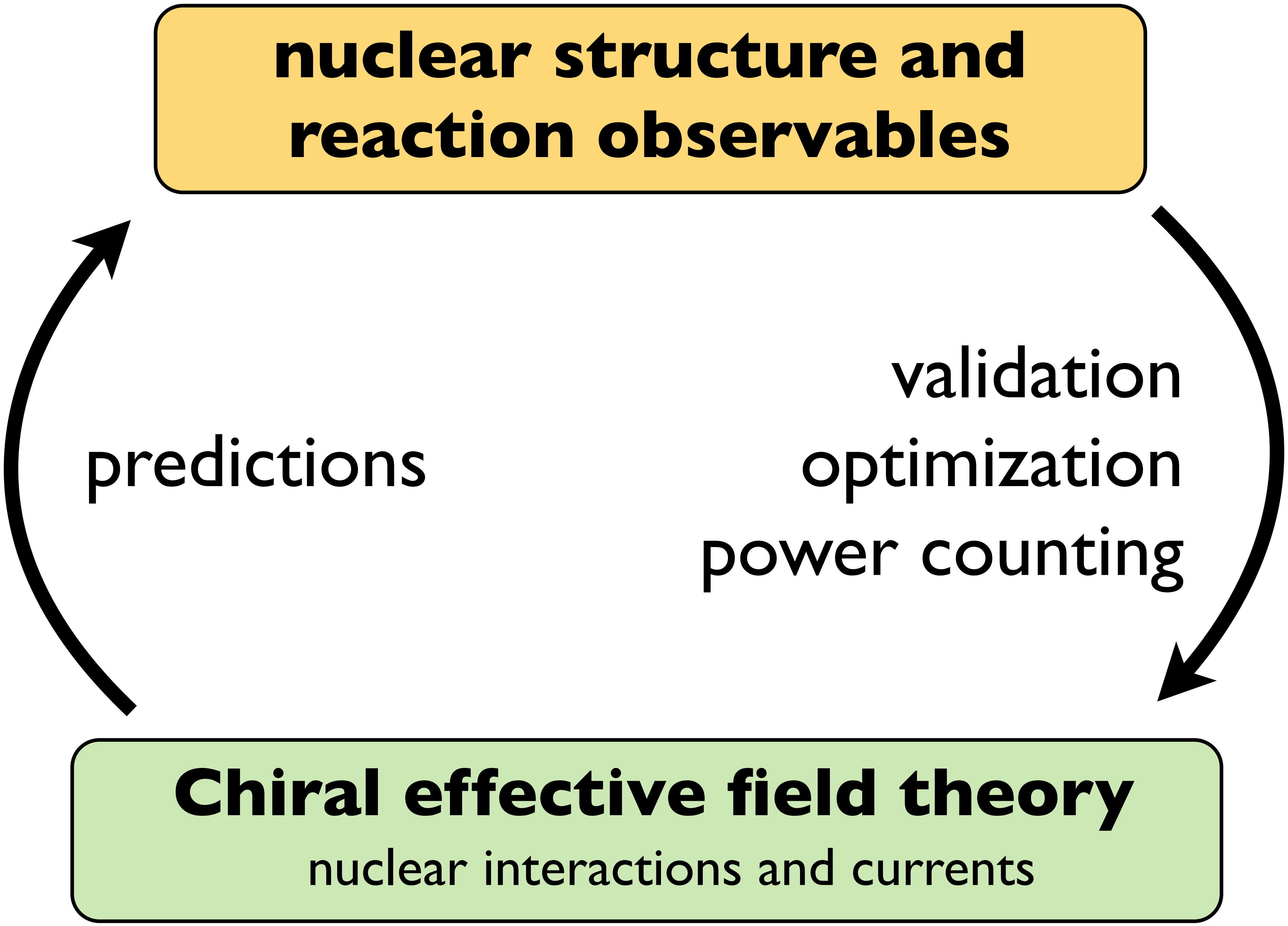
first calculation of matrix elements for ab initio studies of matter and nuclei

KH, Krebs, Epelbaum, Golak, Skibinski, PRC 91, 044001(2015)

Development of nuclear interactions



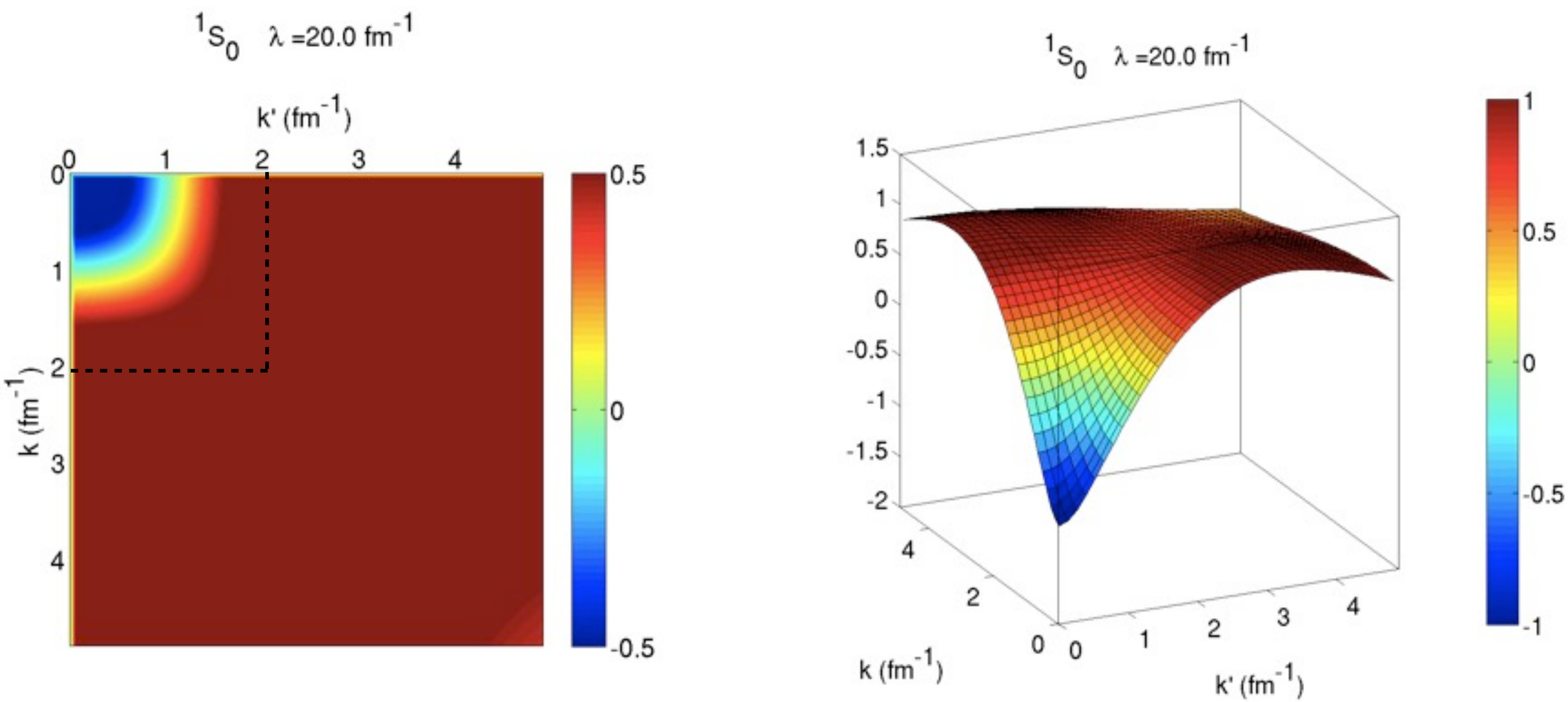
LENPIC

The Similarity Renormalization Group

- generate unitary transformation which decouples low- and high momenta

$$H_\lambda = U_\lambda H U_\lambda^\dagger$$ with the resolution parameter λ

- basic idea: change resolution successively in small steps: $\frac{dH_\lambda}{d\lambda} = [\eta_\lambda, H_\lambda]$
- generator η_λ can be chosen and tailored to different applications
- observables are preserved due to unitarity of transformation



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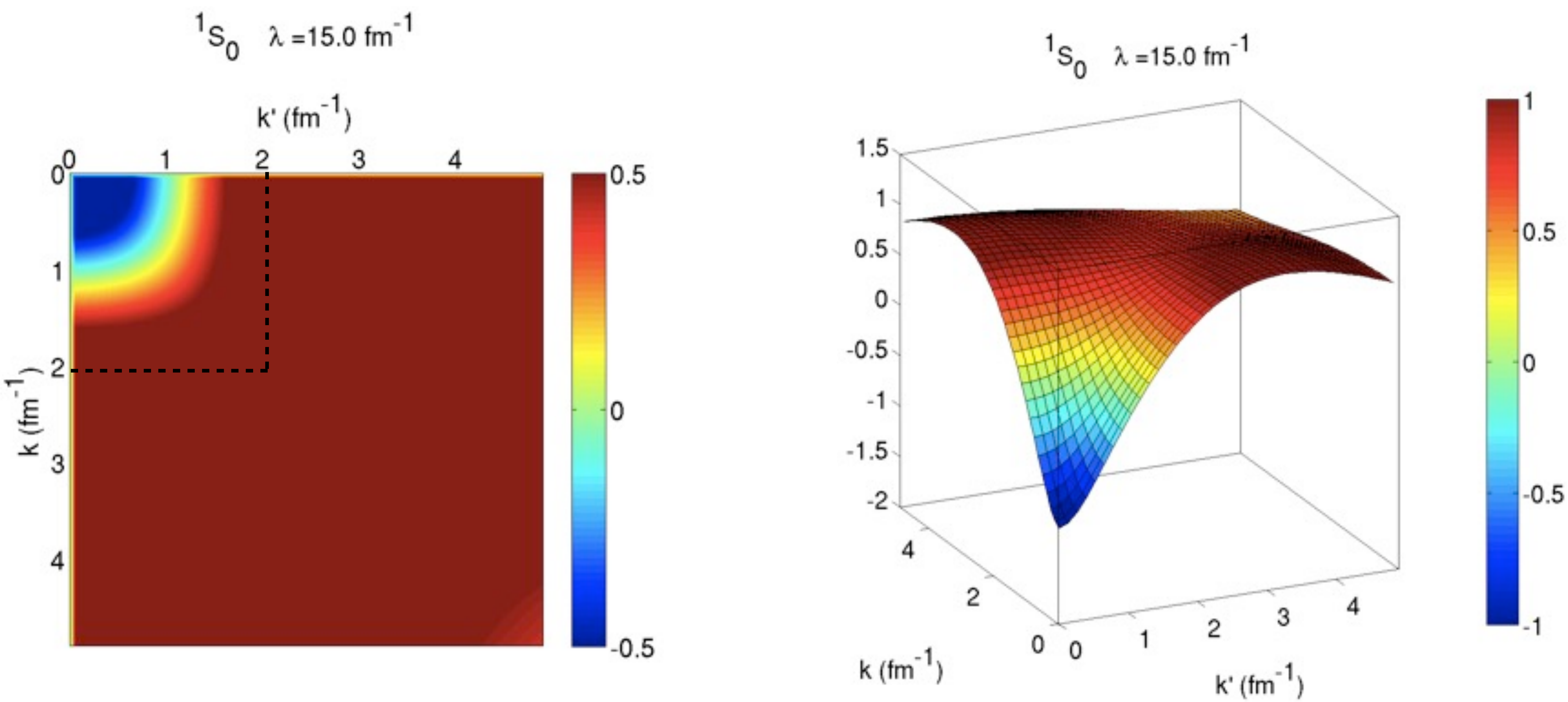

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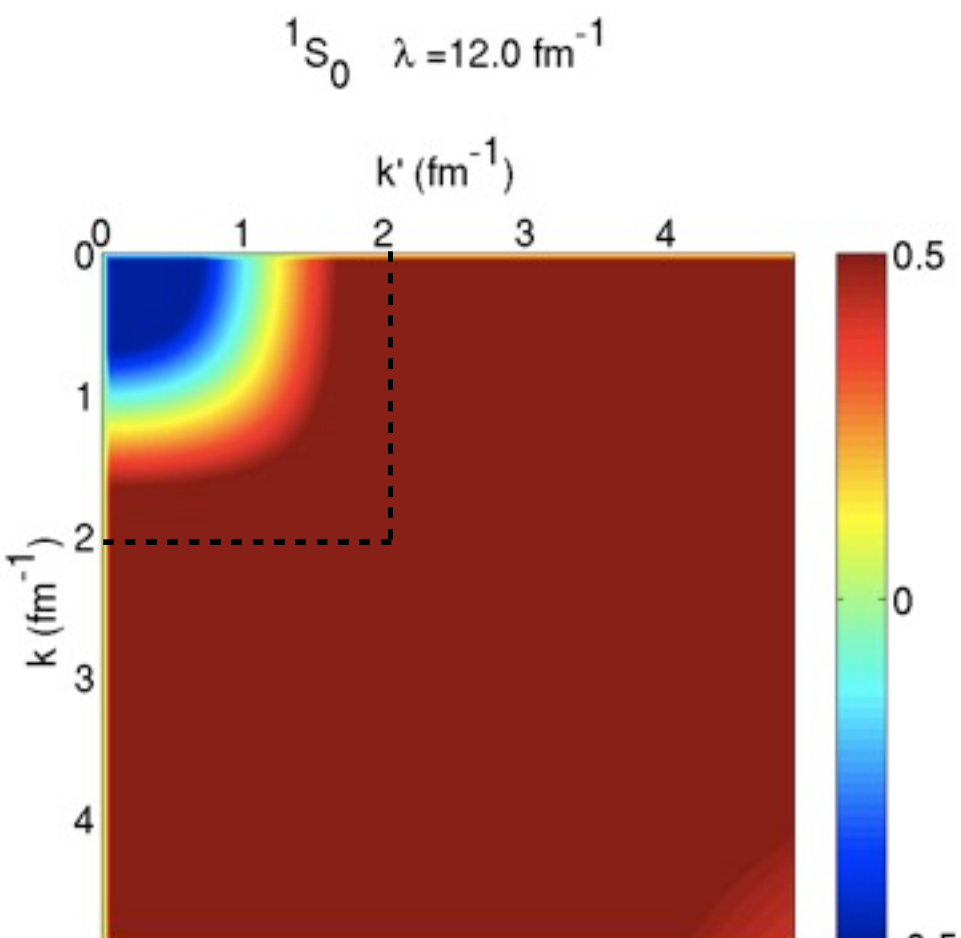

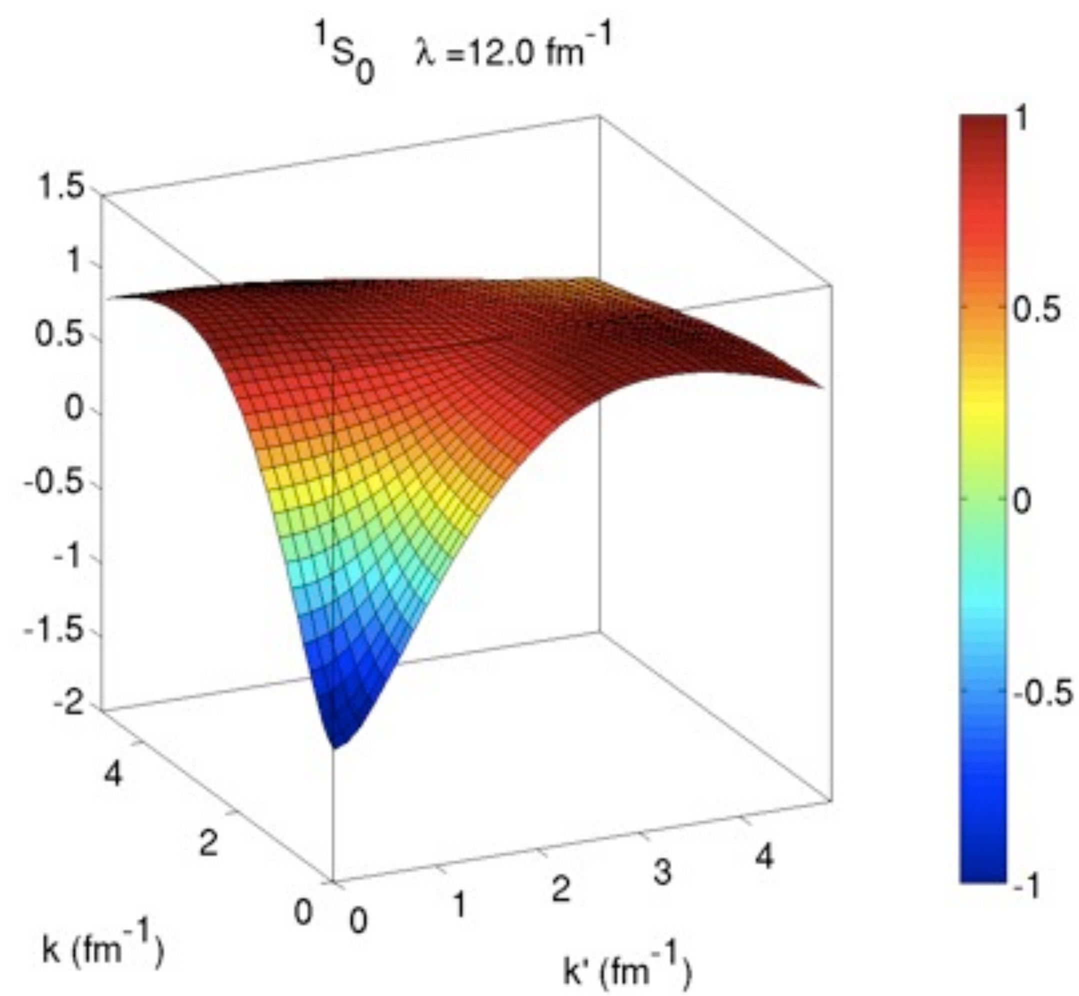

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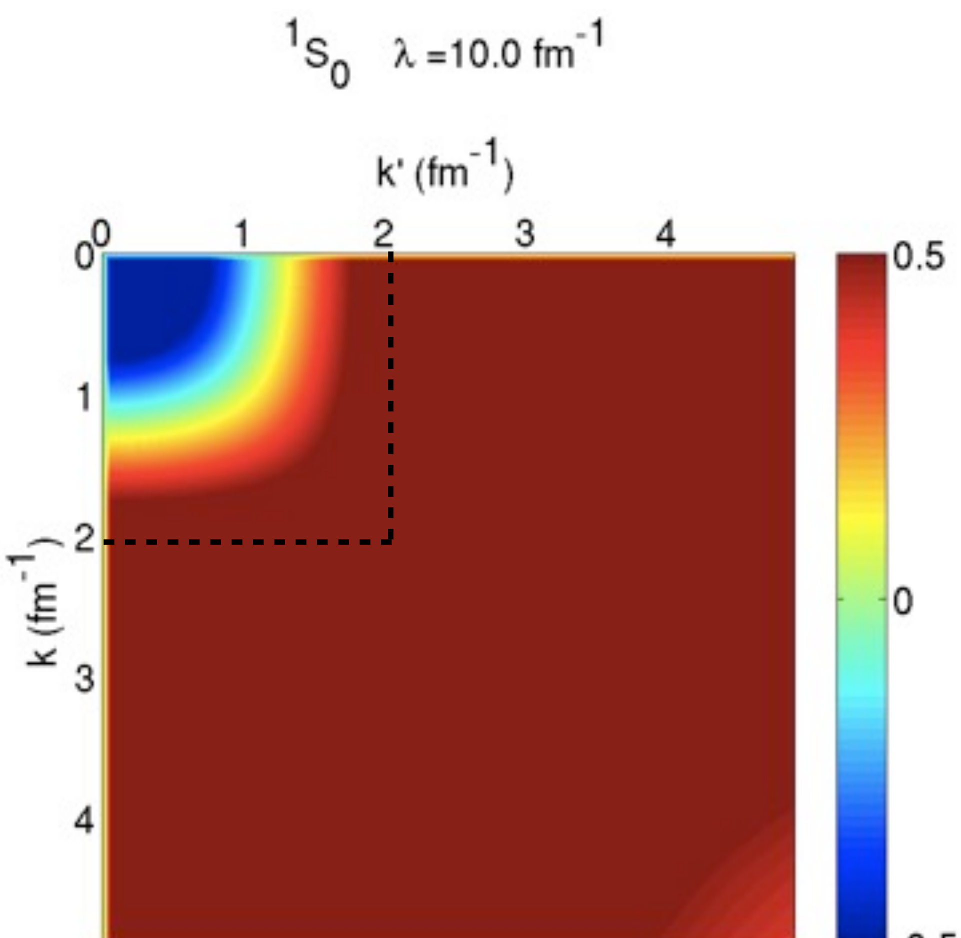

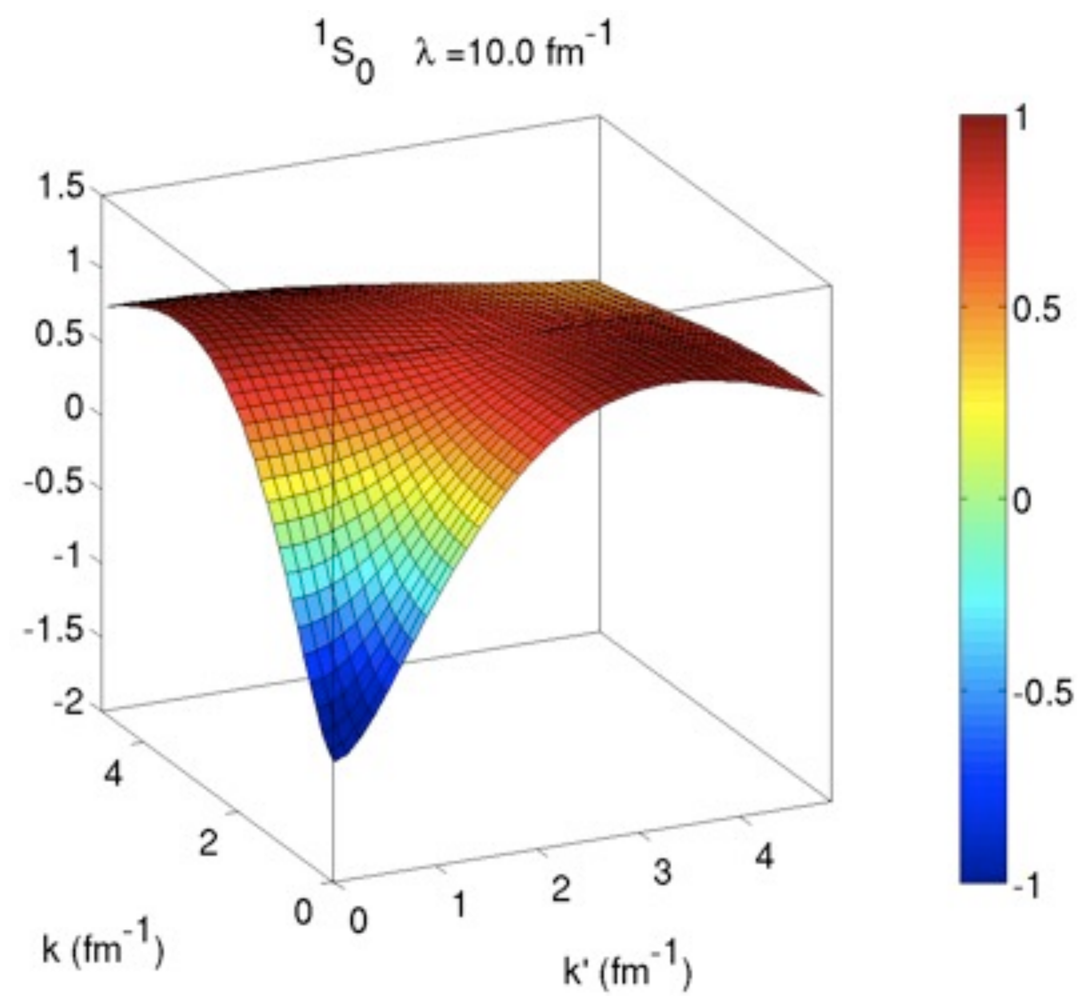

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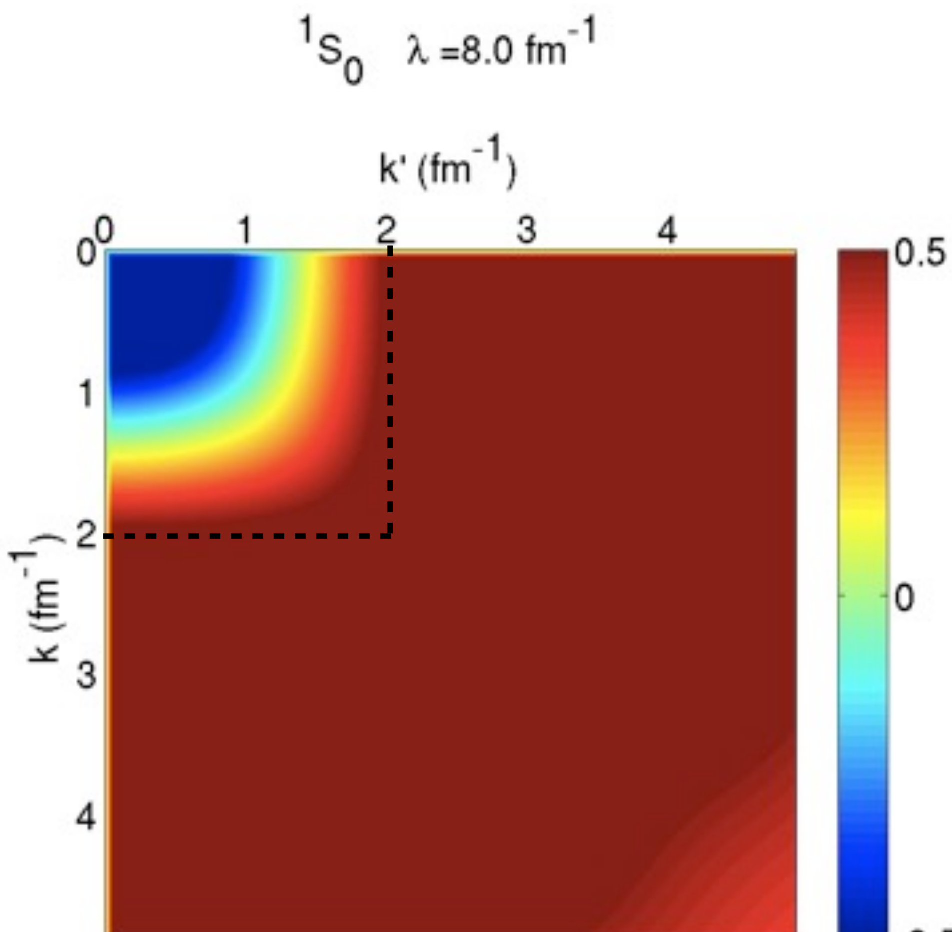

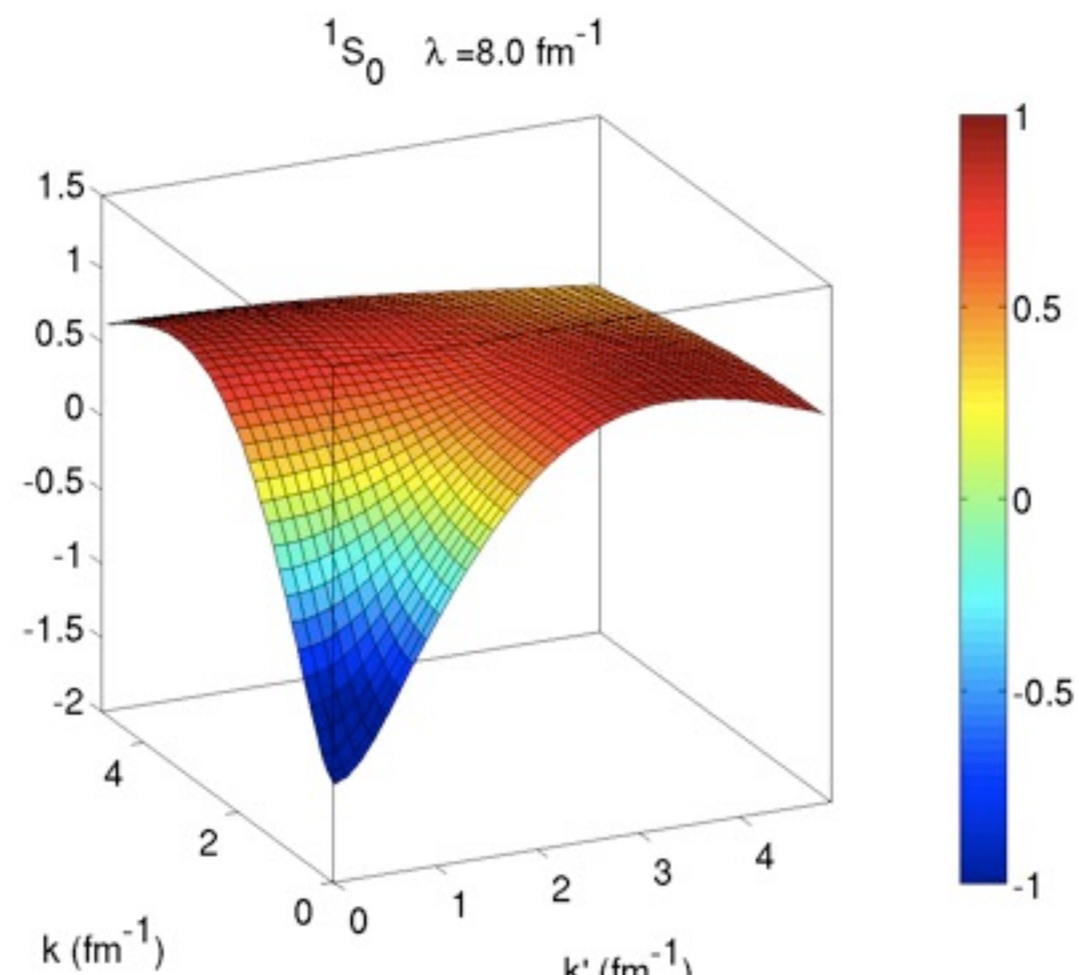

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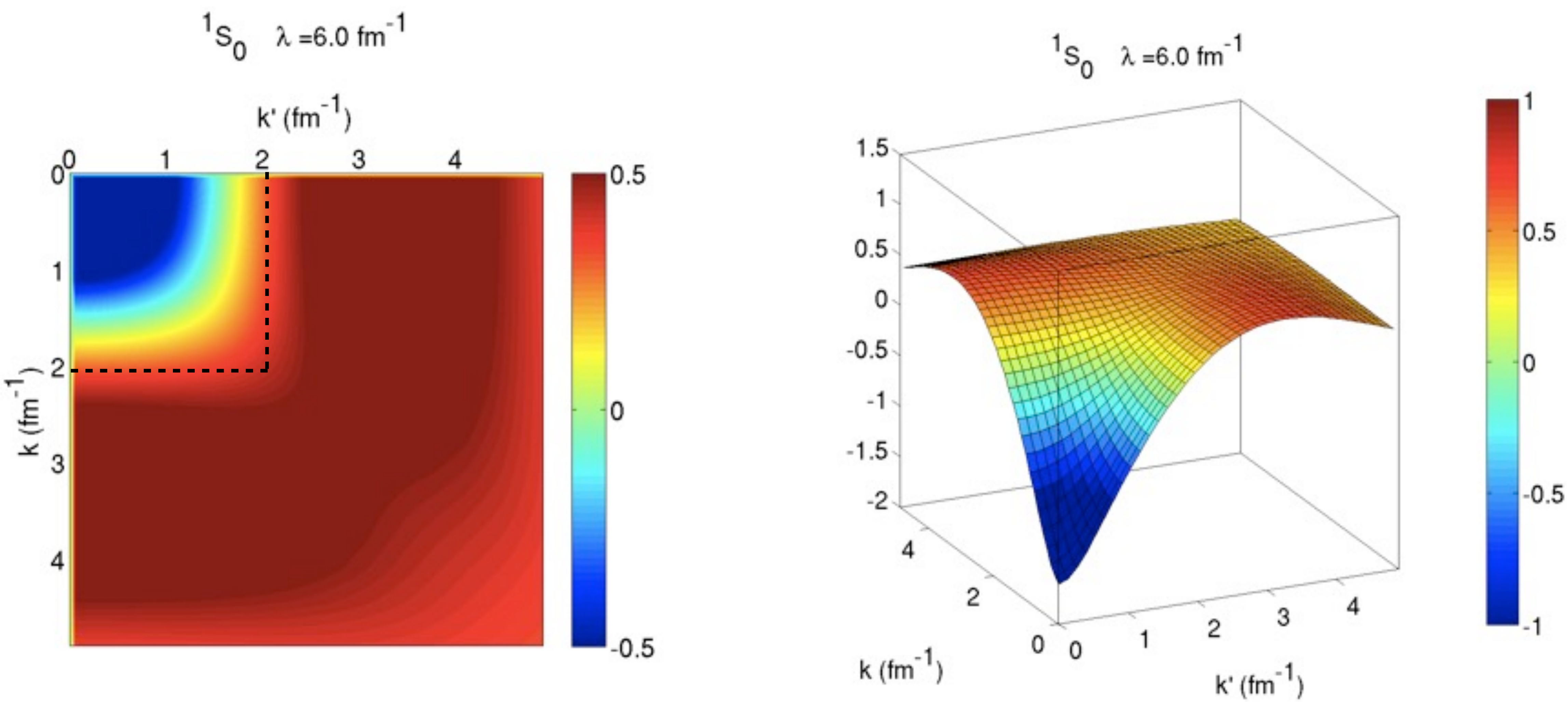

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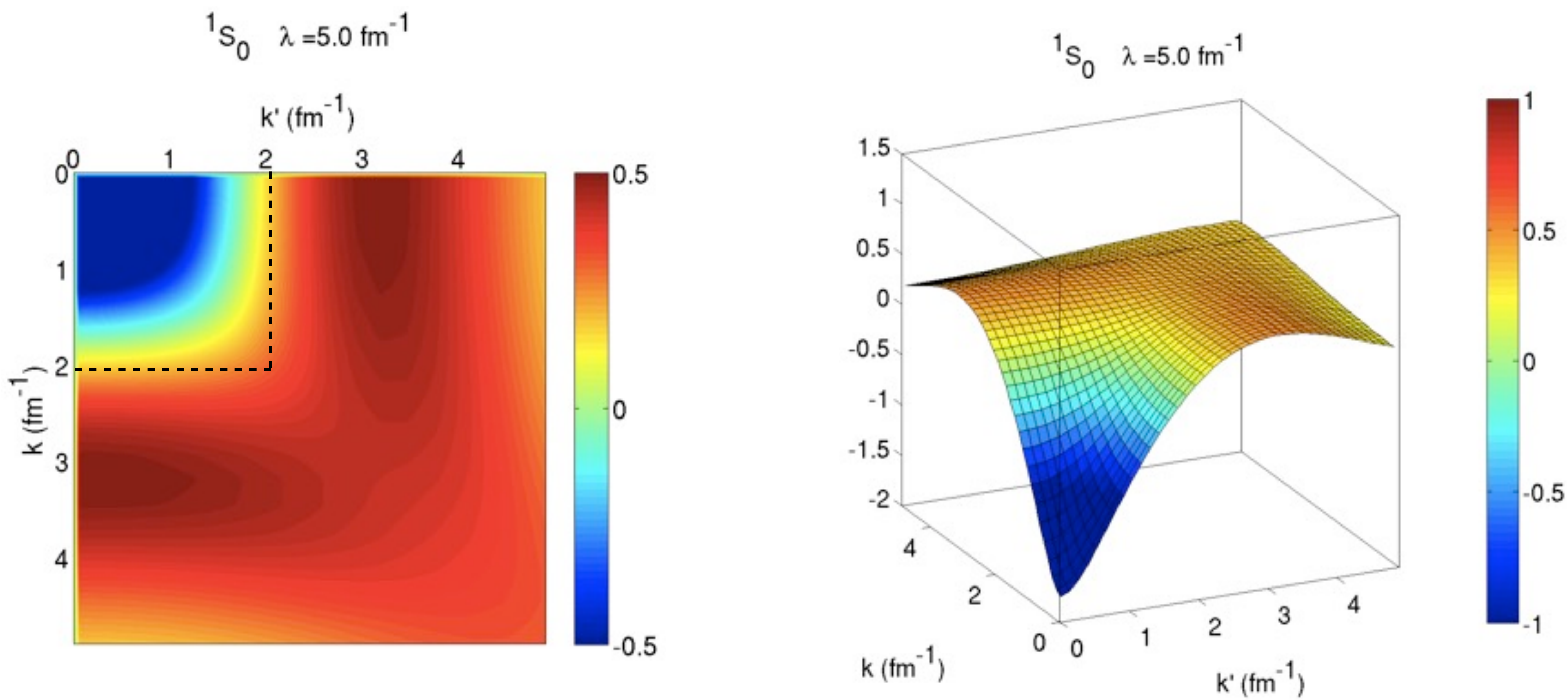

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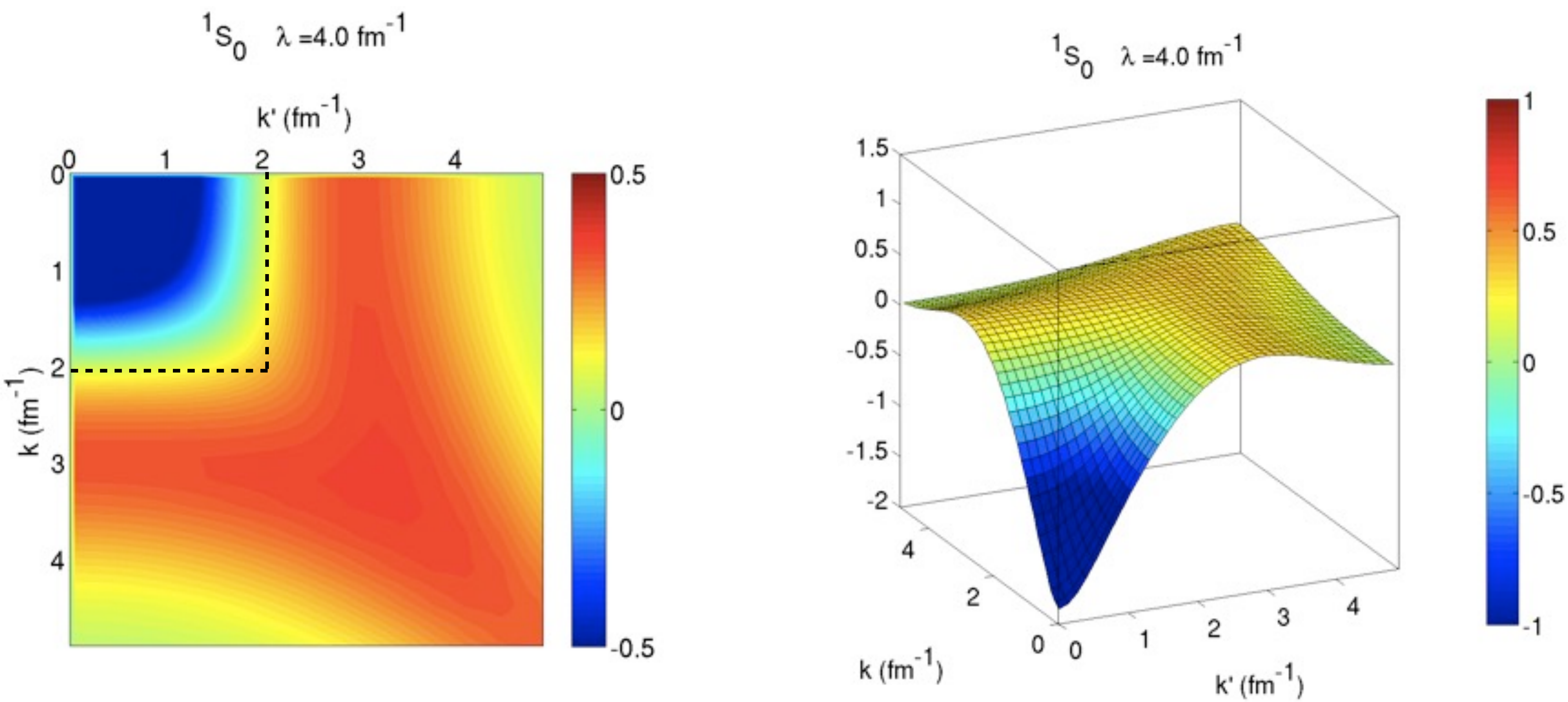

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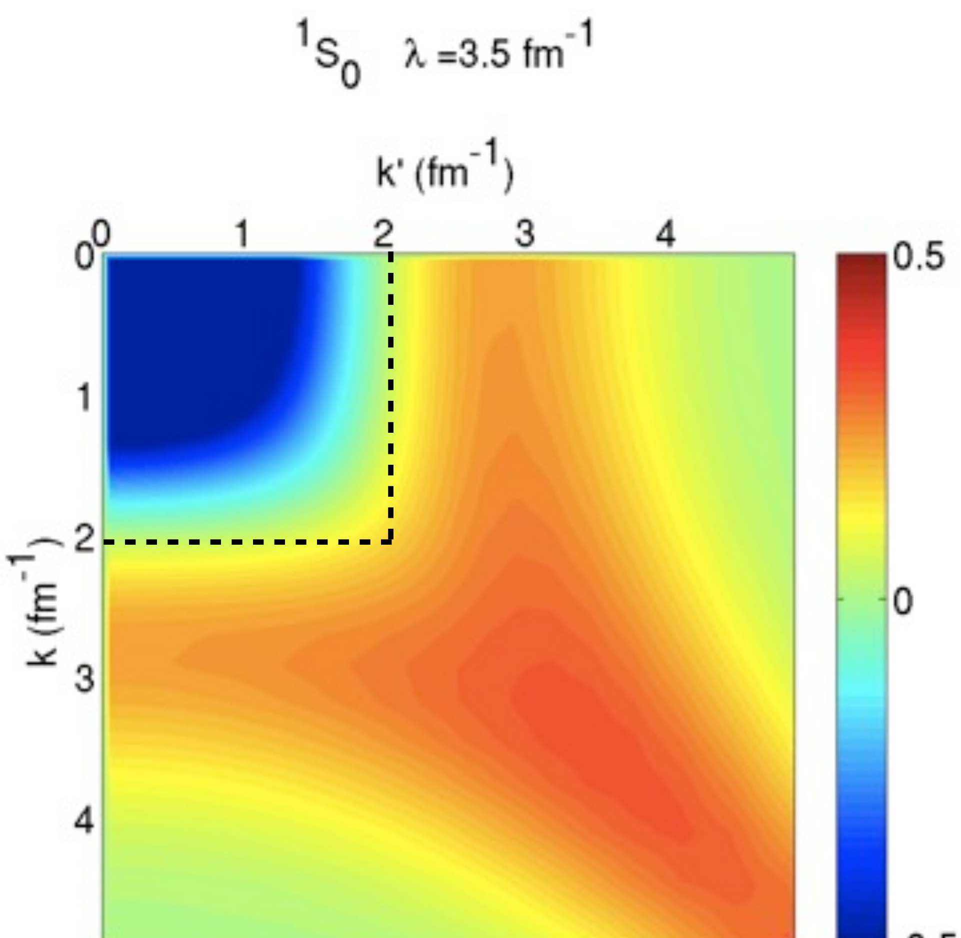

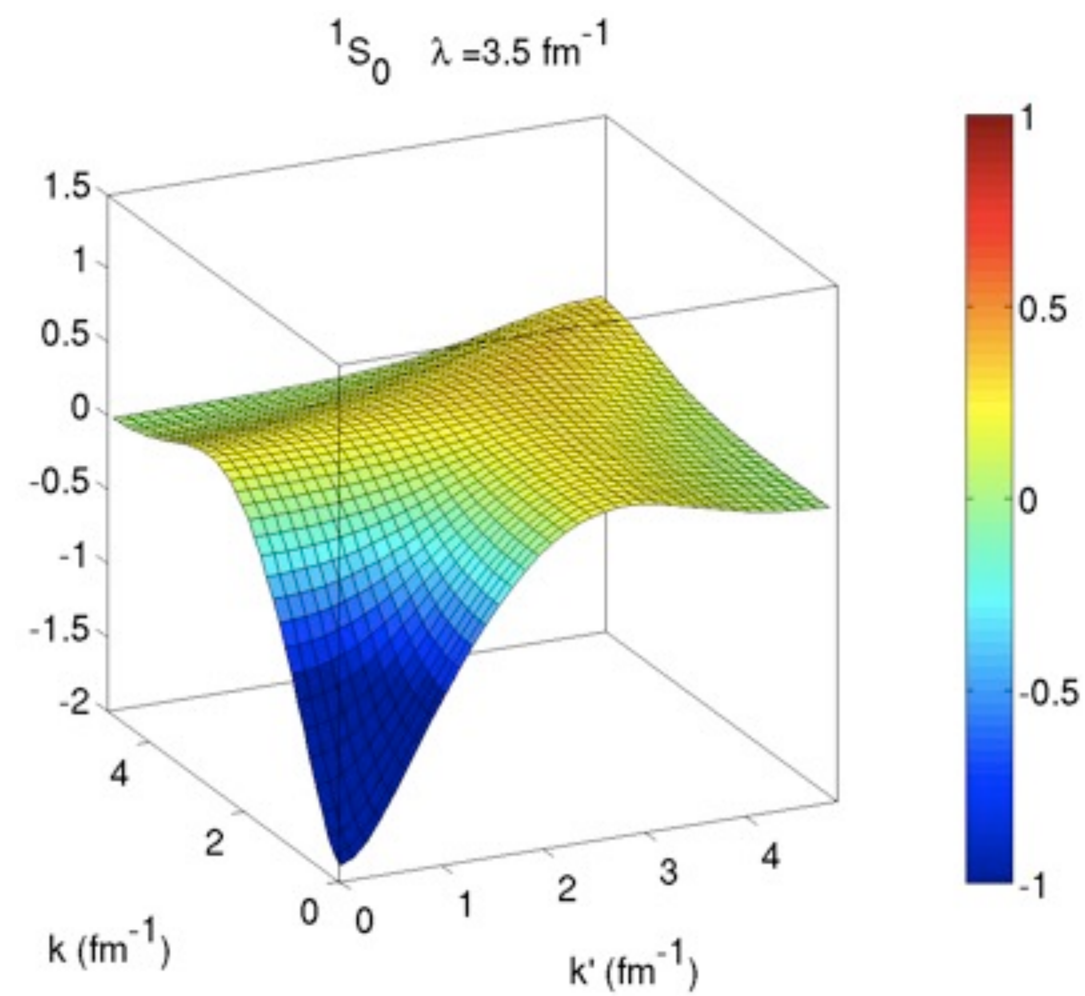

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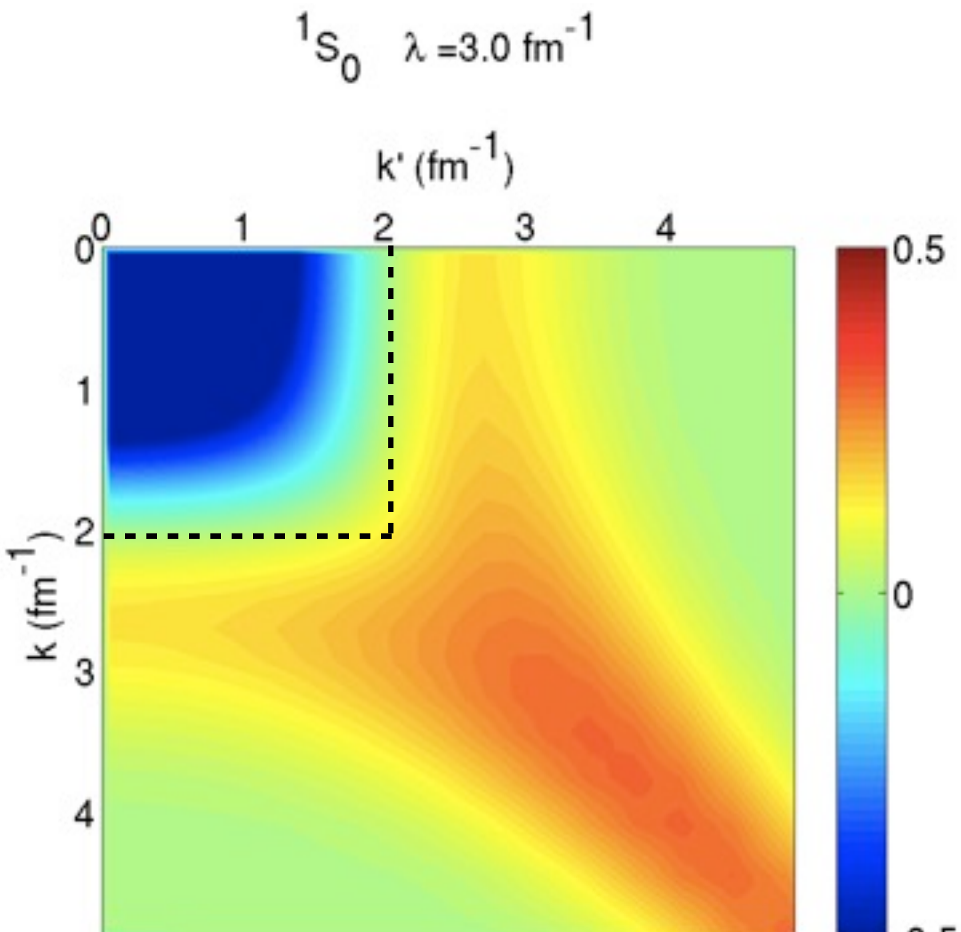

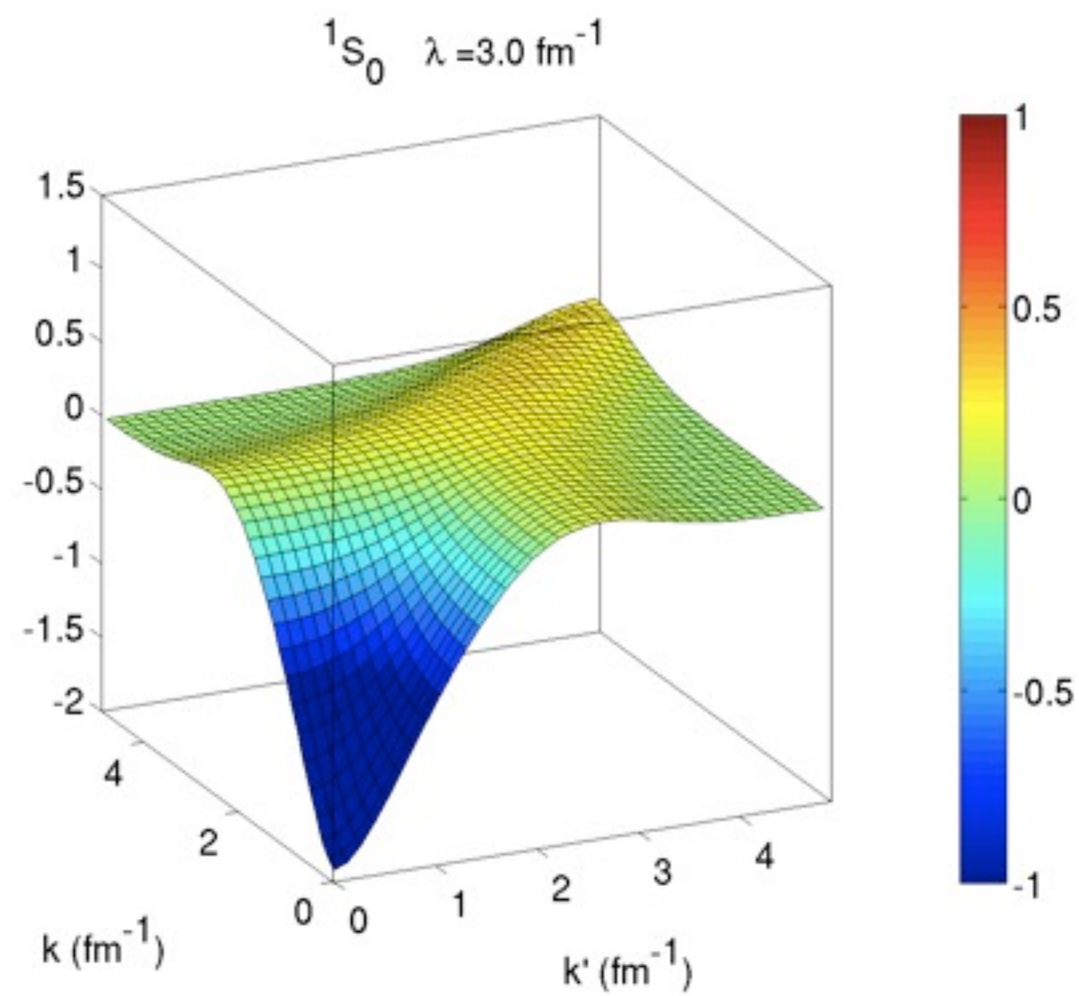

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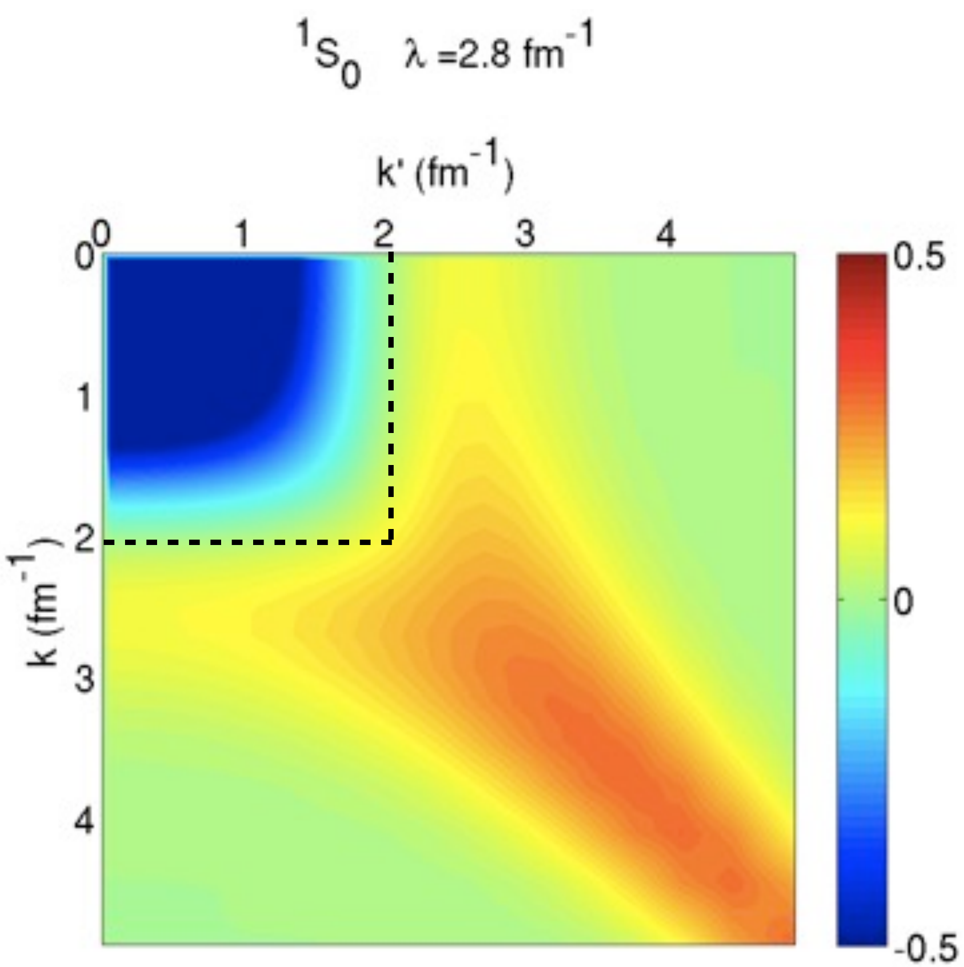

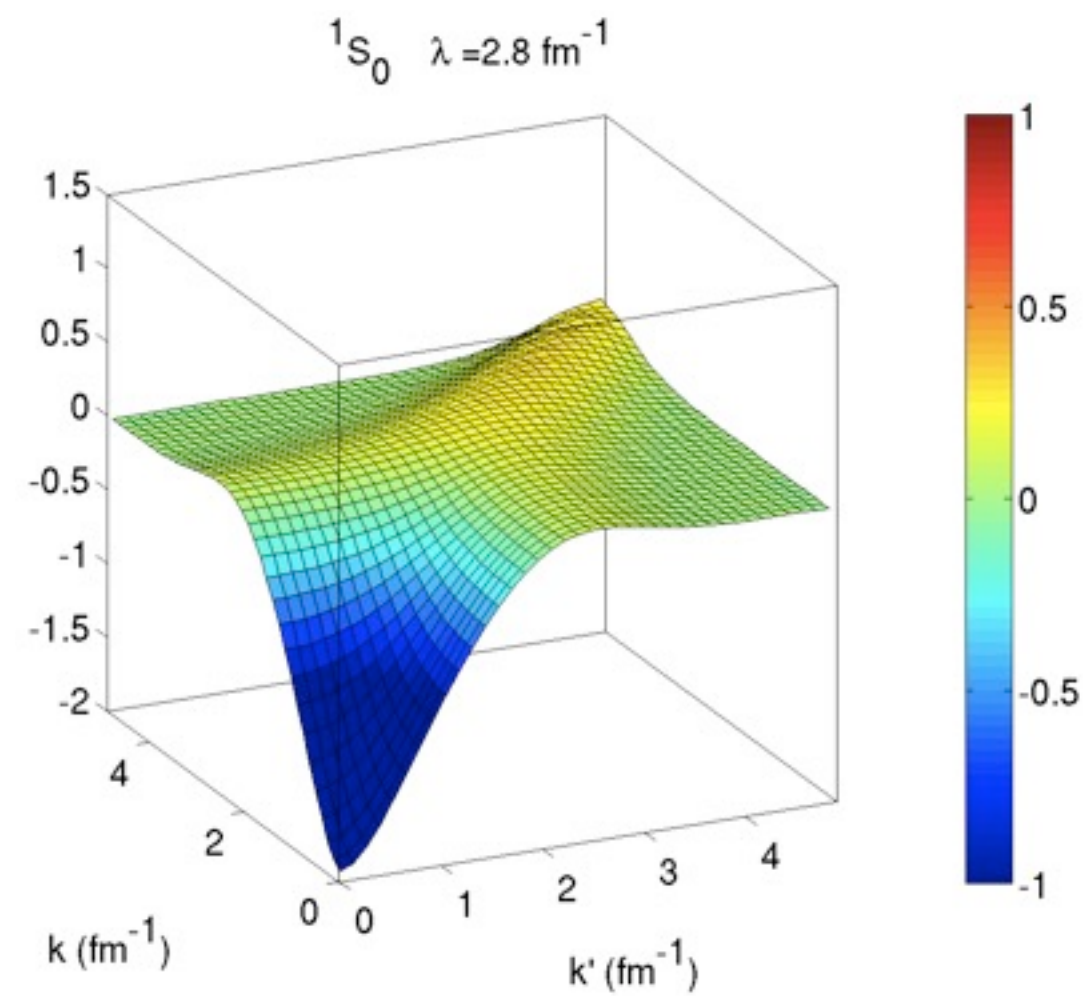

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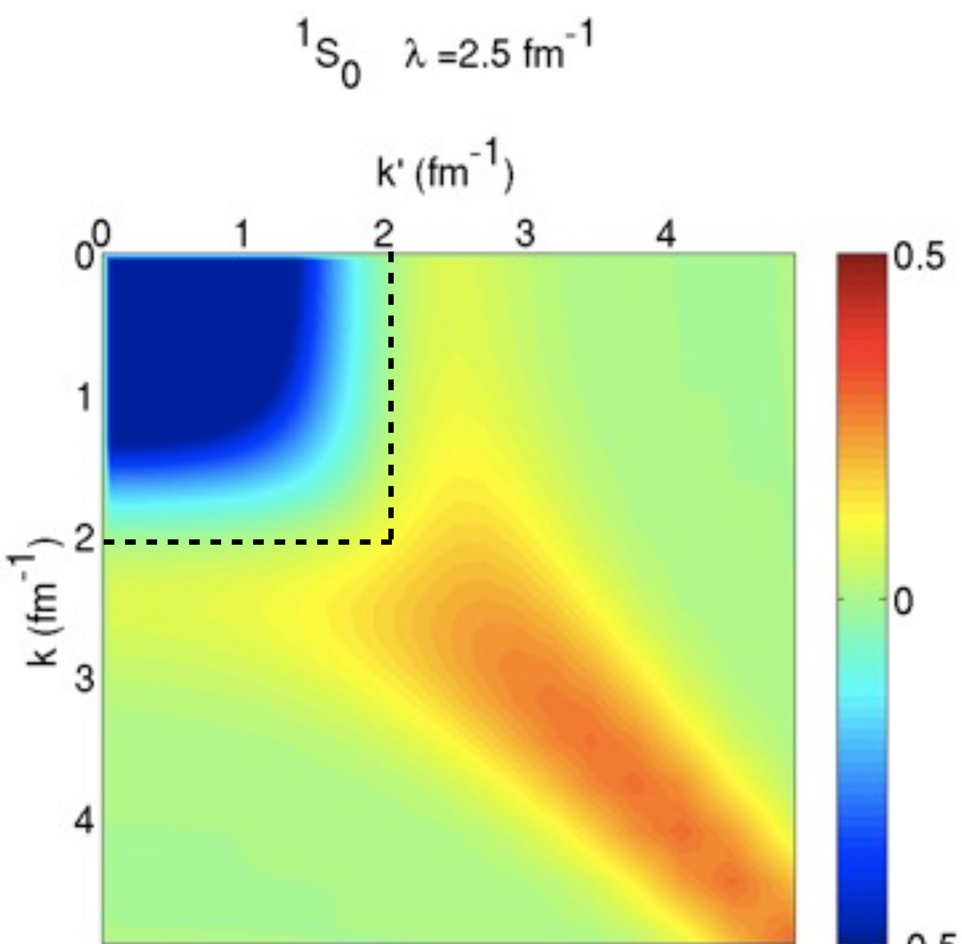

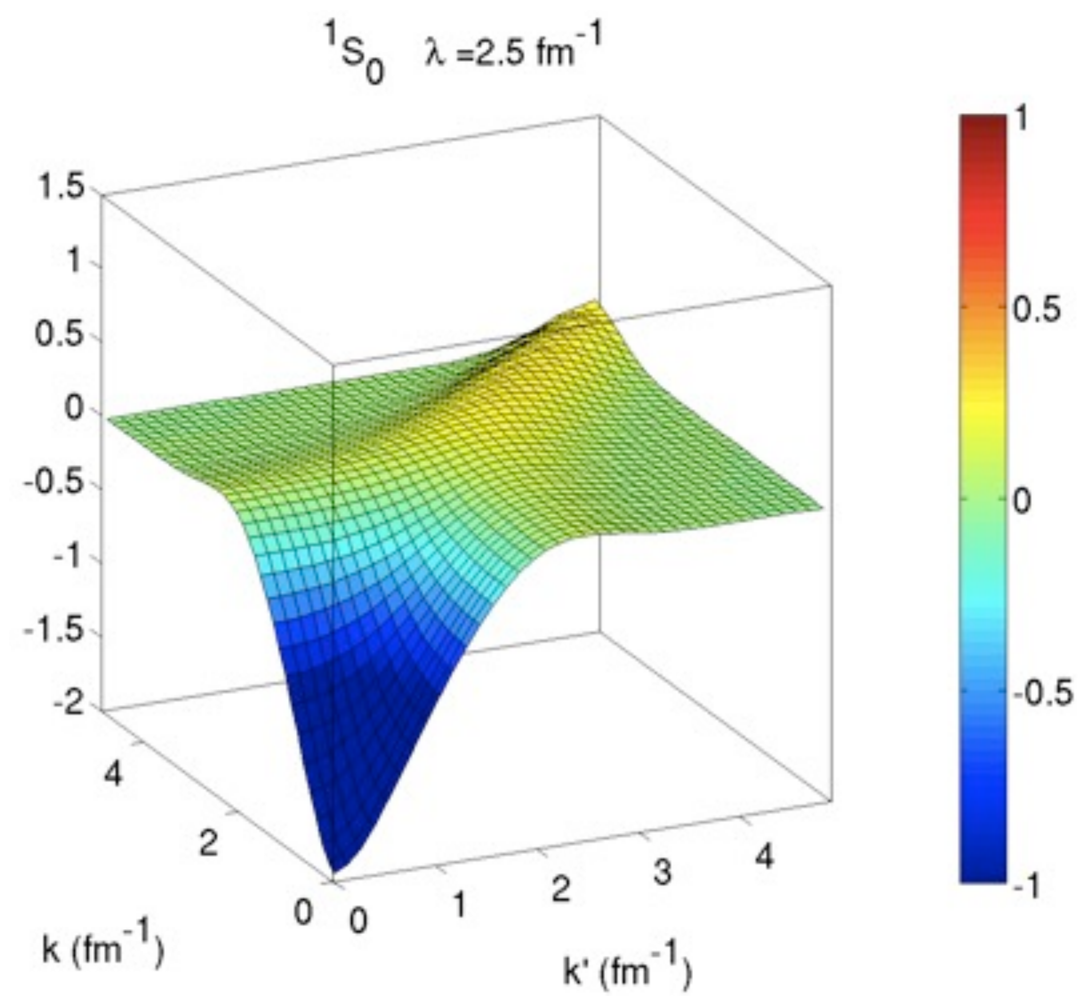

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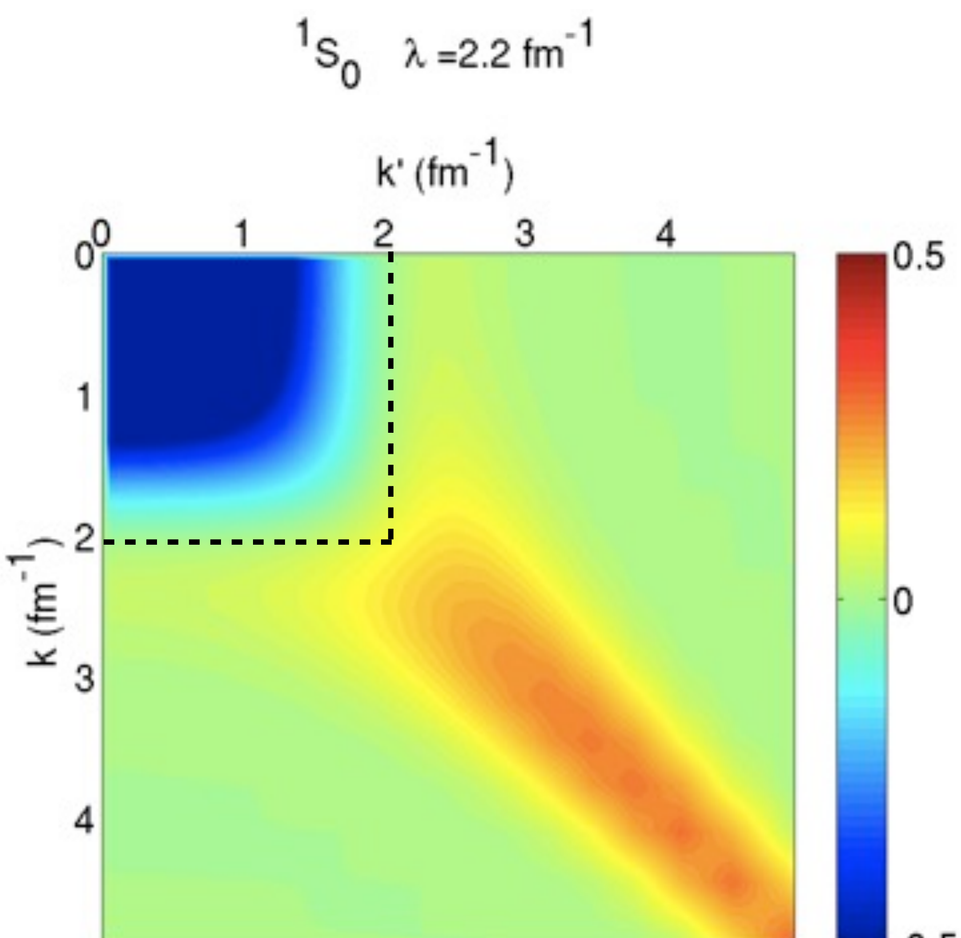

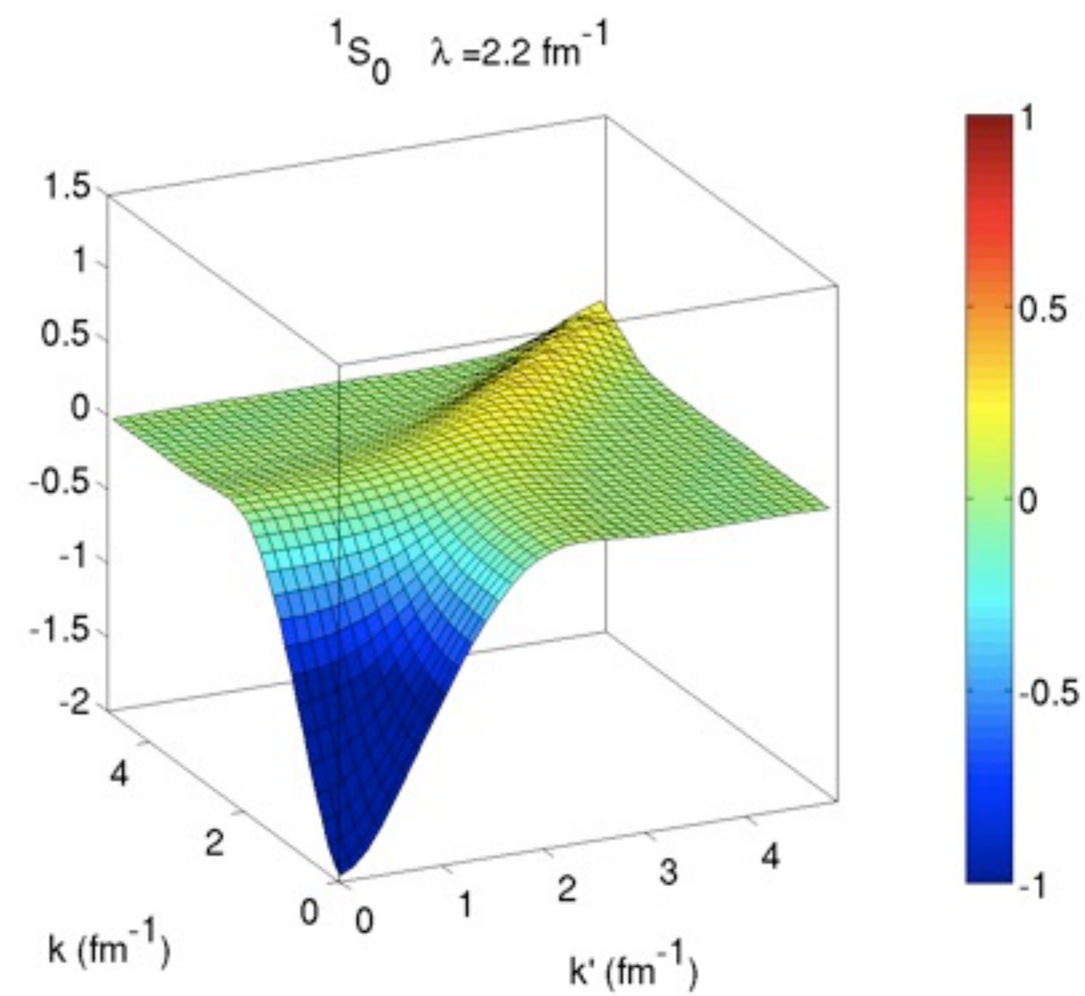

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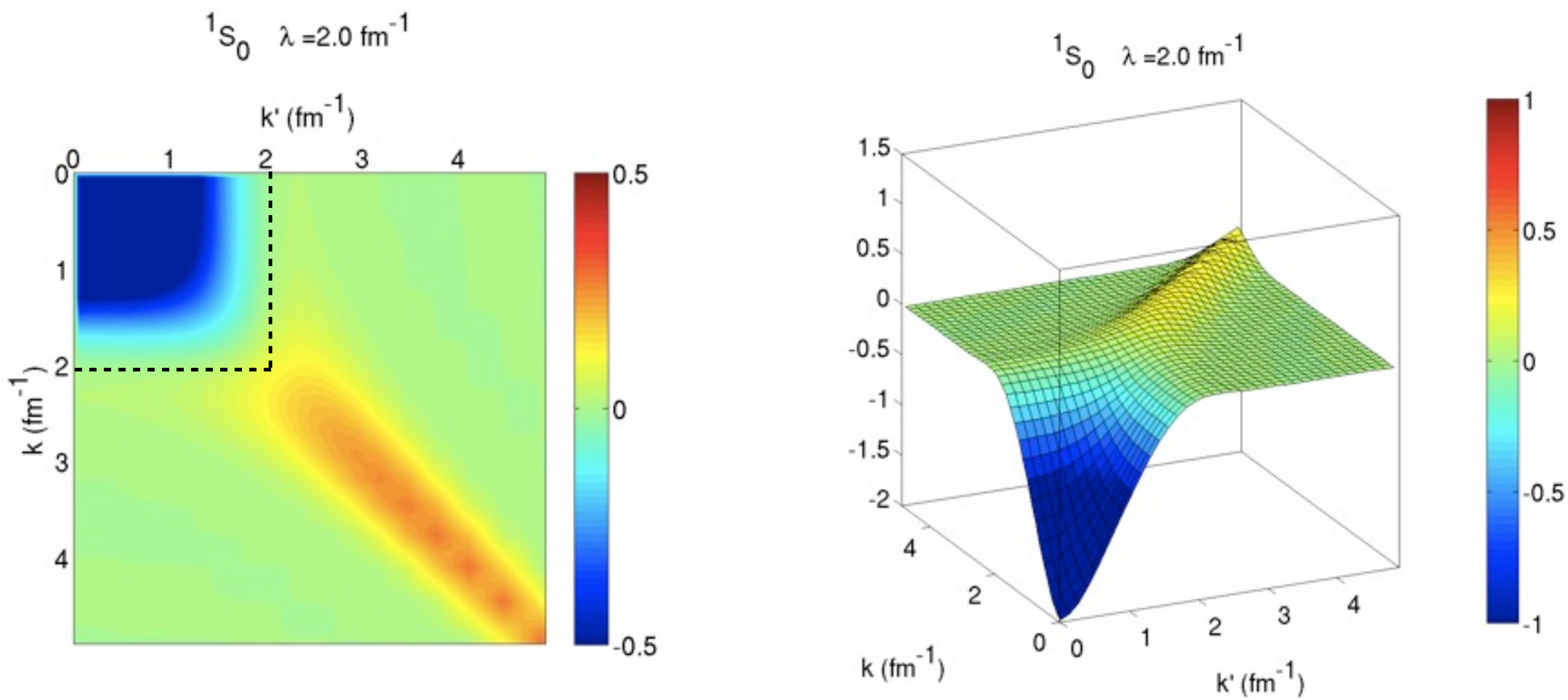
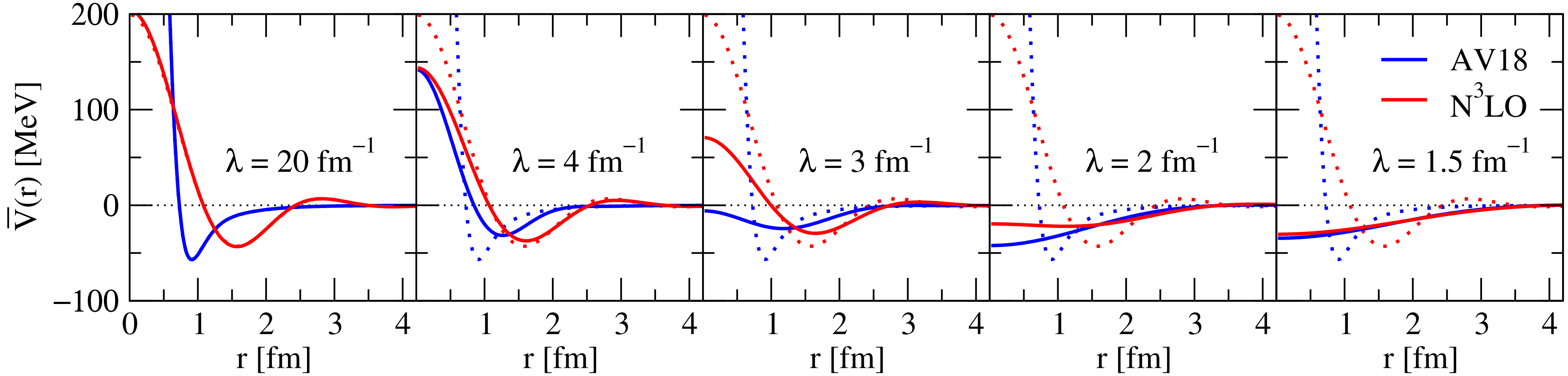
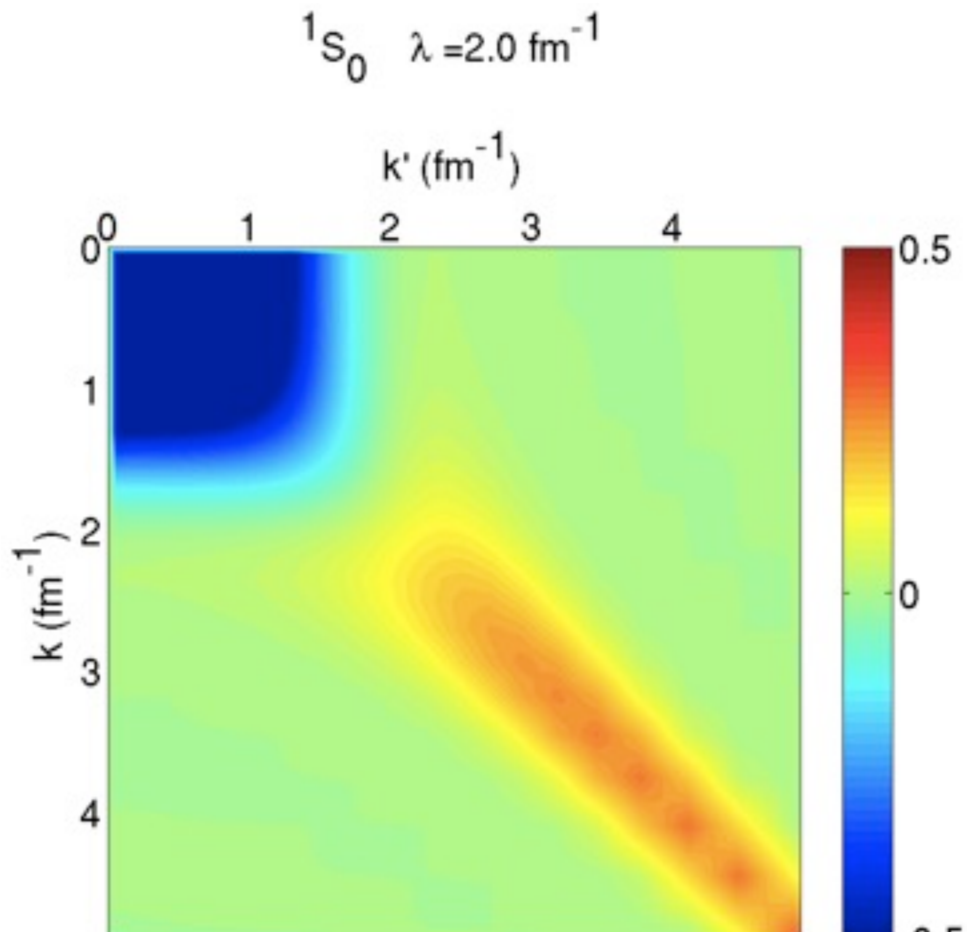
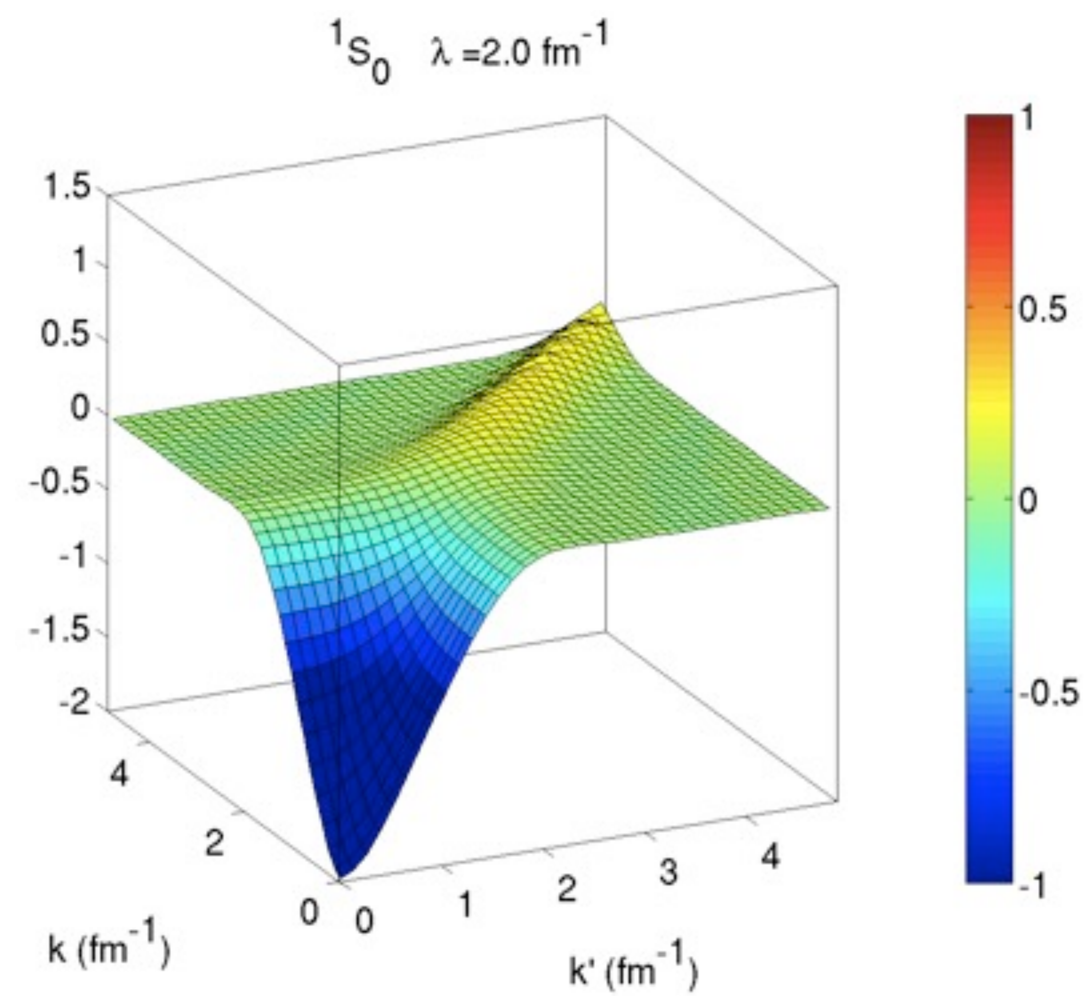


Systematic decoupling of high-momentum physics: The Similarity Renormalization Group

$$\overline{V}_\lambda(r) = \int dr' r'^2 V_\lambda(r, r')$$







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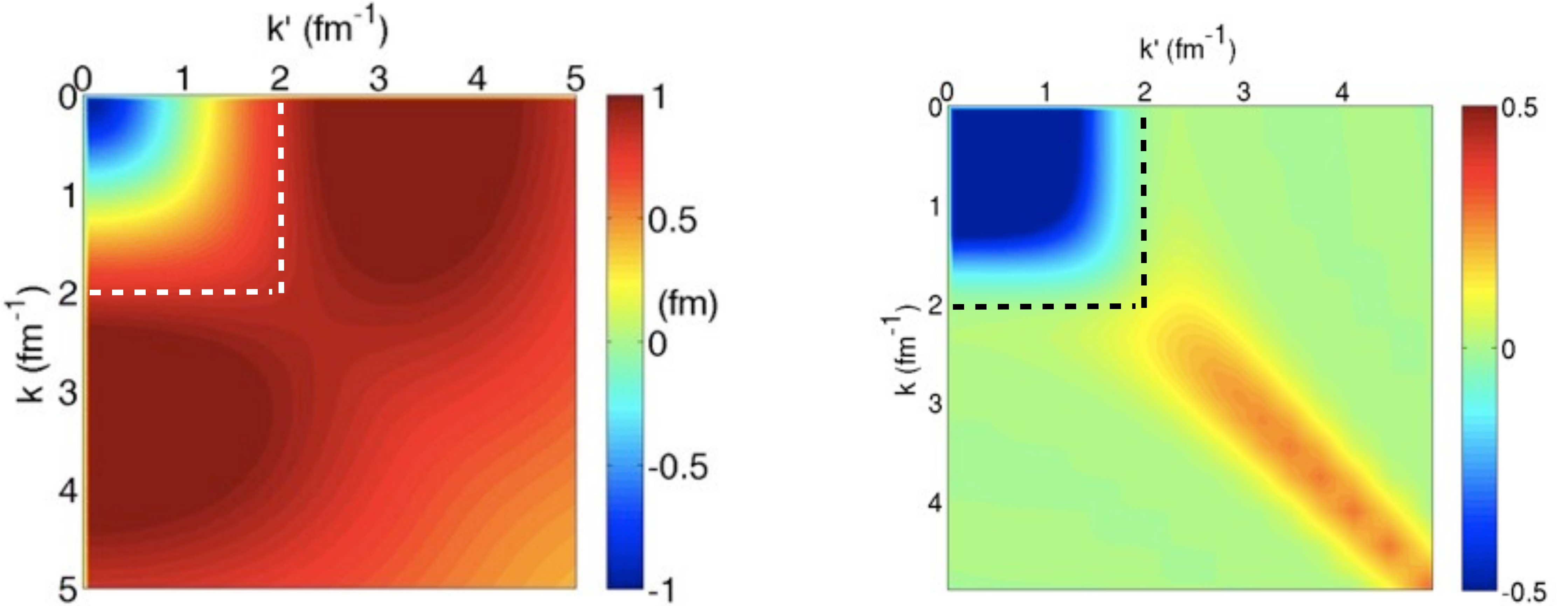
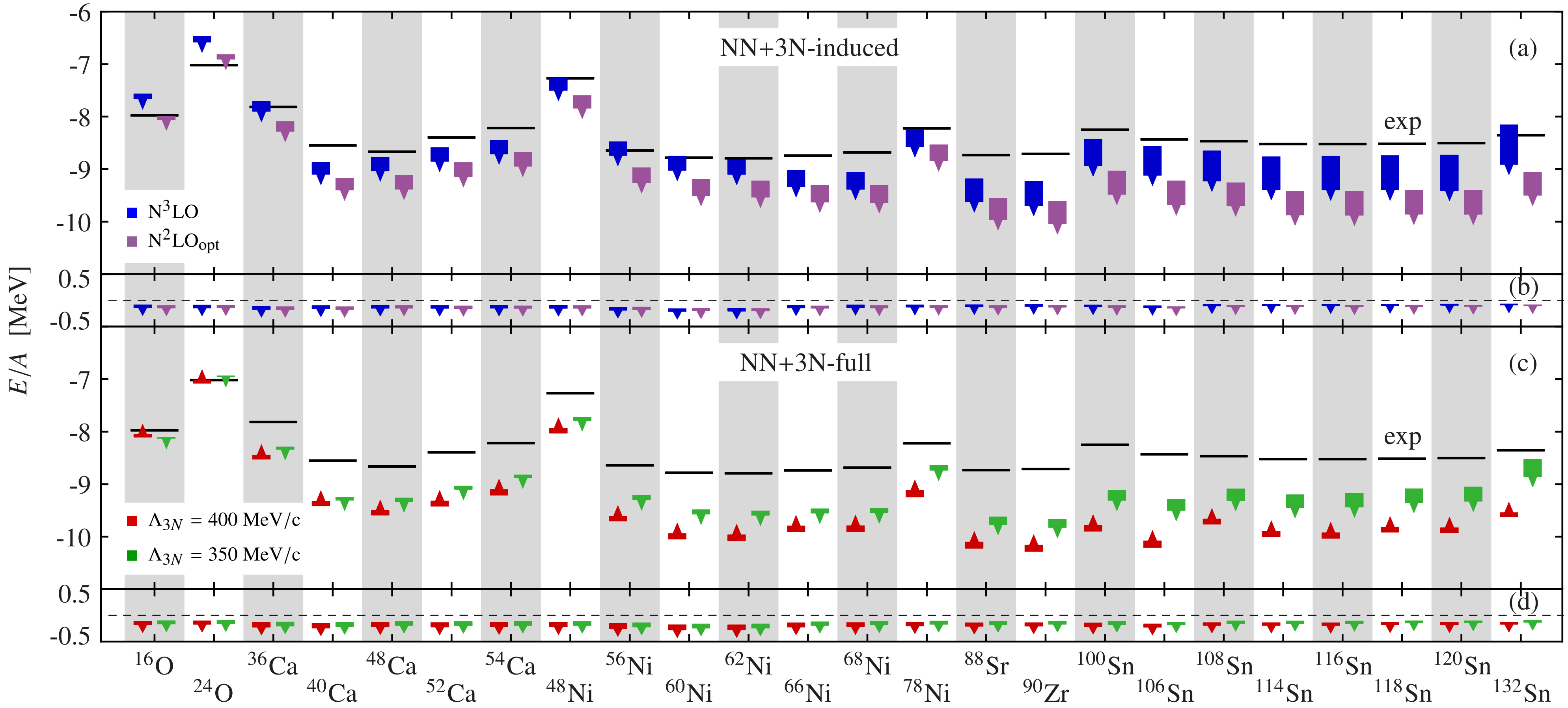
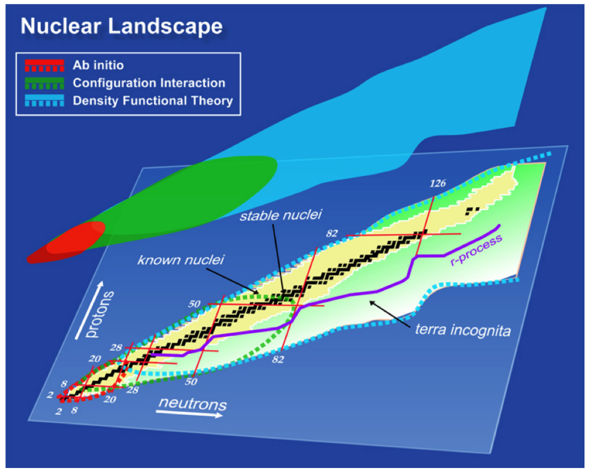


- elimination of coupling between low- and high momentum components,
  $\longrightarrow$ simplified many-body calculations, smaller required model spaces
- observables unaffected by resolution change (for exact calculations)
- residual resolution dependences can be used as tool to test calculations

Not the full story:
RG transformation also changes three-body (and higher-body) interactions.

Recent advances in ab-initio many-body theory





Binder et al. , Phys. Lett B 736, 119 (2014)

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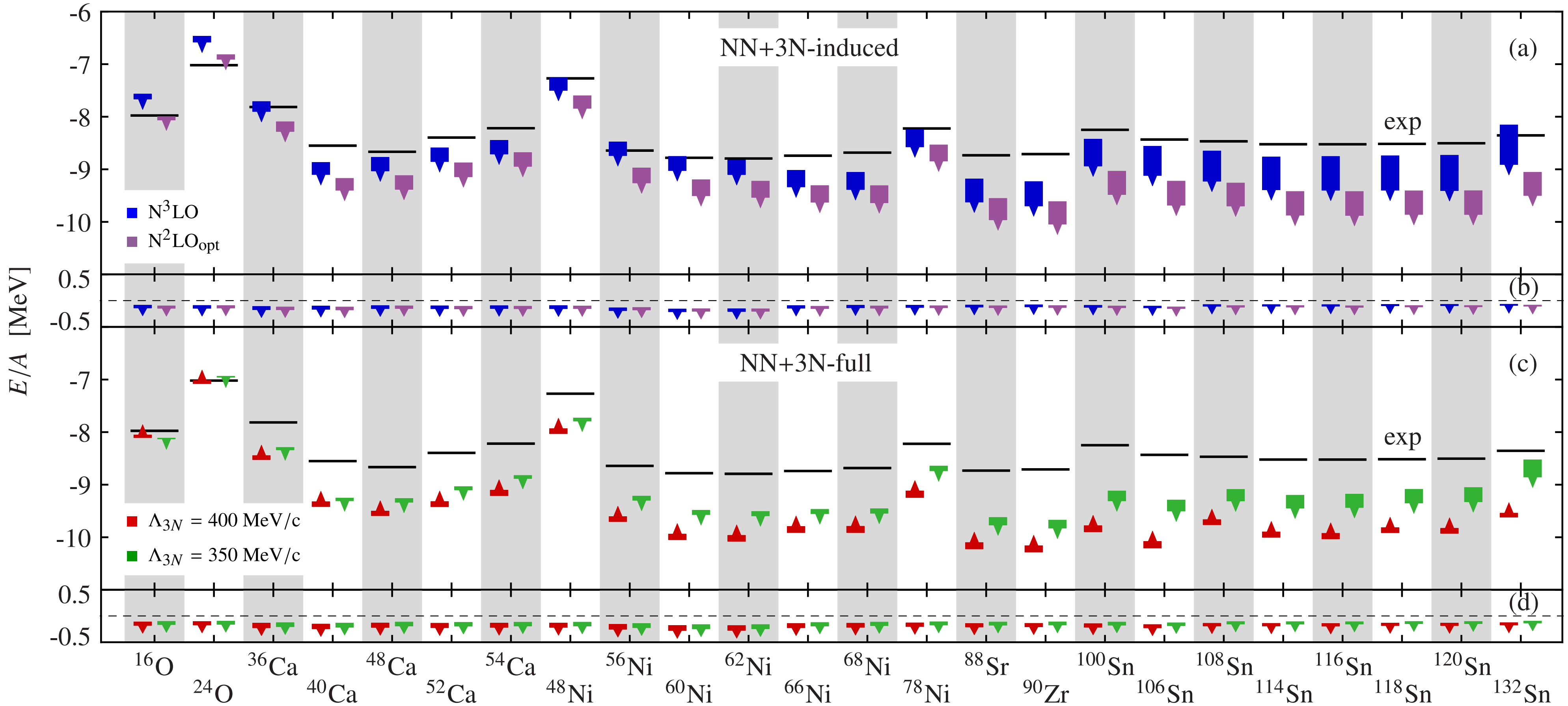

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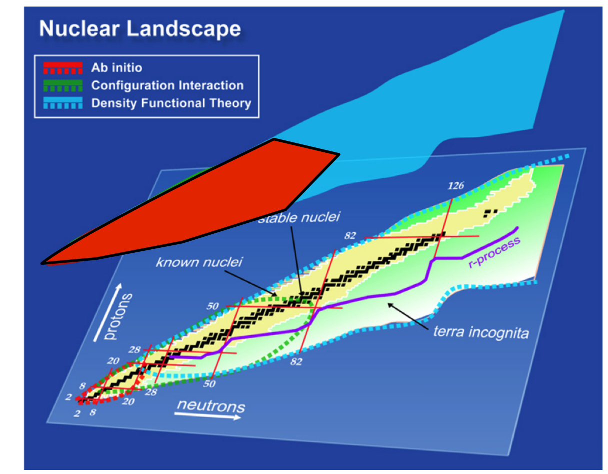
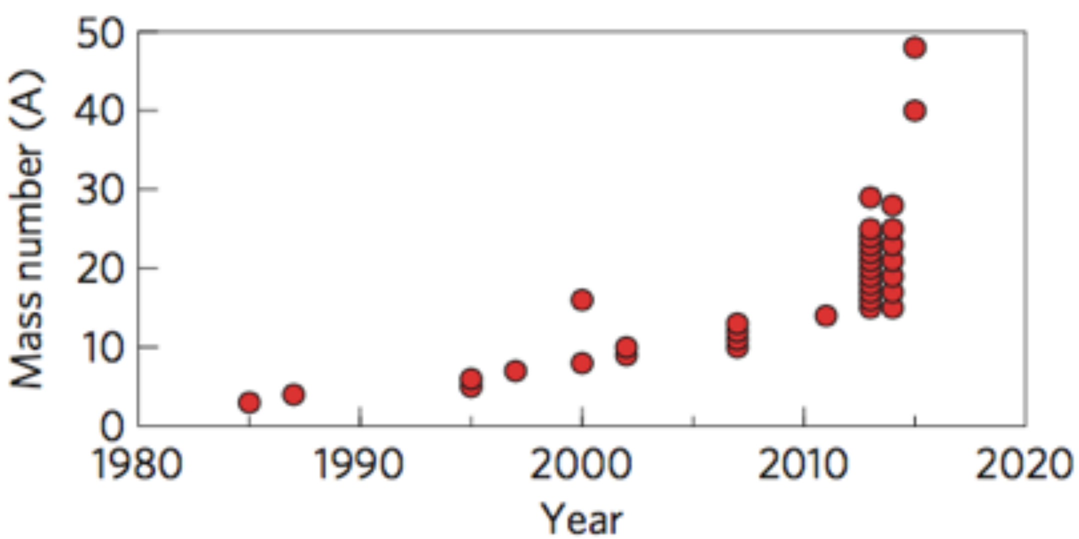
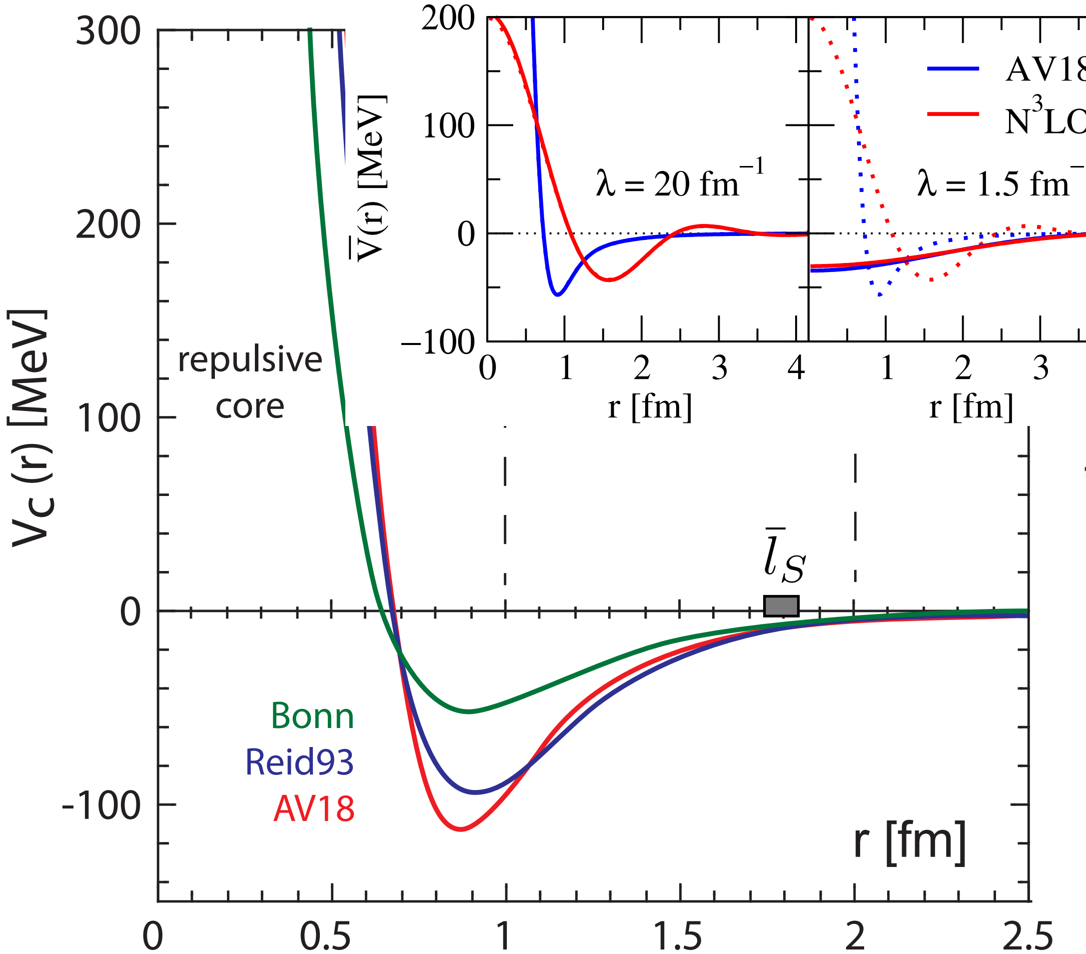


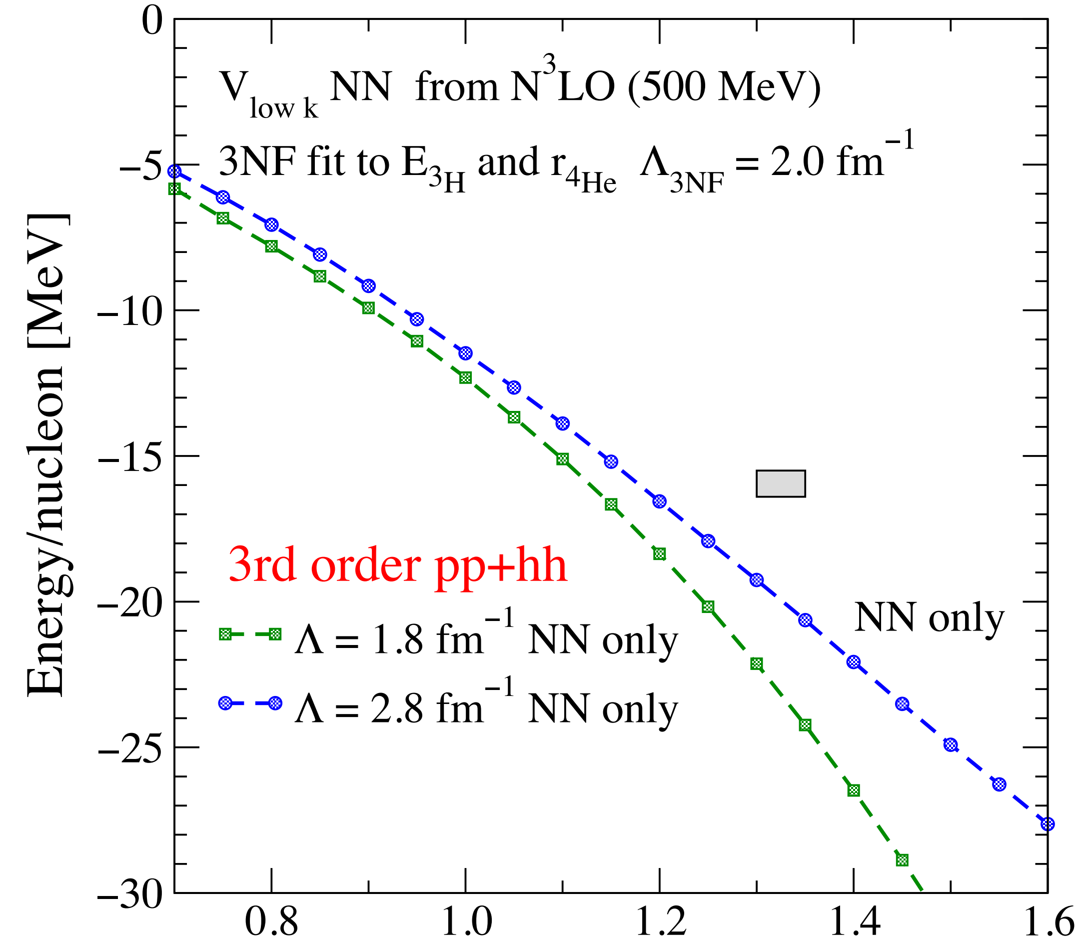


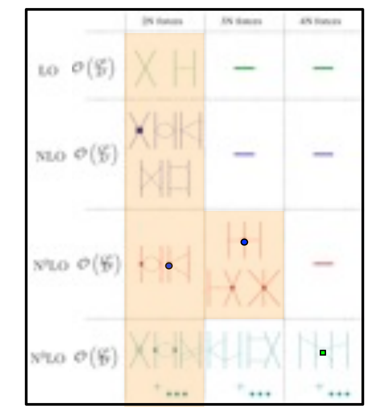
Hagen et al., Nature Physics 12, 186 (2016)

- spectacular increase in range of applicability of *ab initio* many body frameworks
- significant overbinding in heavy nuclei for presently used nuclear interactions

Fitting the 3NF LECs at low resolution scales







KH, Bogner, Furnstahl, Nogga,
PRC(R) 83, 031301 (2011)

"Very soft potentials must be excluded because they do not give saturation;
they give too much binding and too high density. In particular, a substantial tensor force is required."

Hans Bethe (1971)

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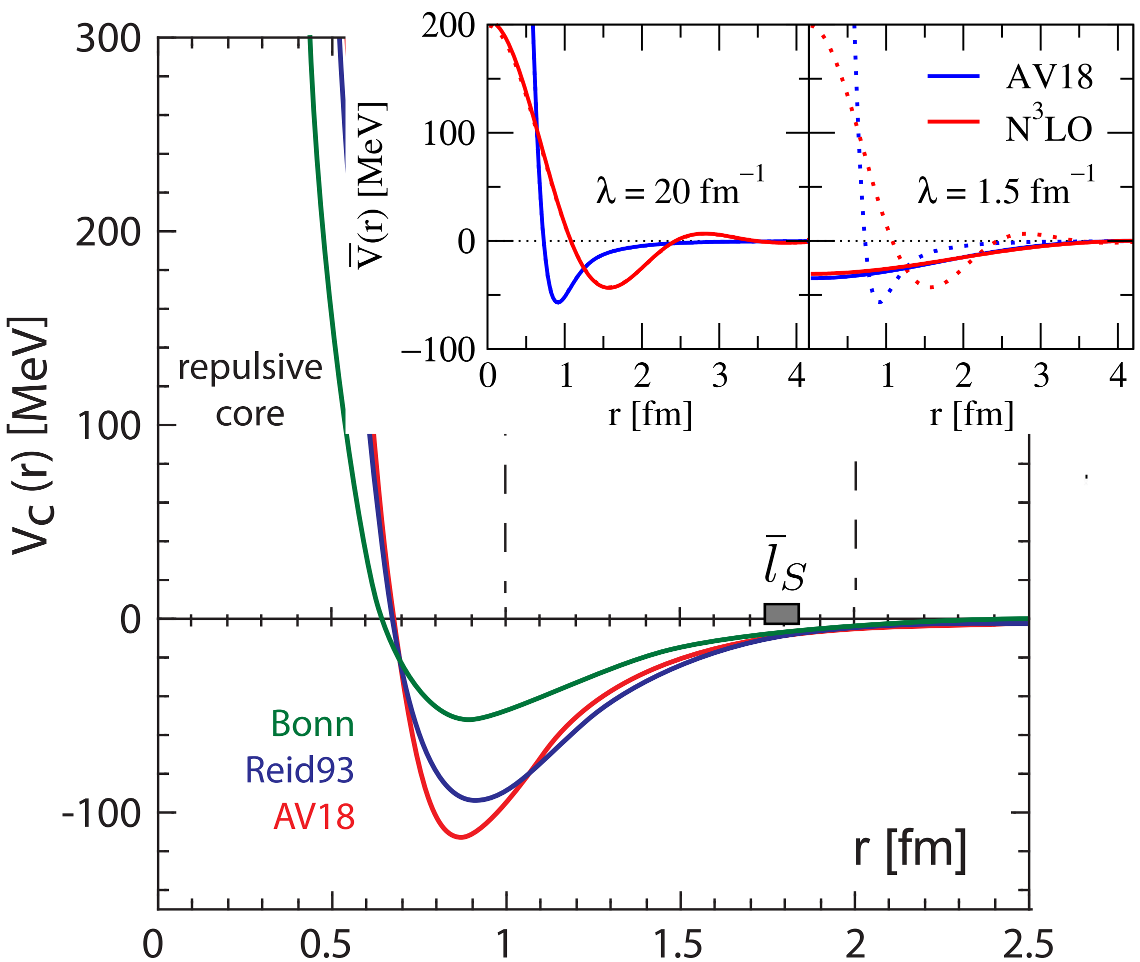

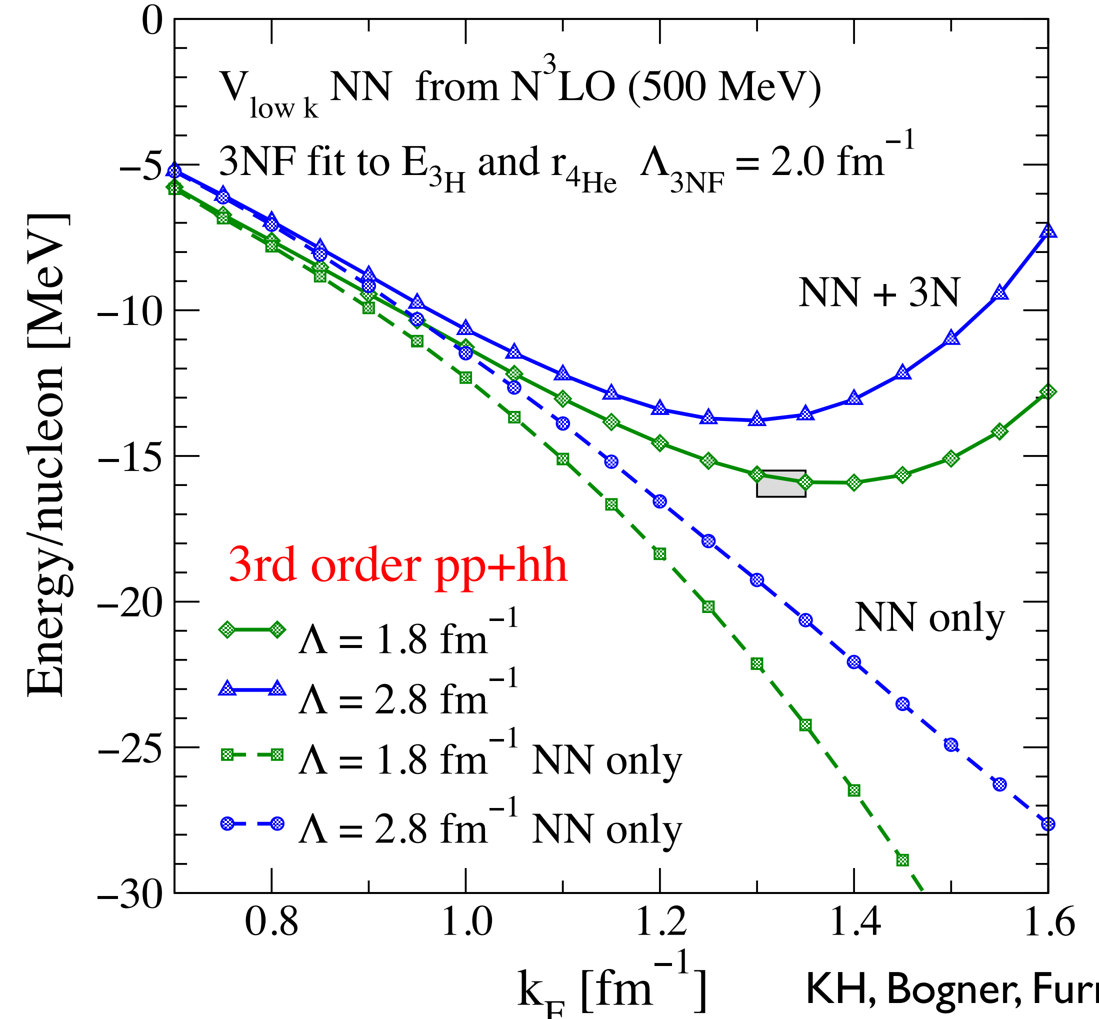

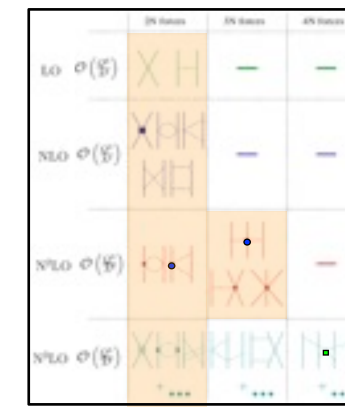

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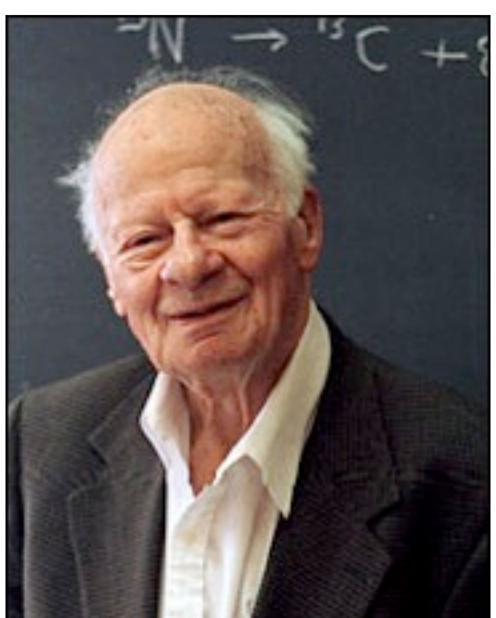

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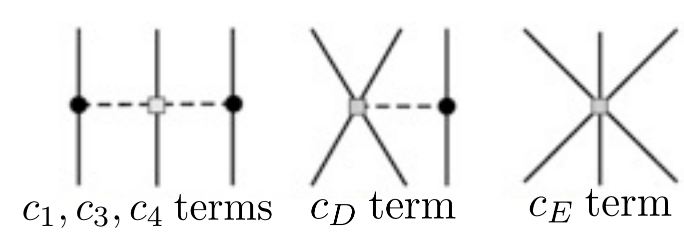
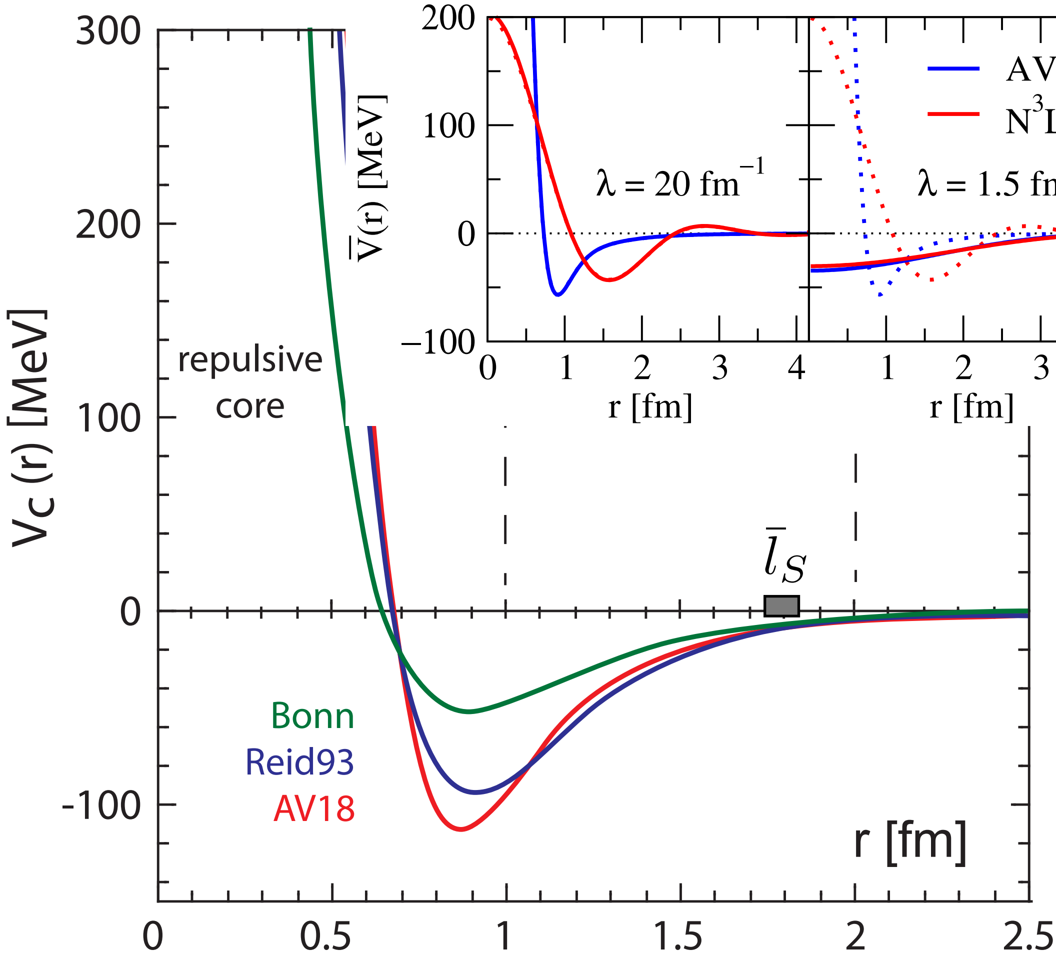
Hans Bethe (1971)

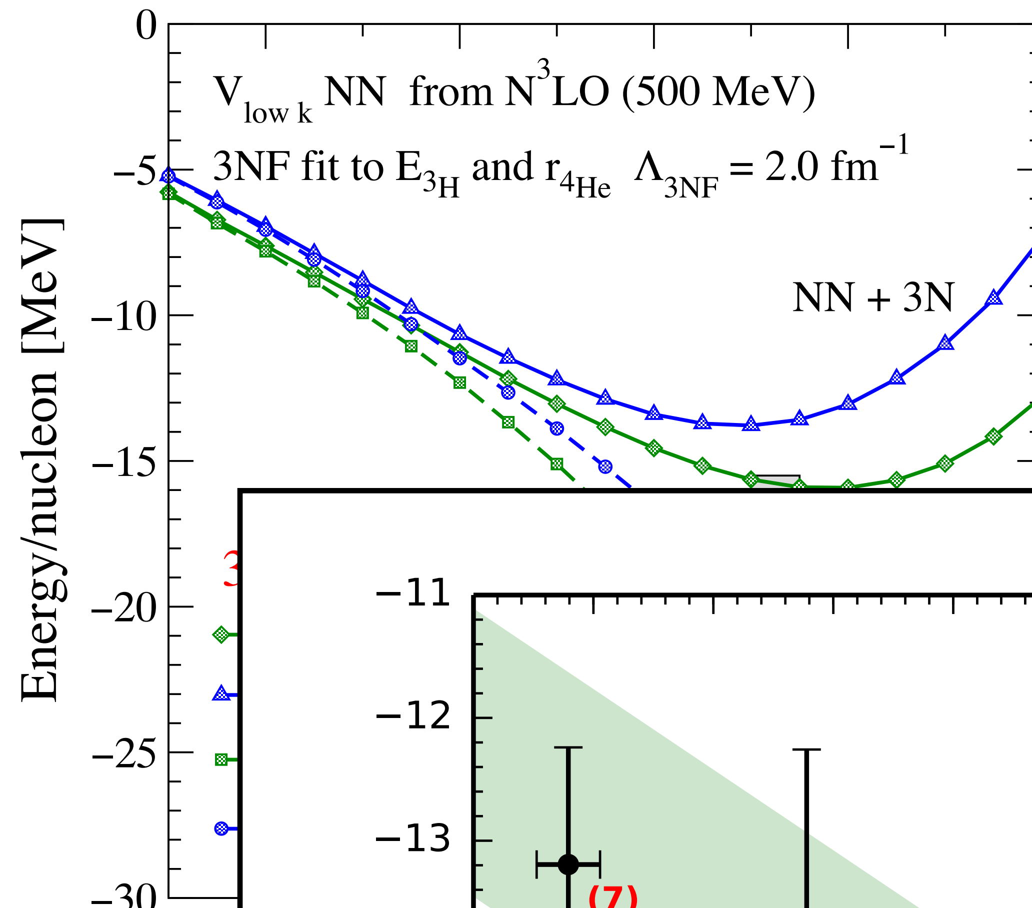
intermediate (c_D) and short-range (c_E) 3NF couplings fitted to few-body systems at different resolution scales:

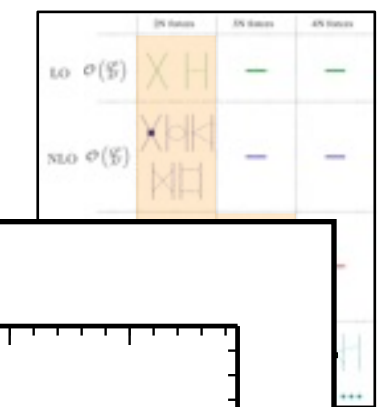
$$E_{^3\mathrm{H}} = -8.482\,\mathrm{MeV} \qquad r_{^4\mathrm{He}} = 1.464\,\mathrm{fm}$$

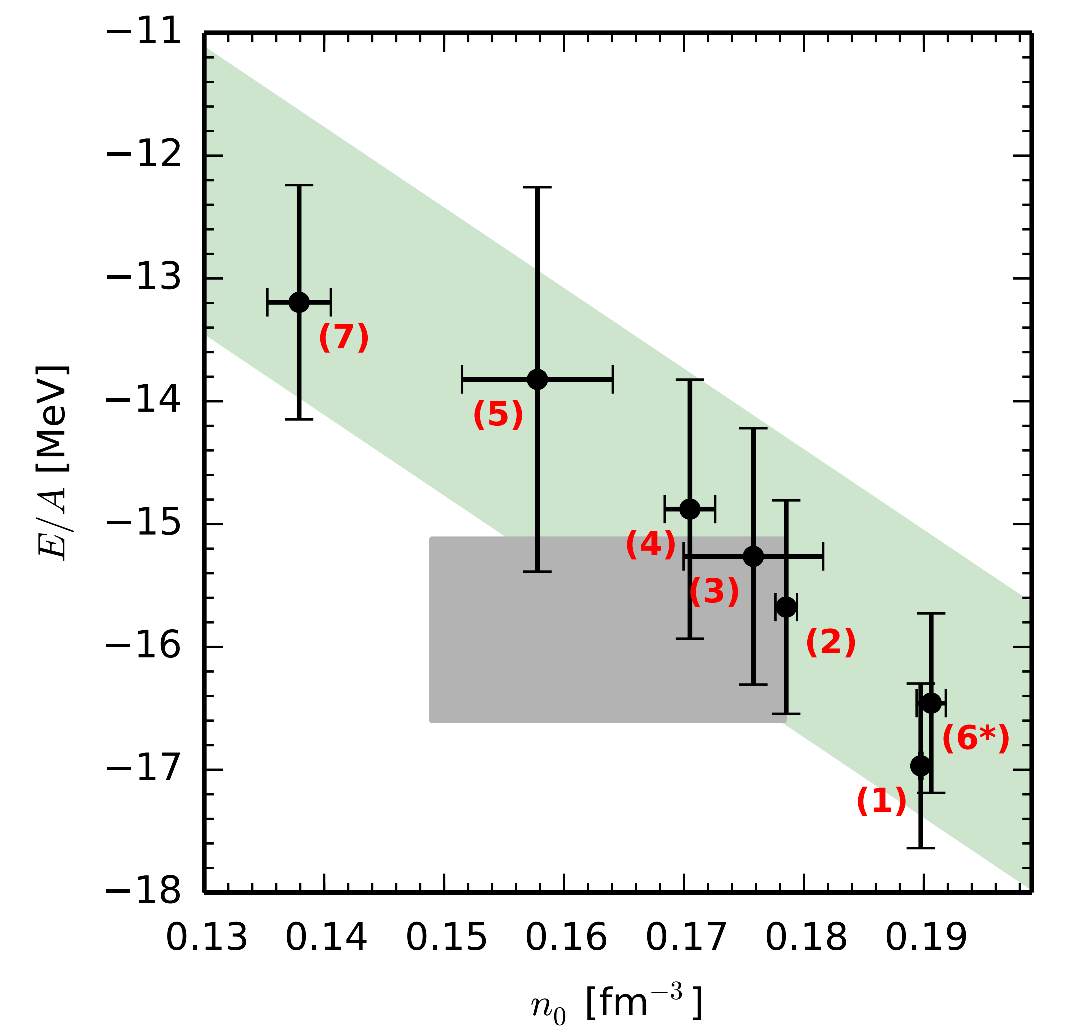


Fitting the 3NF LECs at low resolution scales







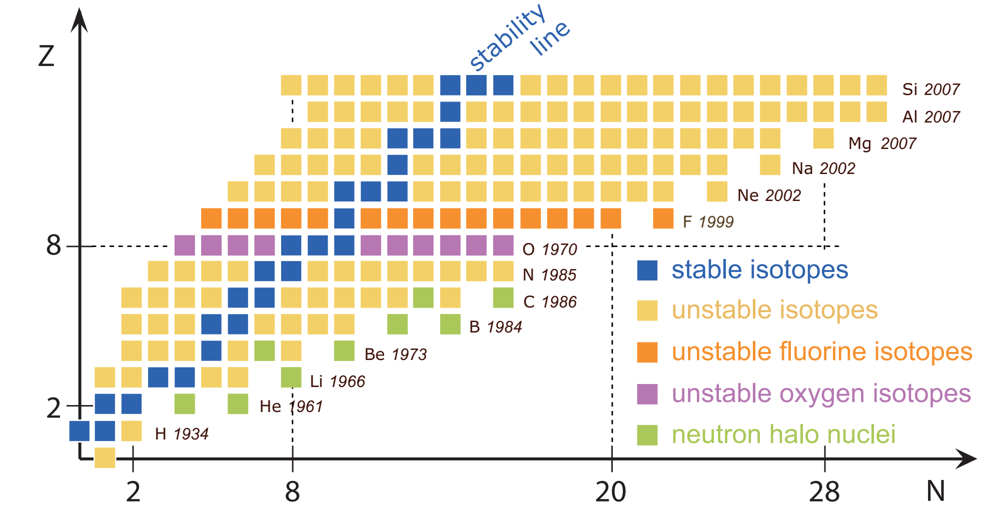


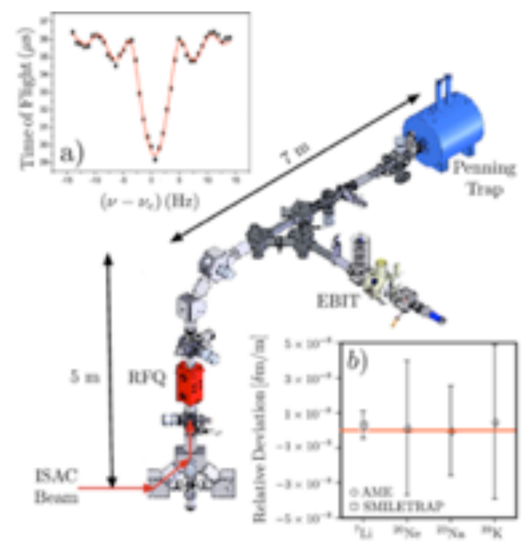
Drischler, KH, Schwenk, PRC93, 054314 (2016)

"Very soft potentials must be excluded because they do not give saturation; they give too much binding and too high density. In particular, a substantial tensor force is required."

Hans Bethe (1971)

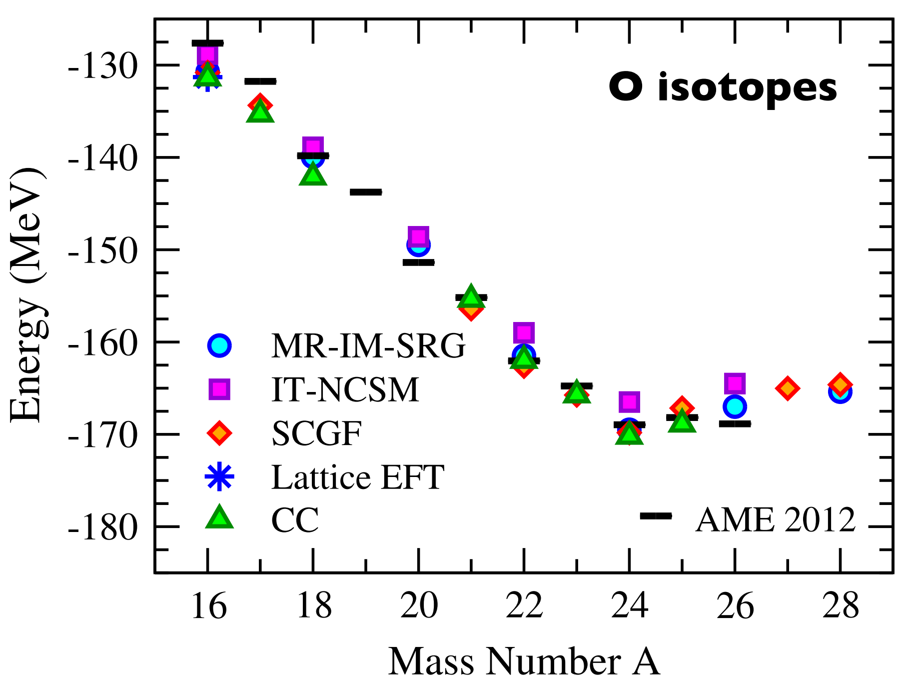
Studies of neutron-rich nuclei

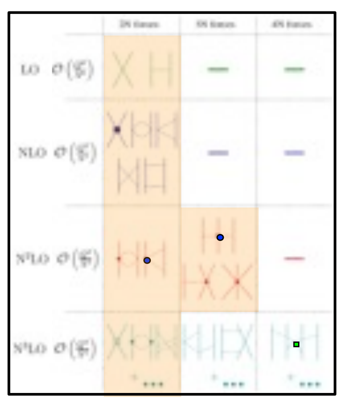


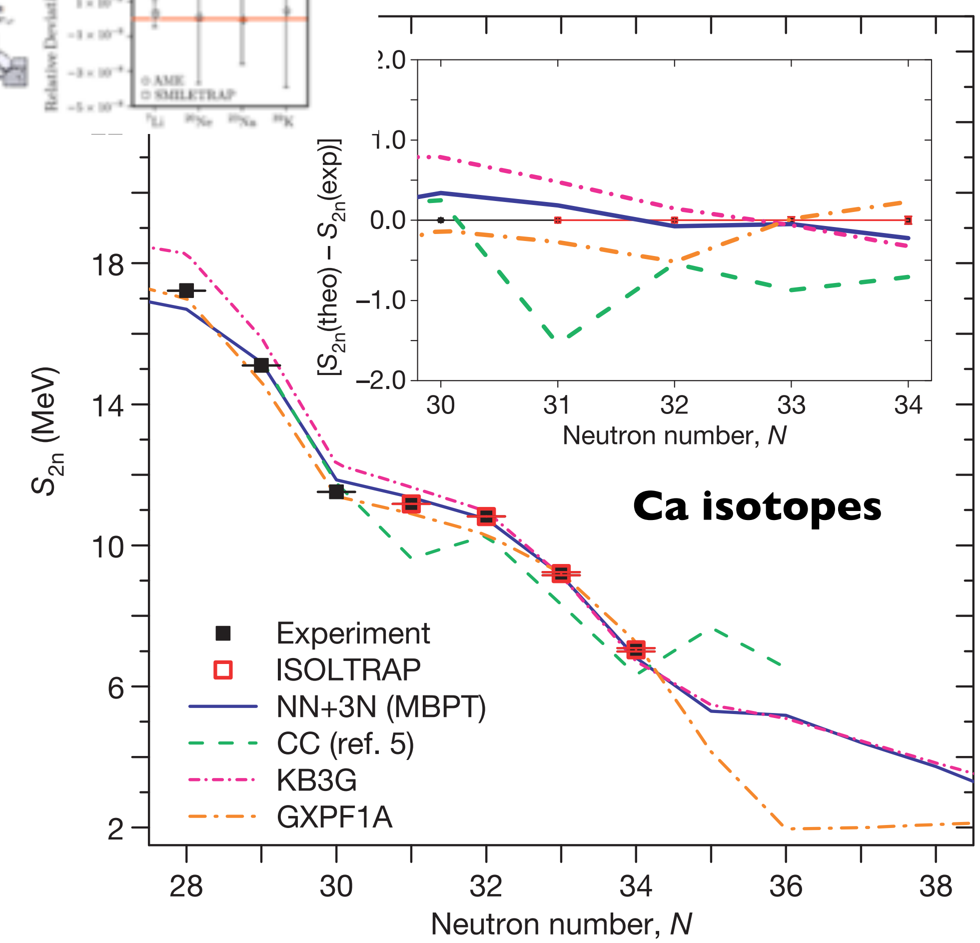


Gallant et al.
PRL 109, 032506 (2012)

Wienholtz et al.
Nature 498, 346 (2013)







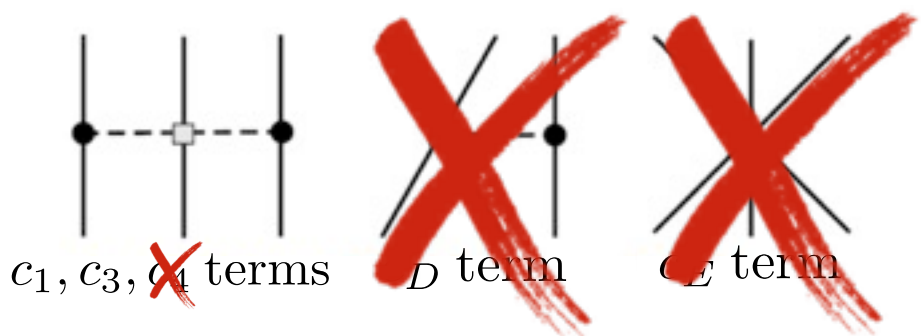
KH et al., Ann. Rev. Nucl. Part. Sci. 165, 457 (2015)

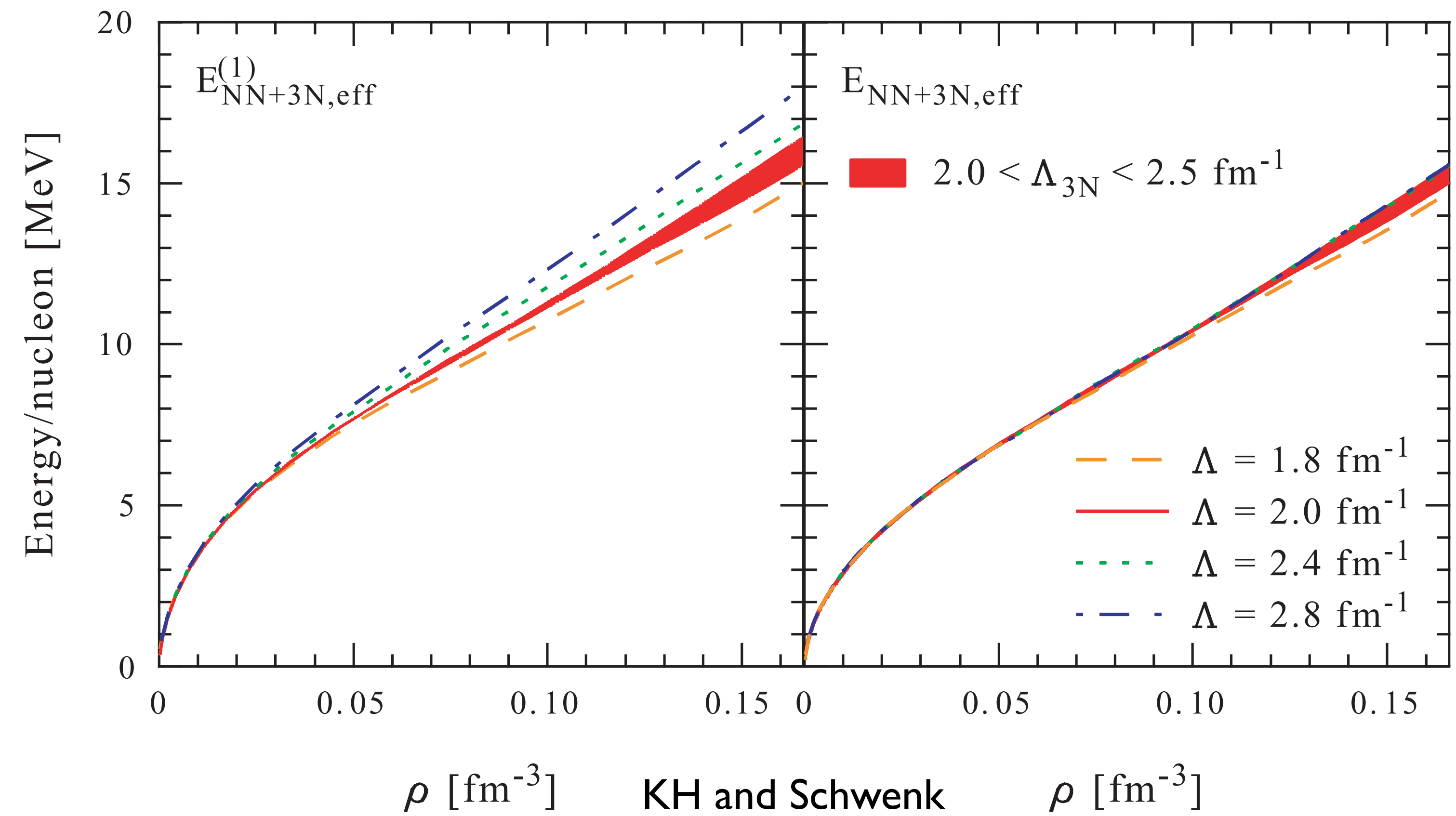
- remarkable agreement between different many-body frameworks
- excellent agreement between theory and experiment for masses of oxygen and calcium isotopes based on specific chiral interactions
- need to quantify theoretical uncertainties

Results for the neutron matter equation of state

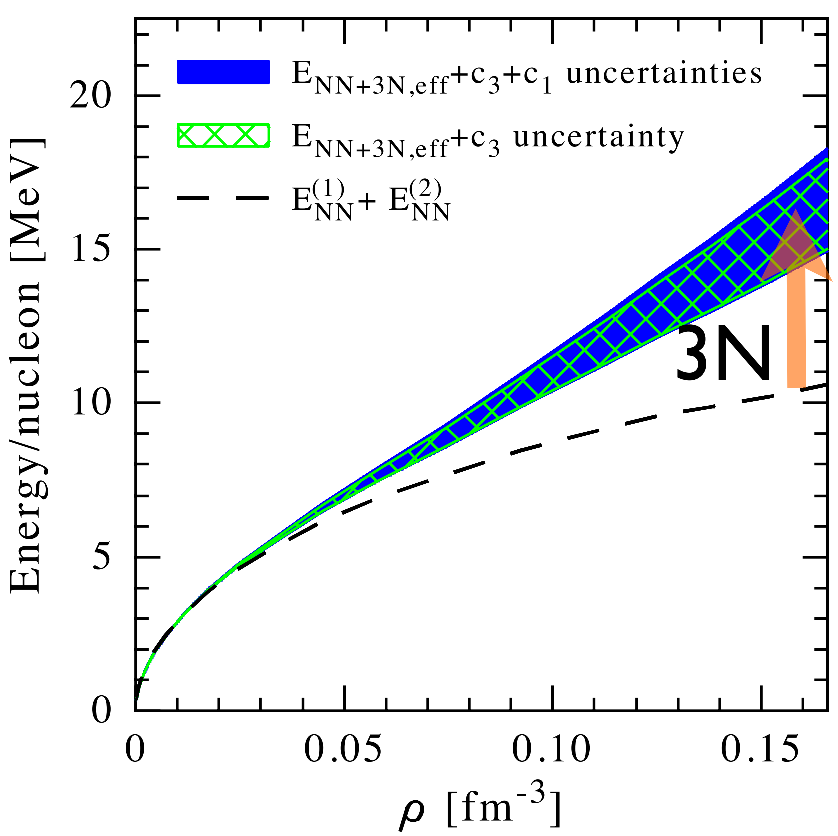
neutron matter is a **unique system** for chiral EFT:

only long-range 3NF contribute in leading order



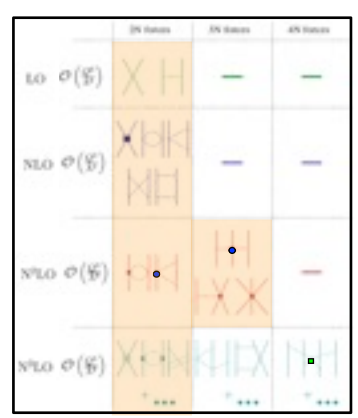


KH and Schwenk
PRC 82, 014314 (2010)



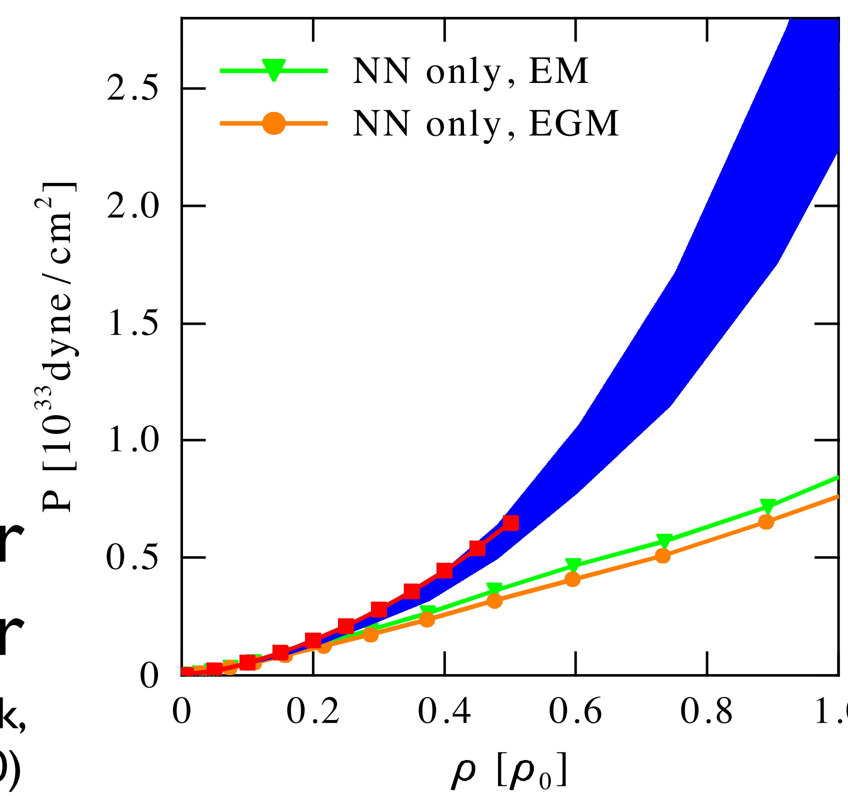
pure neutron matter

KH and Schwenk PRC 82, 014314 (2010)

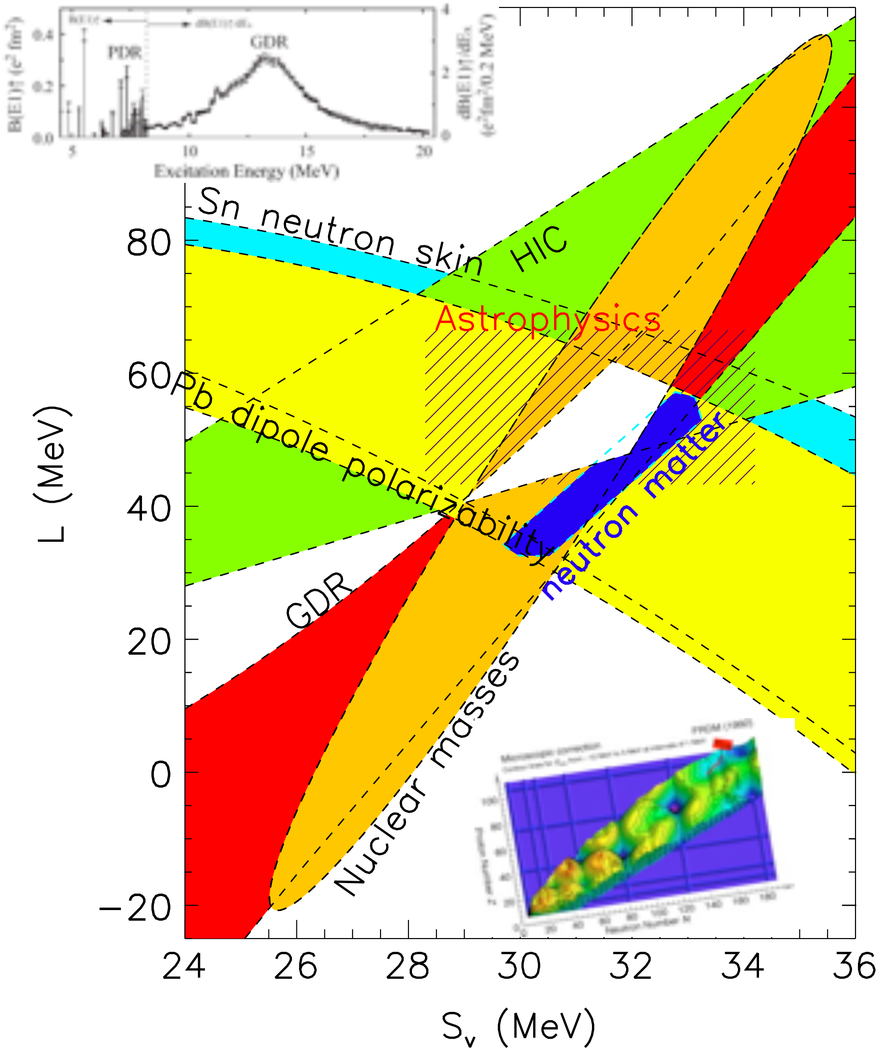


neutron star matter

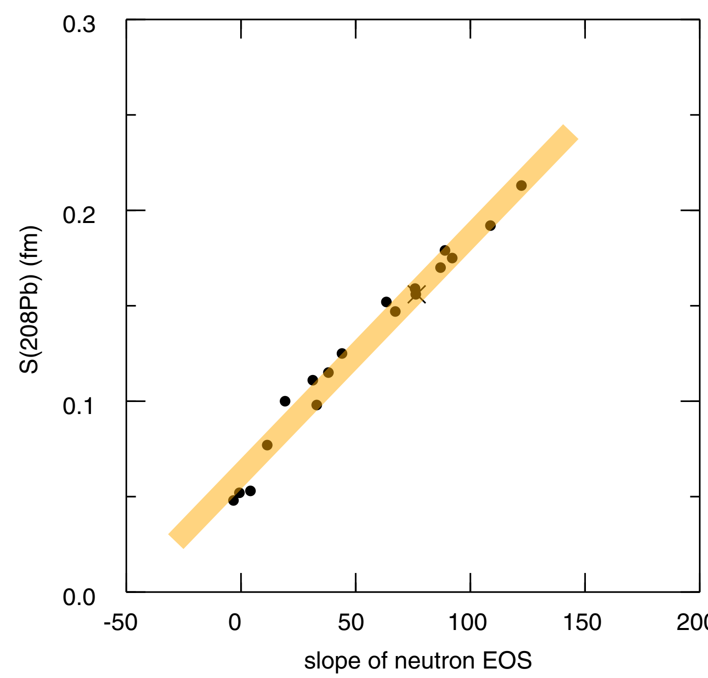
KH, Lattimer, Pethick, Schwenk,
PRL 105, 161102 (2010)



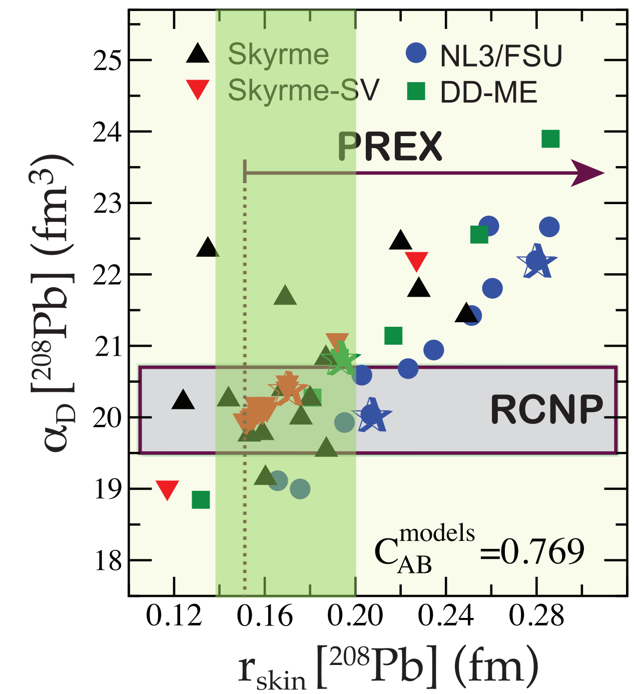
Symmetry energy and neutron skin constraints



KH, Lattimer, Pethick, Schwenk, ApJ 773, 11 (2013)



Brown,
PRL 85, 5296 (2000)



Piekarewicz,
PRC 85, 041302 (2012)

$$S_v = \left.\frac{\partial^2 E/N}{\partial^2 x}\right|_{\rho=\rho_0, x=1/2}$$

$$L = \frac{3}{8} \left.\frac{\partial^3 E/N}{\partial\rho\partial^2 x}\right|_{\rho=\rho_0, x=1/2}$$

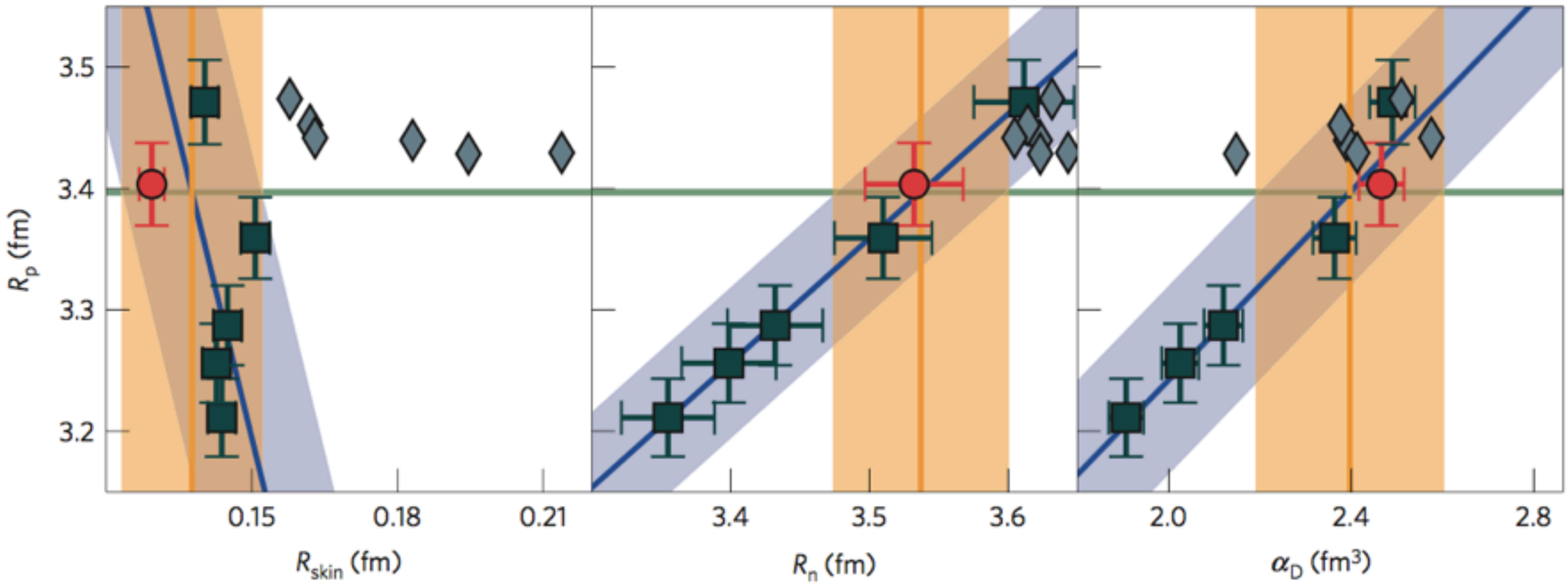
neutron skin constraint from neutron matter results:

$$r_{\rm skin}[^{208}{\rm Pb}] = 0.14 - 0.2 \text{ fm}$$

KH, Lattimer, Pethick, Schwenk, PRL 105, 161102 (2010)

- neutron matter give tightest constraints
- in agreement with all other constraints

Predictions for the neutron skin of ^{48}Ca

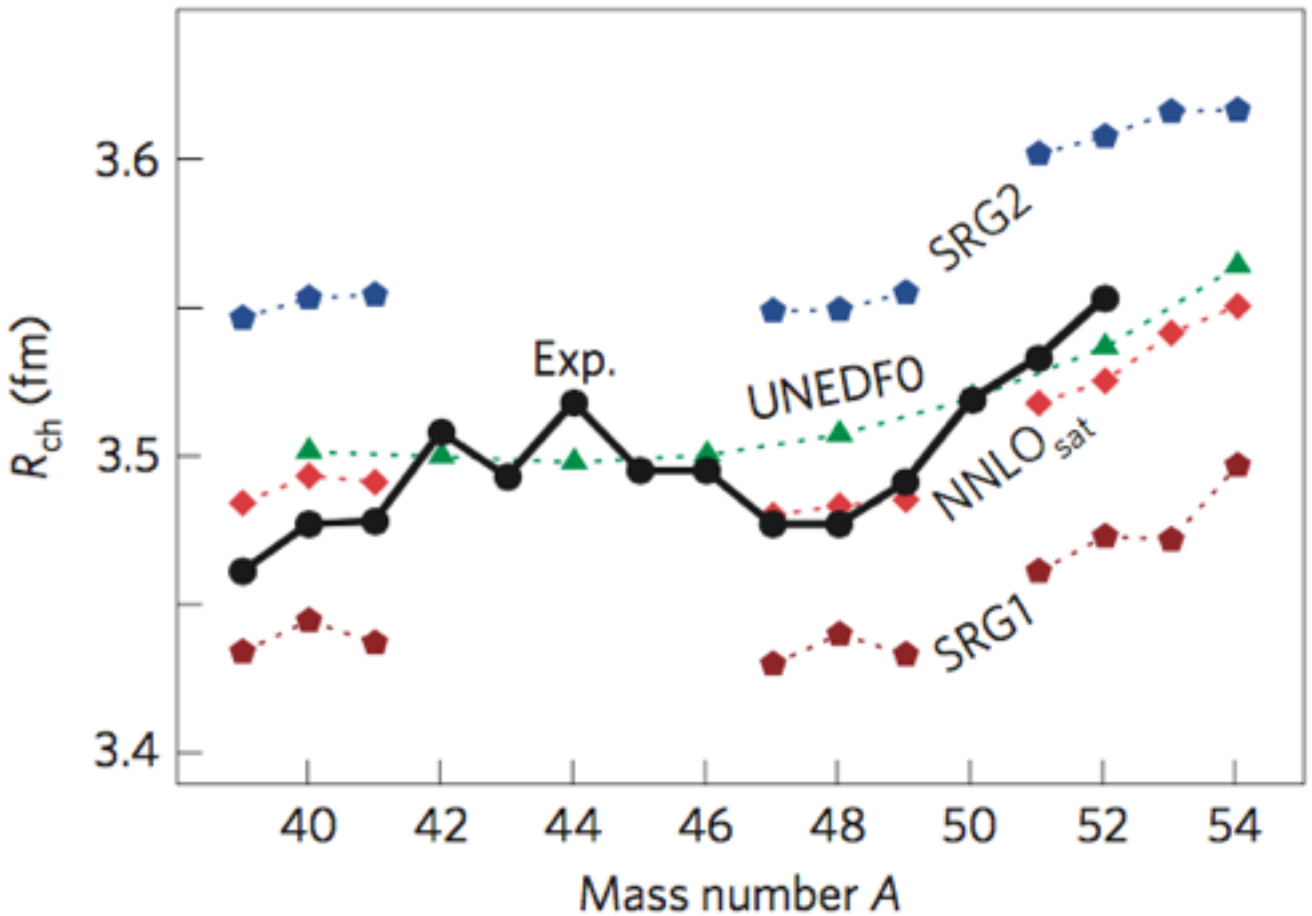


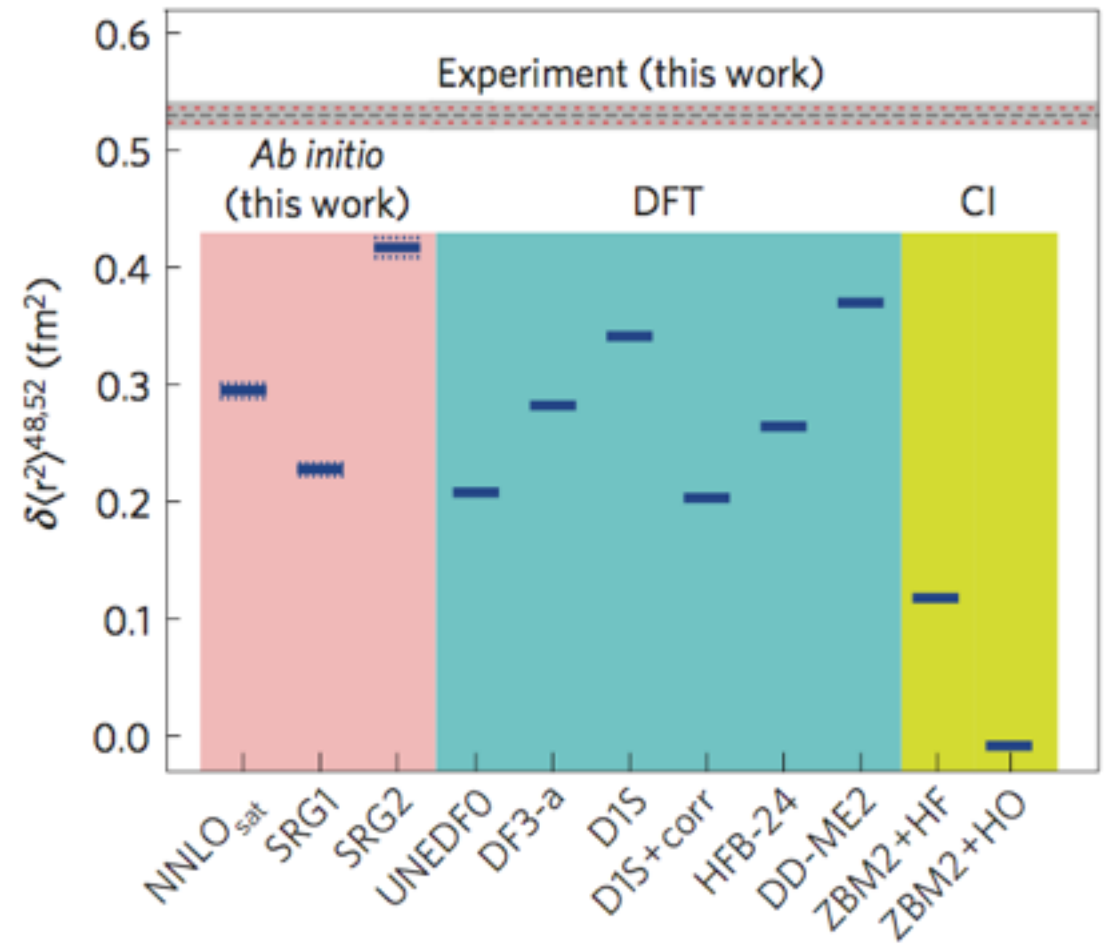
Hagen et al., Nature Physics 12, 186 (2016)

- microscopic coupled cluster results based on a set of different nuclear NN+3N interactions (see also Phys. Rev. C91, 051301 (2015))
- correlations between different observables and the precisely measured R_p
- prediction of significantly smaller neutron skin compared to EDF results:

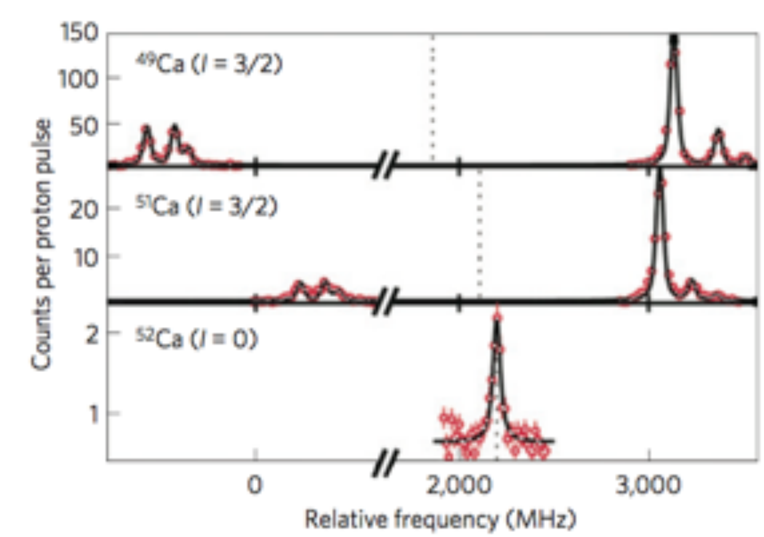
$$0.12 \lesssim R_{\text{skin}} \lesssim 0.15\,\text{fm}$$

Charge radii of calcium isotopes





Garcia Ruiz at al., Nature Physics (advanced online, 2016)



- novel precise measurements of Ca isotope shifts
- unexpectedly large radii of neutron-rich isotopes
- reasonable theoretical reproduction of radius trends

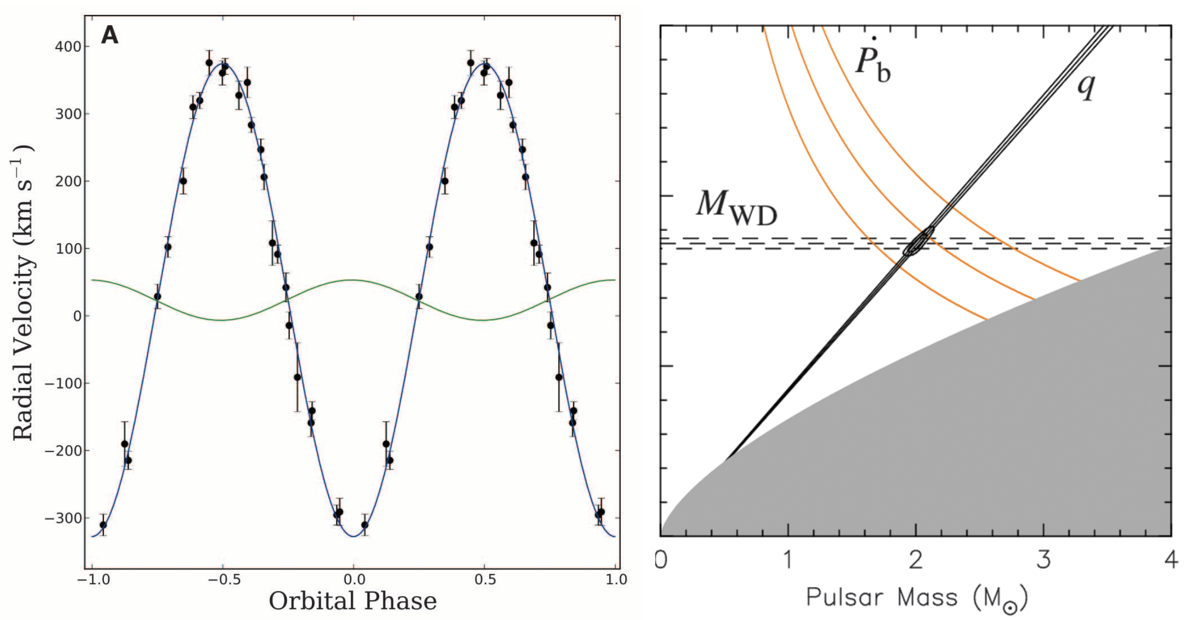
in coupled cluster calculations based on chiral EFT interactions

- radius increase quantitatively underestimated in all theoretical studies

Constraints on the nuclear equation of state (EOS)

Science

A Massive Pulsar in a Compact Relativistic Binary

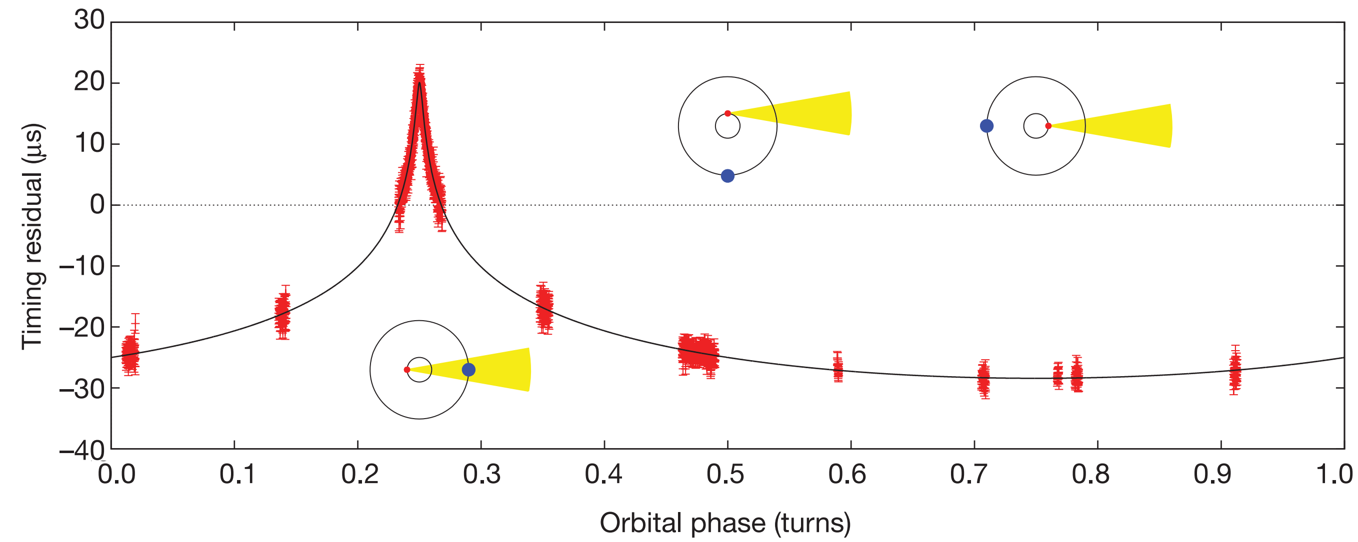


Antoniadis et al., Science 340, 448 (2013)

nature

A two-solar-mass neutron star measured using Shapiro delay

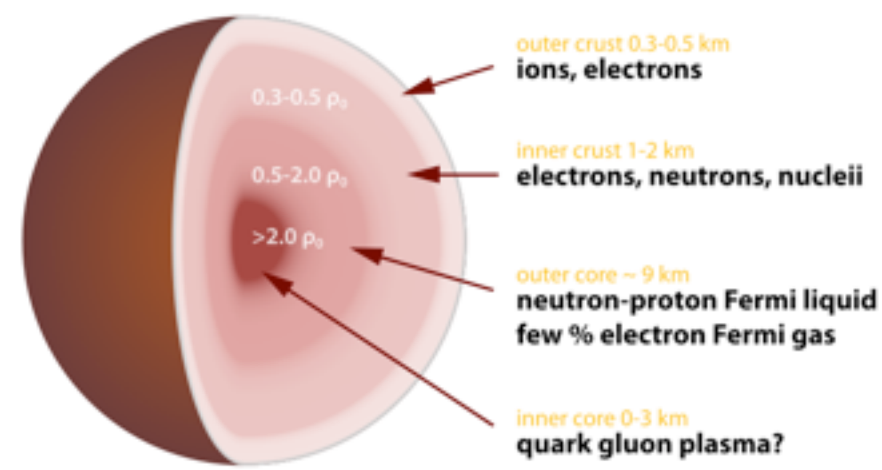
P. B. Demorest[1], T. Pennucci[2], S. M. Ransom[1], M. S. E. Roberts[3] & J. W. T. Hessels[4,5]



Demorest et al., Nature 467, 1081 (2010)

New constraints from recent observations:

$$M_{\rm max} = 1.65 M_\odot \rightarrow 1.97 \pm 0.04\, M_\odot$$
$$\rightarrow 2.01 \pm 0.04\, M_\odot$$



Calculation of neutron star properties require EOS up to high densities.

Strategy:

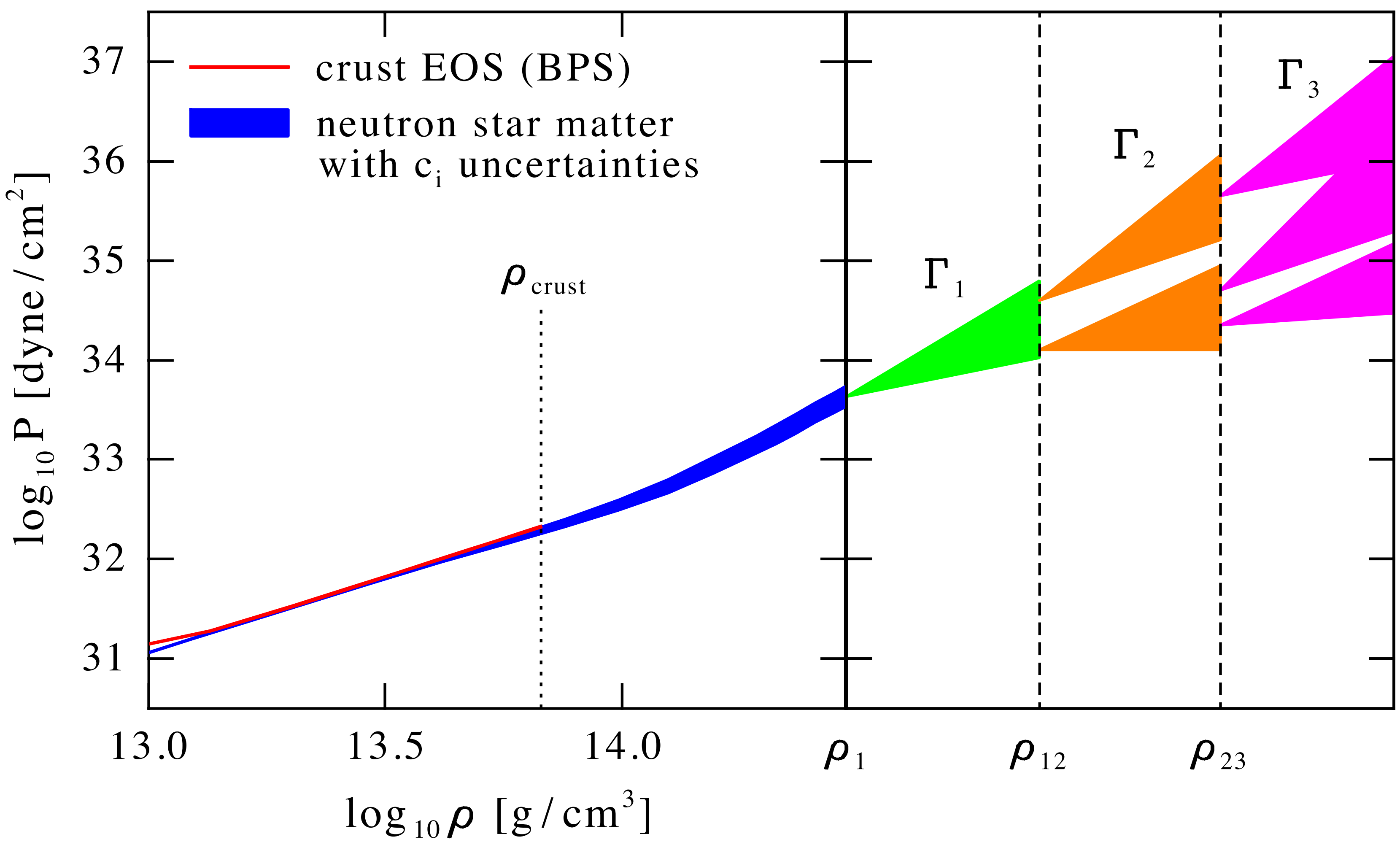
Use observations to constrain the high-density part of the nuclear EOS.

Neutron star radius constraints

incorporation of beta-equilibrium: neutron matter $\longrightarrow$ neutron star matter

parametrize piecewise high-density extensions of EOS:

- use polytropic ansatz $p \sim \rho^{\Gamma}$
- range of parameters $\Gamma_1, \rho_{12}, \Gamma_2, \rho_{23}, \Gamma_3$ limited by physics



KH, Lattimer, Pethick, Schwenk, ApJ 773, 11 (2013)
KH, Lattimer, Pethick, Schwenk, PRL 105, 161102 (2010)

Constraints on the nuclear equation of state

use the constraints:

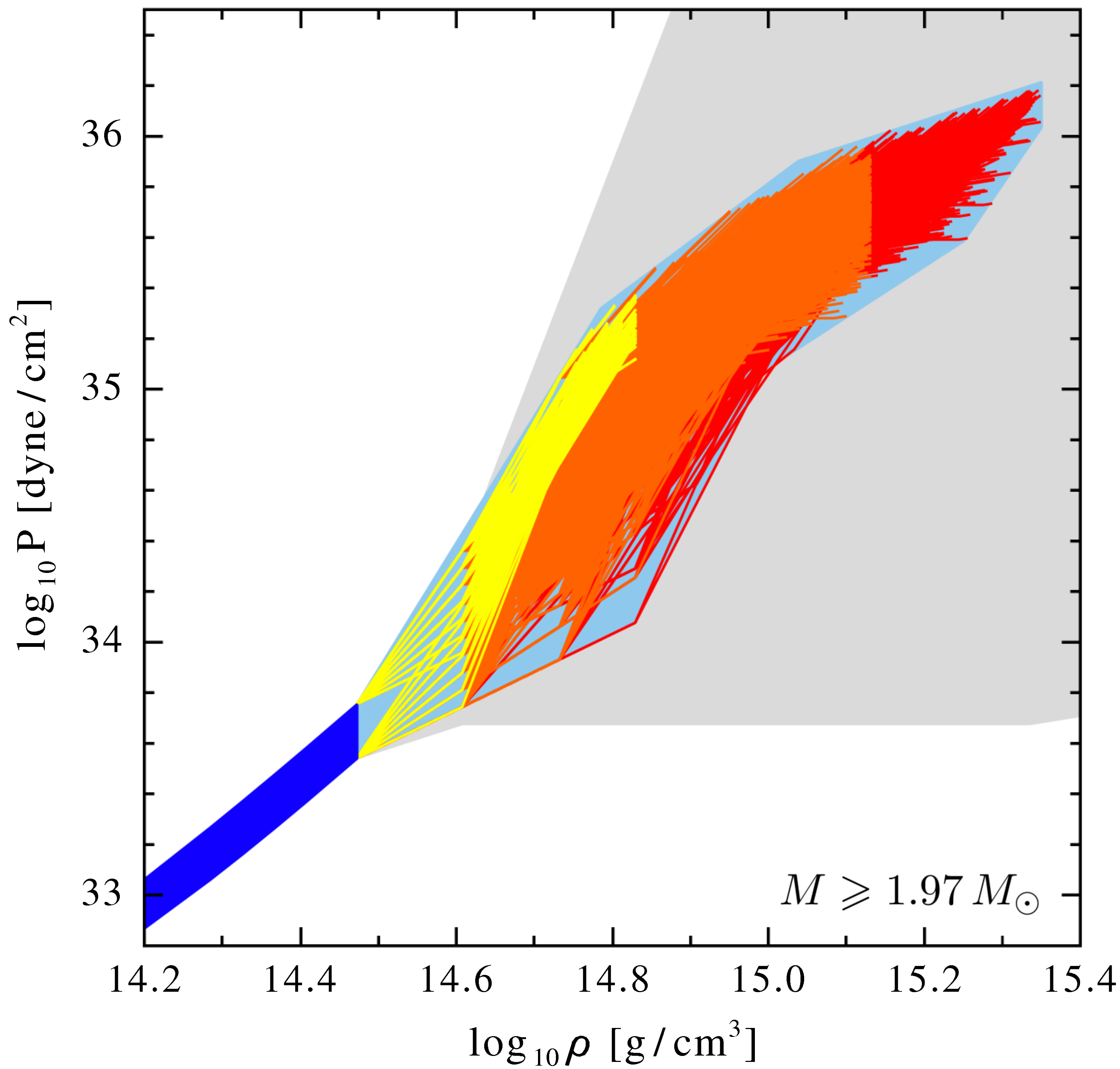
recent NS observations

$$M_{\max} > 1.97\, M_\odot$$

causality

$$v_s(\rho) = \sqrt{dP/d\varepsilon} < c$$

KH, Lattimer, Pethick, Schwenk, ApJ 773, 11 (2013)



constraints lead to significant reduction of EOS uncertainty band

Constraints on the nuclear equation of state

use the constraints:

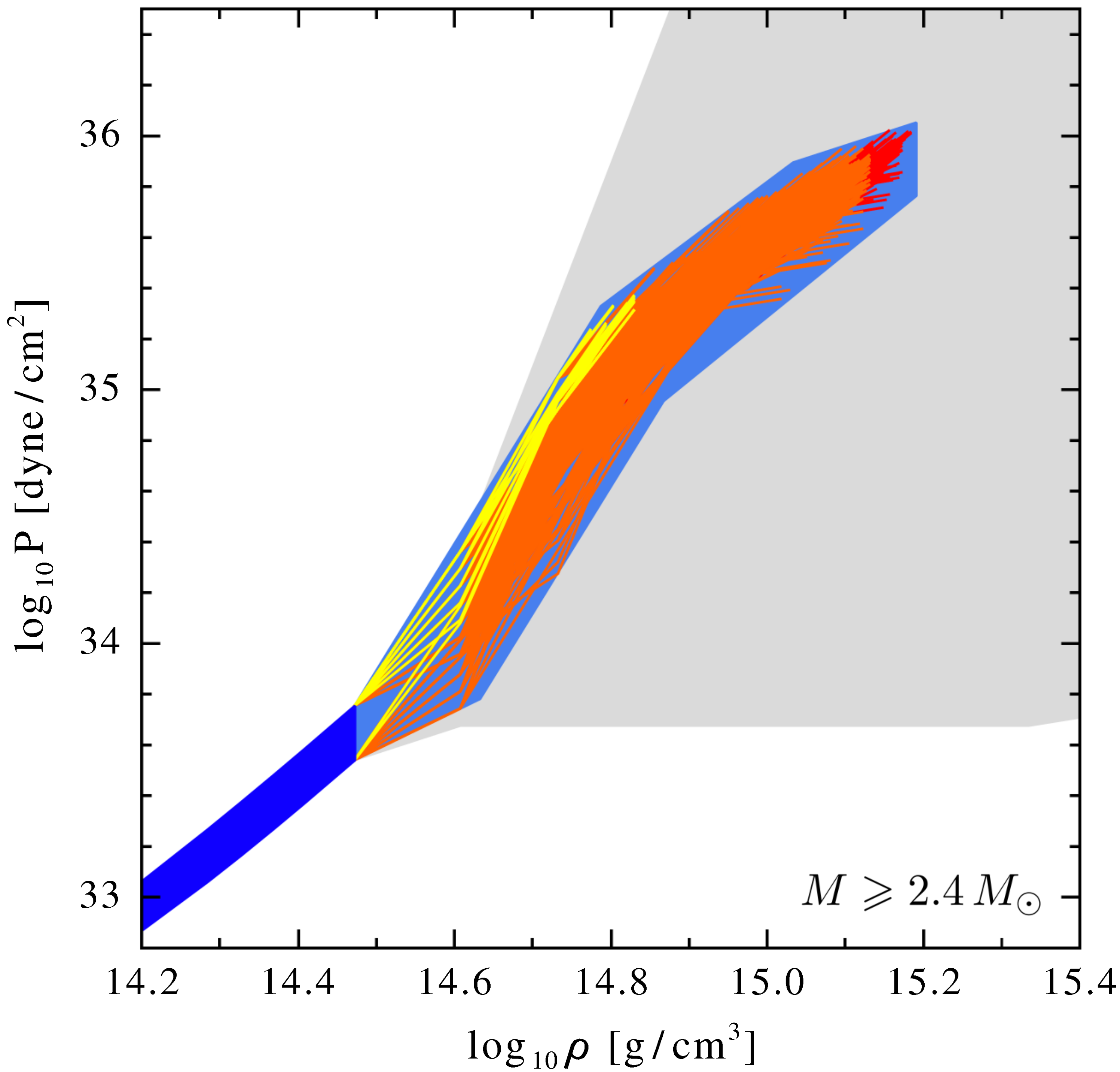
fictitious NS mass

$$M_{\max} > 2.4\,M_{\odot}$$

causality

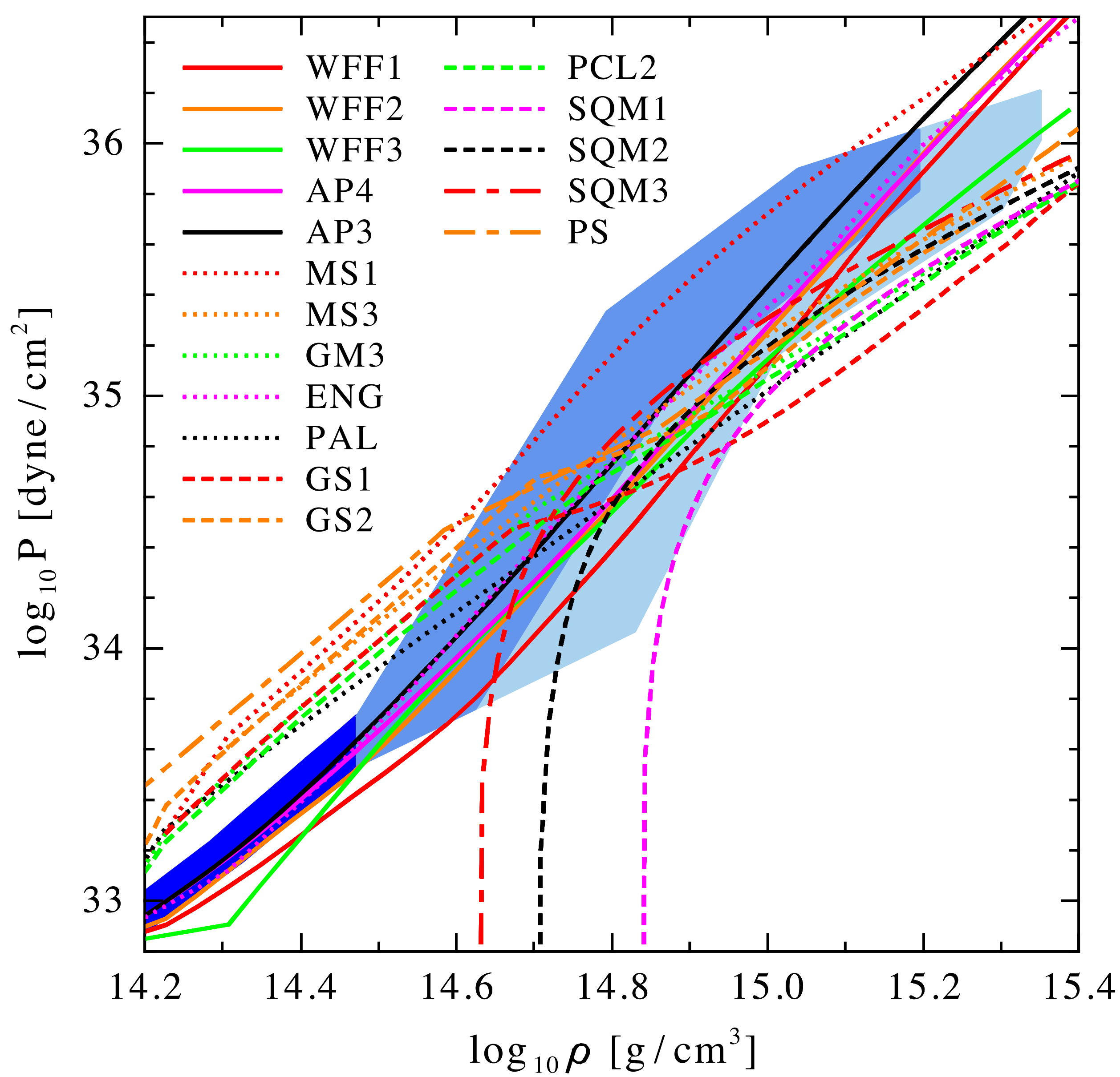
$$v_s(\rho) = \sqrt{dP/d\varepsilon} < c$$

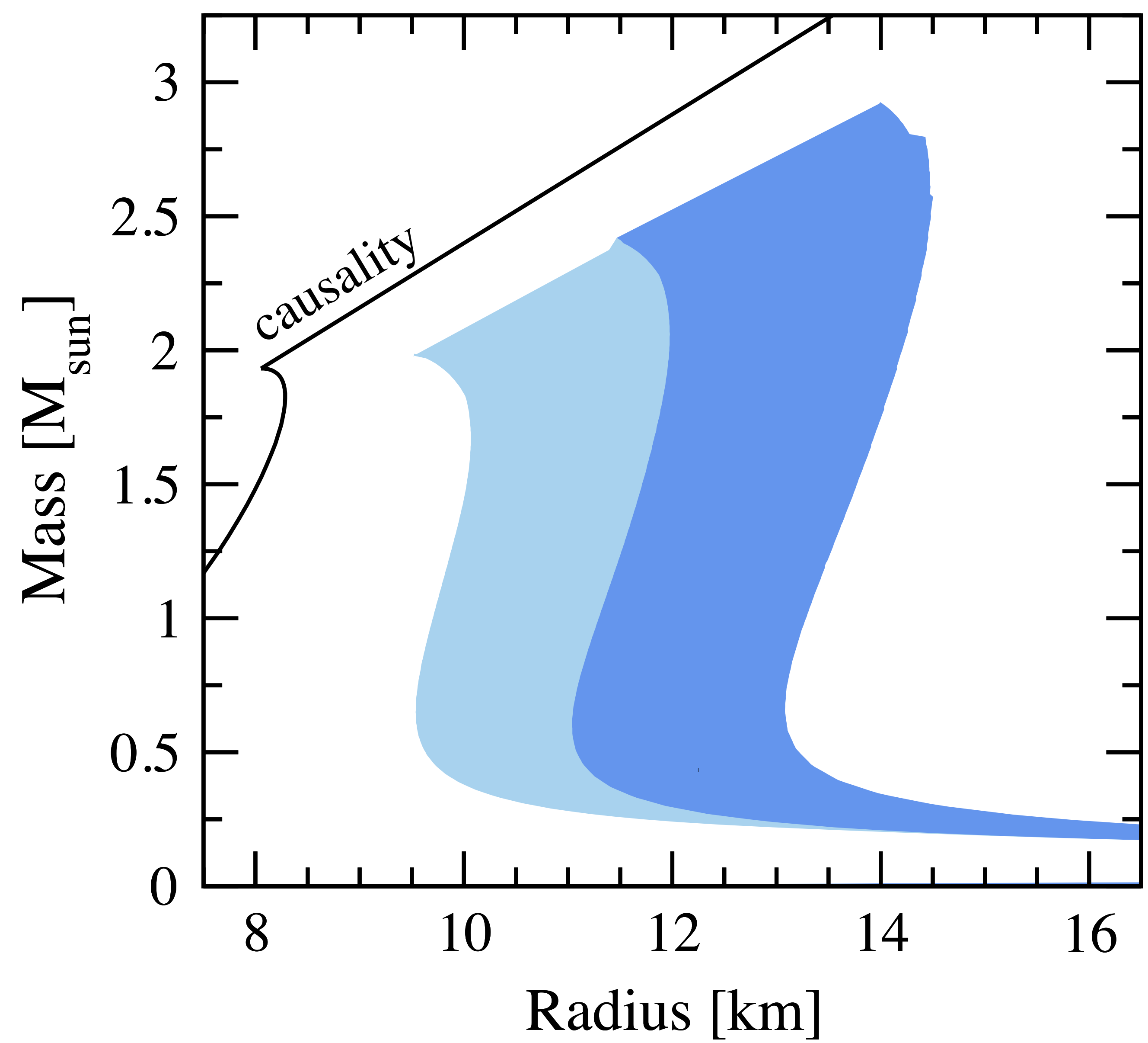
KH, Lattimer, Pethick, Schwenk, ApJ 773, 11 (2013)



increased $M_{\max}$ systematically reduces width of band

Constraints on neutron star radii

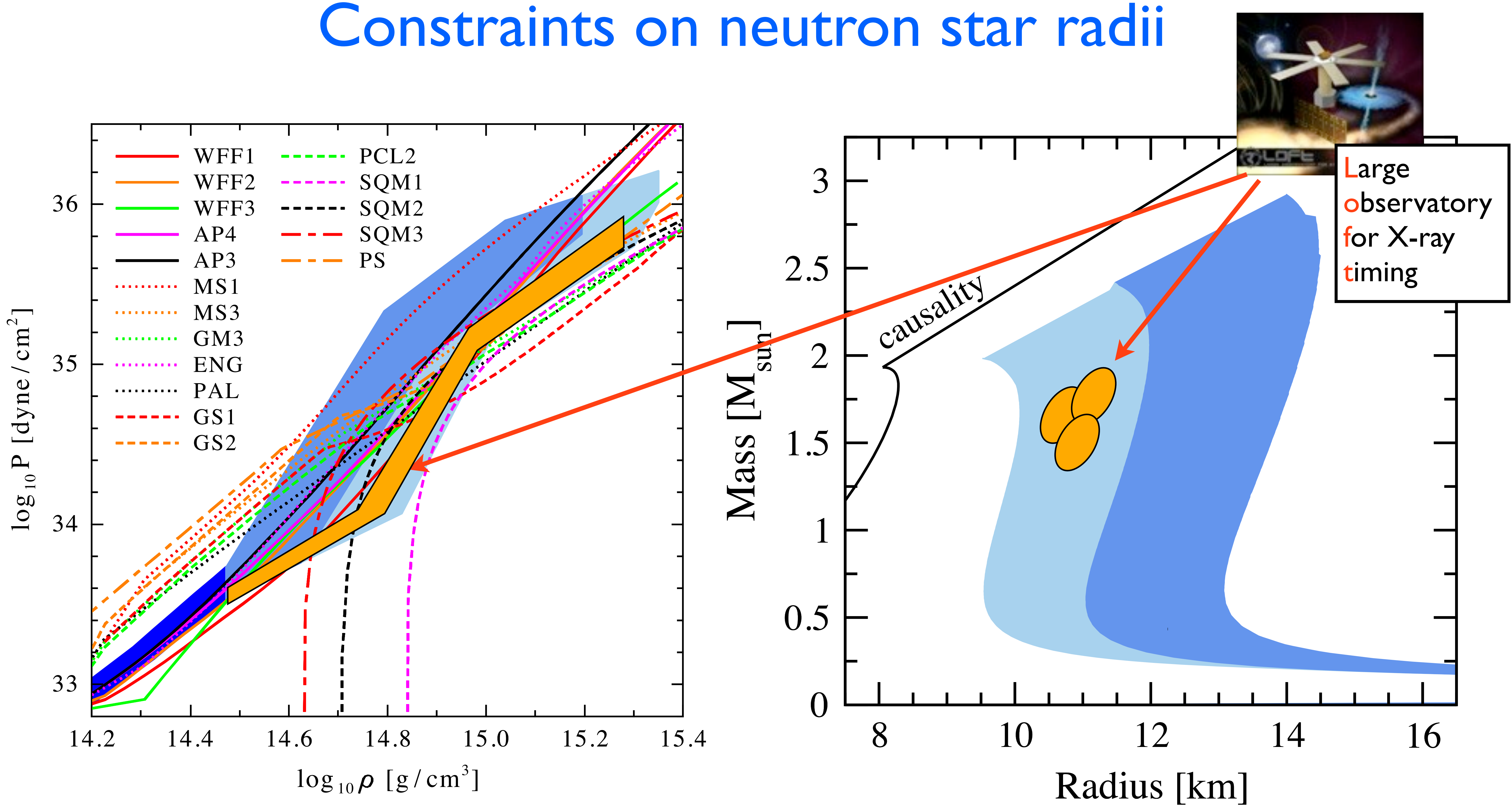




KH, Lattimer, Pethick, Schwenk, ApJ 773, 11 (2013)
see also KH, Lattimer, Pethick, Schwenk, PRL 105, 161102 (2010)

- low-density part of EOS sets scale for allowed high-density extensions
- current radius prediction for typical $1.4\,M_\odot$ neutron star: $9.7 - 13.9\,\mathrm{km}$

Constraints on neutron star radii



KH, Lattimer, Pethick, Schwenk, ApJ 773, 11 (2013)
see also KH, Lattimer, Pethick, Schwenk, PRL 105, 161102 (2010)

- low-density part of EOS sets scale for allowed high-density extensions
- current radius prediction for typical $1.4\,M_\odot$ neutron star: $9.7 - 13.9\,\mathrm{km}$
- radius measurements could significantly improve constraints

Summary

• recent advances allow ab initio studies of medium-mass nuclei

• remarkable agreement between different methods for given interaction, uncertainties dominated by differences in nuclear interactions

• results presented for properties of neutron-rich nuclei and matter based on sets of current chiral EFT NN+3N interactions

Future directions

• derivation of systematic uncertainty estimates for many-body observables, order-by-order convergence studies

• exploration of different fitting strategies, include bayesian analysis for statistical interpretation of uncertainties?

• role of regulators, clean separation of short- and long-range physics, naturalness of coupling constants, power counting schemes, inclusion of delta excitations...

LENPIC

In collaboration with:

C. Drischler, T. Krüger, R. Roth, A. Schwenk

R. Furnstahl, S. More

S. Bogner

E. Epelbaum, H. Krebs

A. Gezerlis

A. Nogga

J. Lattimer

C. Pethick

J. Golak, R. Skibinski

G. Hagen, T. Papenbrock

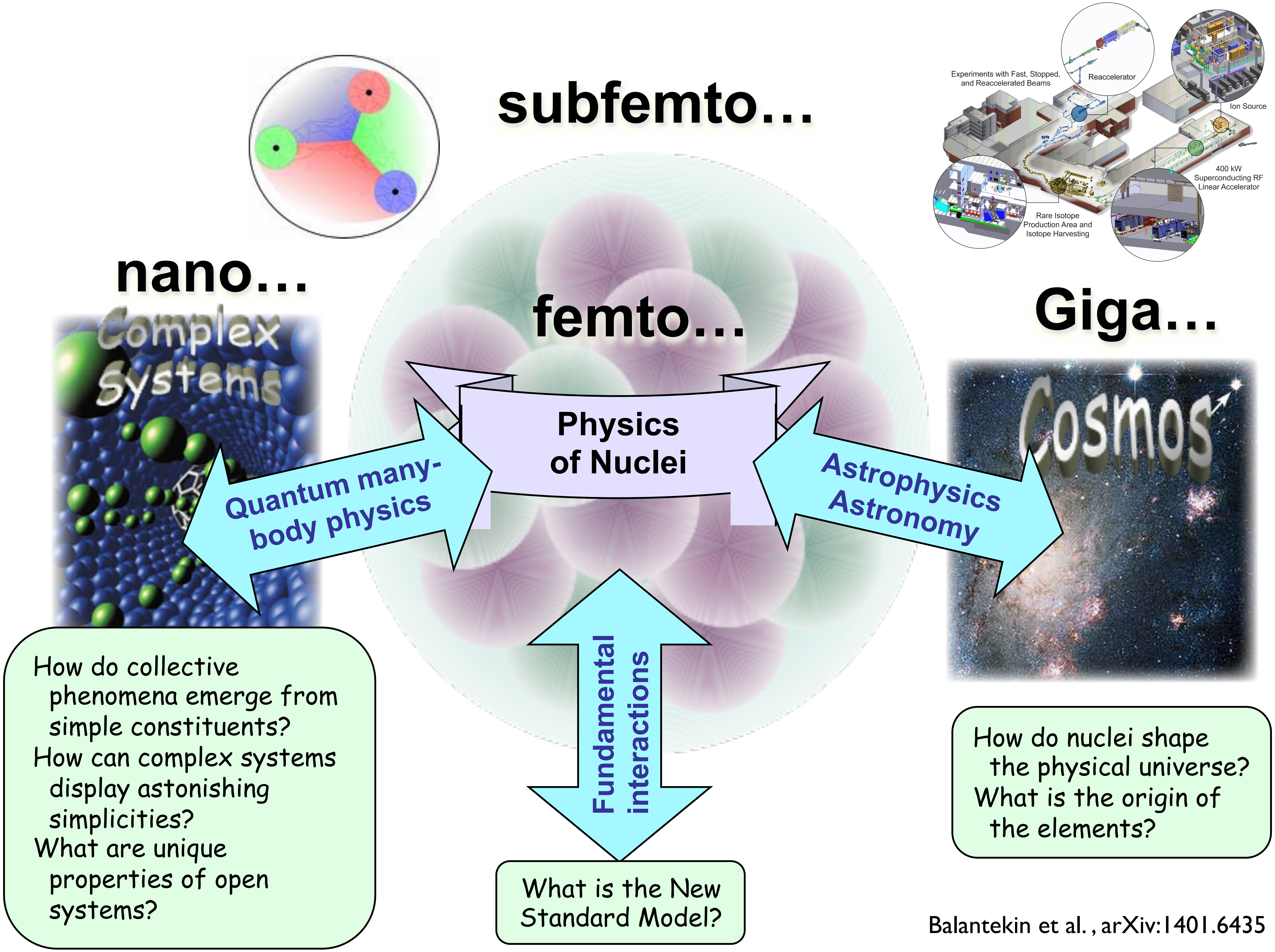
international collaborator in

---

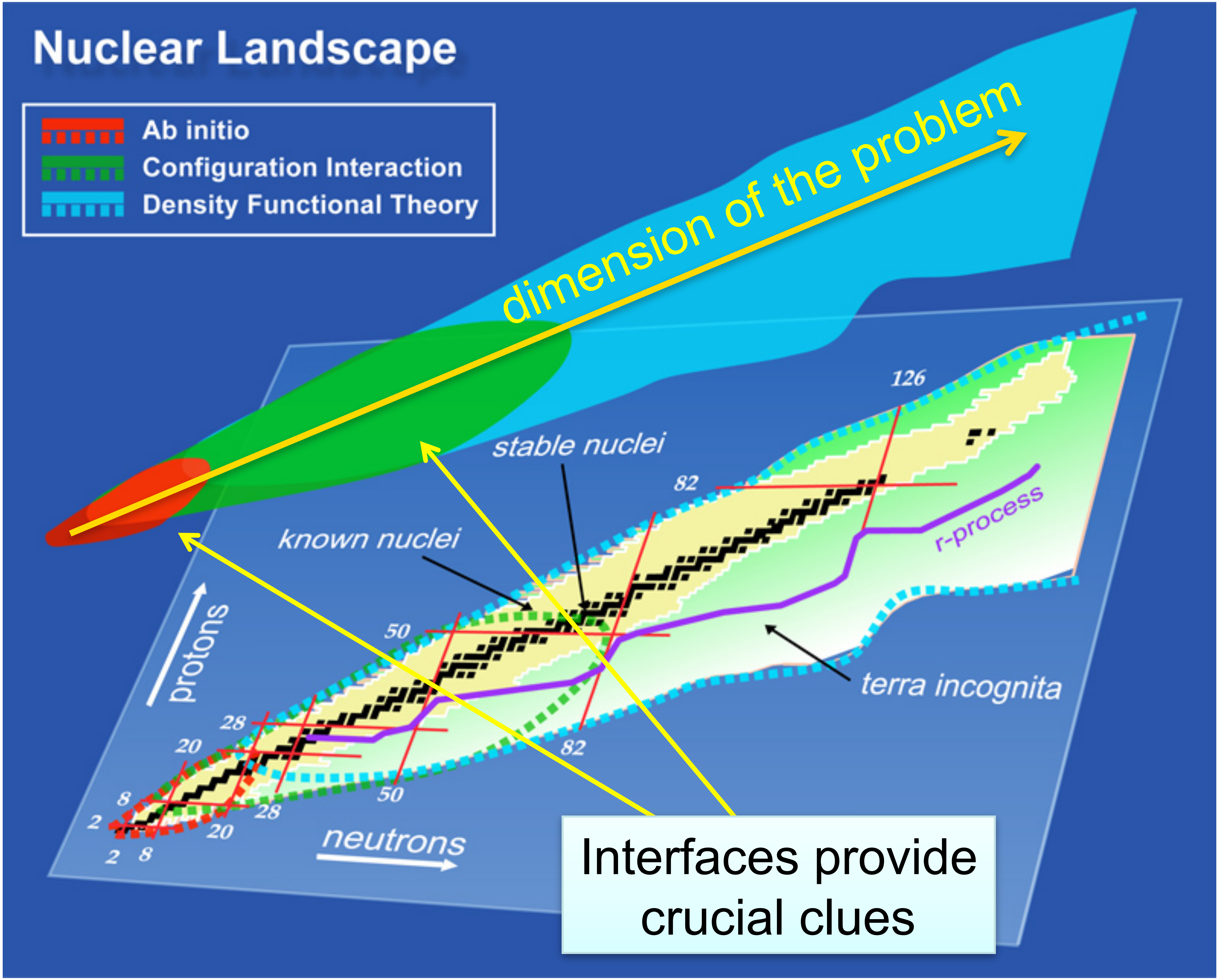
computing support:

Thank you!

New frontiers from rare isotope facilities

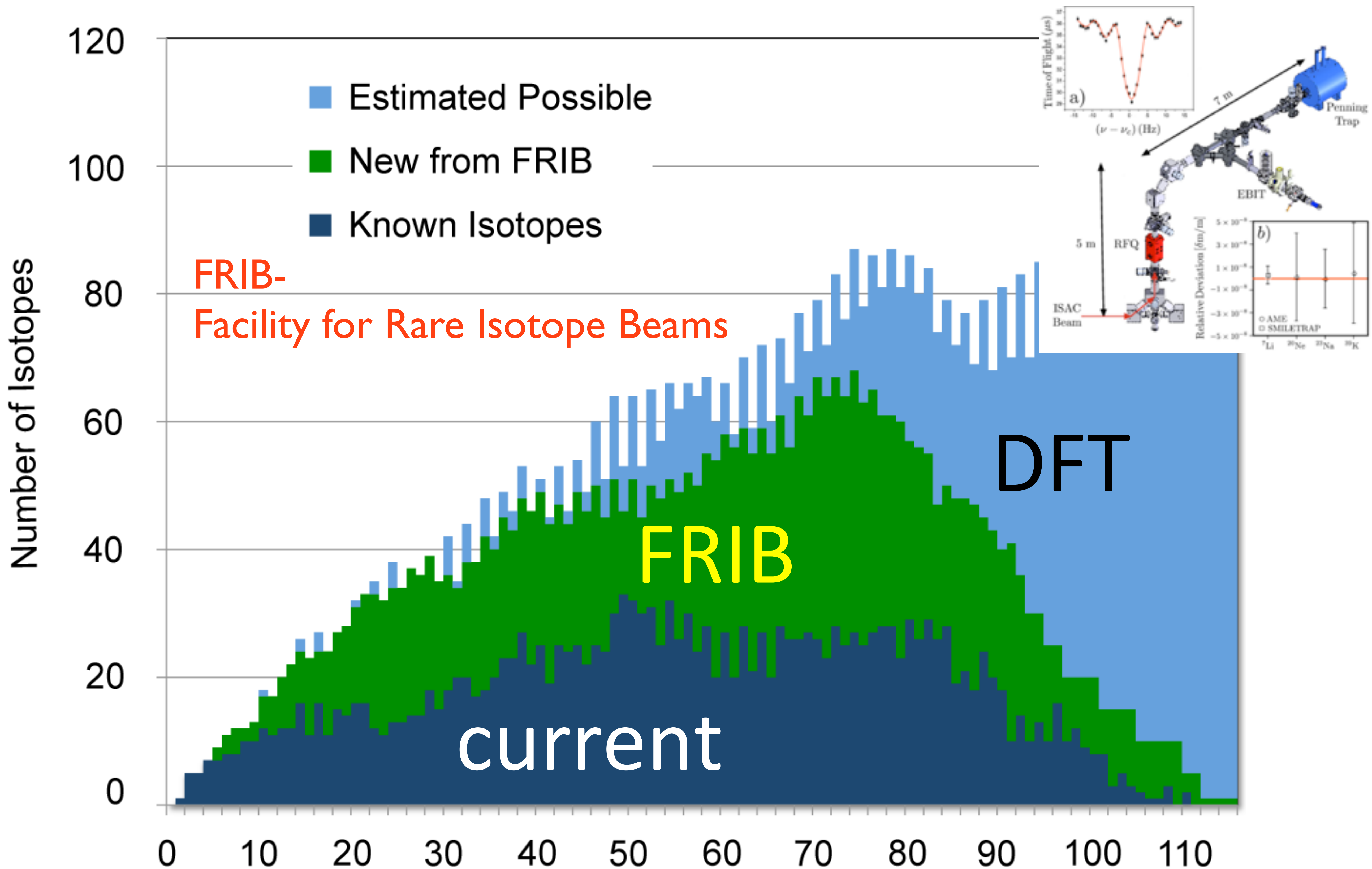


Balantekin et al., arXiv:1401.6435



Nuclear Landscape
Ab initio
Configuration Interaction
Density Functional Theory
dimension of the problem
stable nuclei
known nuclei
126
82
50
28
20
8
2
r-process
terra incognita
protons
neutrons
Interfaces provide crucial clues

New frontiers from rare isotope facilities



Balantekin et al. , arXiv:1401.6435

Exciting recent developments on many fronts...

LETTER

nature

doi:10.1038/nature12522

Evidence for a new nuclear 'magic number' from the level structure of ^{54}Ca

D. Steppenbeck[1], S. Takeuchi[2], N. Aoi[3], P. Doornenbal[2], M. Matsushita[1], H. Wang[2], H. Baba[2], N. Fukuda[2], S. Go[1], M. Honma[4], J. Lee[2], K. Matsui[5], S. Michimasa[1], T. Motobayashi[2], D. Nishimura[6], T. Otsuka[1,5], H. Sakurai[2,5], Y. Shiga[7], P.-A. Söderström[2], T. Sumikama[8], H. Suzuki[2], R. Taniuchi[5], Y. Utsuno[9], J. J. Valiente-Dobón[10] & K. Yoneda[2]

LETTER

nature

doi:10.1038/nature11188

The limits of the nuclear landscape

Jochen Erler[1,2], Noah Birge[1], Markus Kortelainen[1,2,3], Witold Nazarewicz[1,2,4], Erik Olsen[1,2], Alexander M. Perhac[1] & Mario Stoitsov[1,2]‡

LETTER

nature

doi:10.1038/nature12226

Masses of exotic calcium isotopes pin down nuclear forces

F. Wienholtz[1], D. Beck[2], K. Blaum[3], Ch. Borgmann[3], M. Breitenfeldt[4], R. B. Cakirli[3,5], S. George[1], F. Herfurth[2], J. D. Holt[6,7], M. Kowalska[8], S. Kreim[3,8], D. Lunney[9], V. Manea[9], J. Menéndez[6,7], D. Neidherr[2], M. Rosenbusch[1], L. Schweikhard[1], A. Schwenk[7,6], J. Simonis[6,7], J. Stanja[10], R. N. Wolf[1] & K. Zuber[10]

LETTER

nature

doi:10.1038/nature09466

A two-solar-mass neutron star measured using Shapiro delay

P. B. Demorest[1], T. Pennucci[2], S. M. Ransom[1], M. S. E. Roberts[3] & J. W. T. Hessels[4,5]

RESEARCH ARTICLE SUMMARY

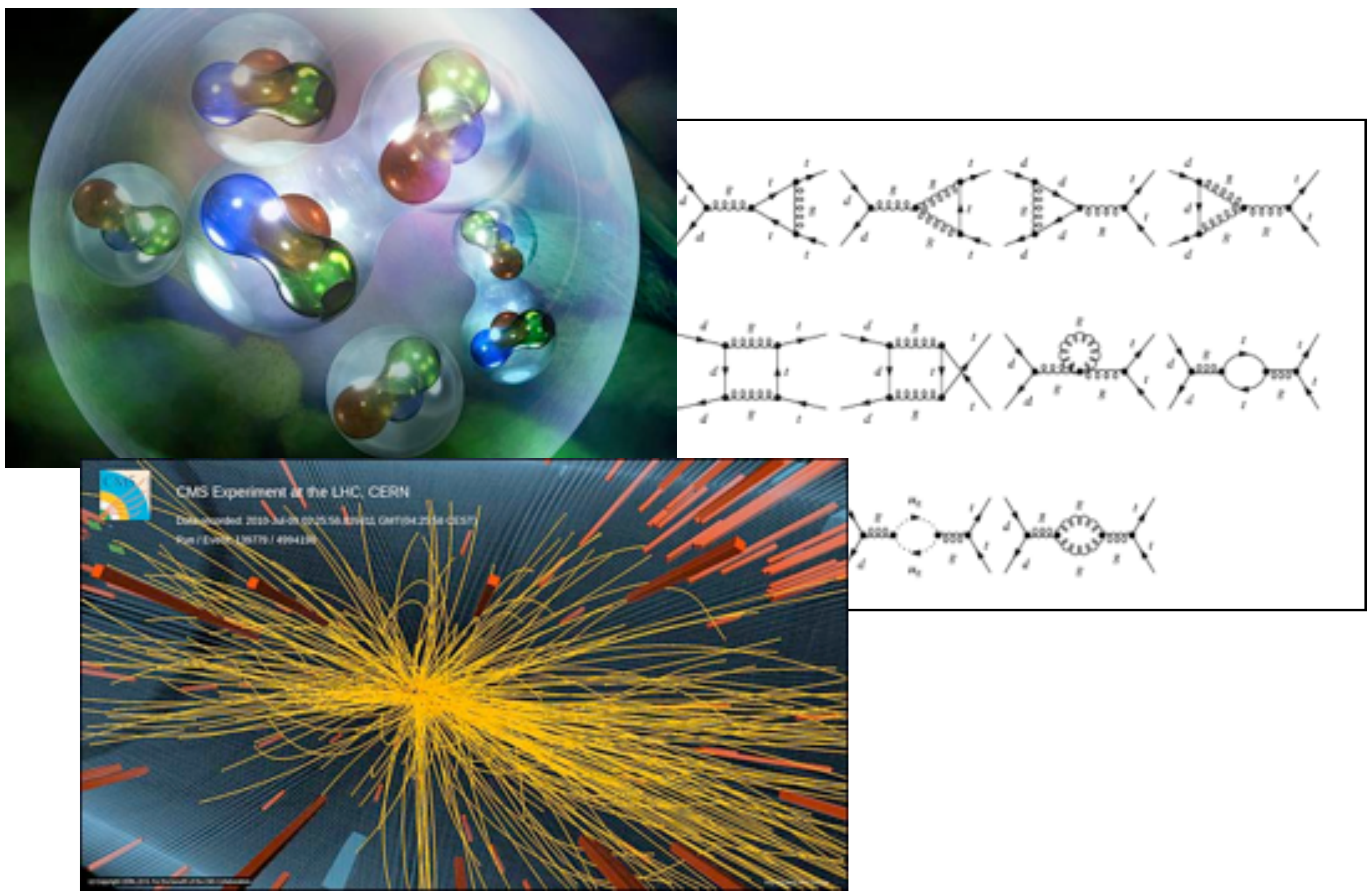
Science

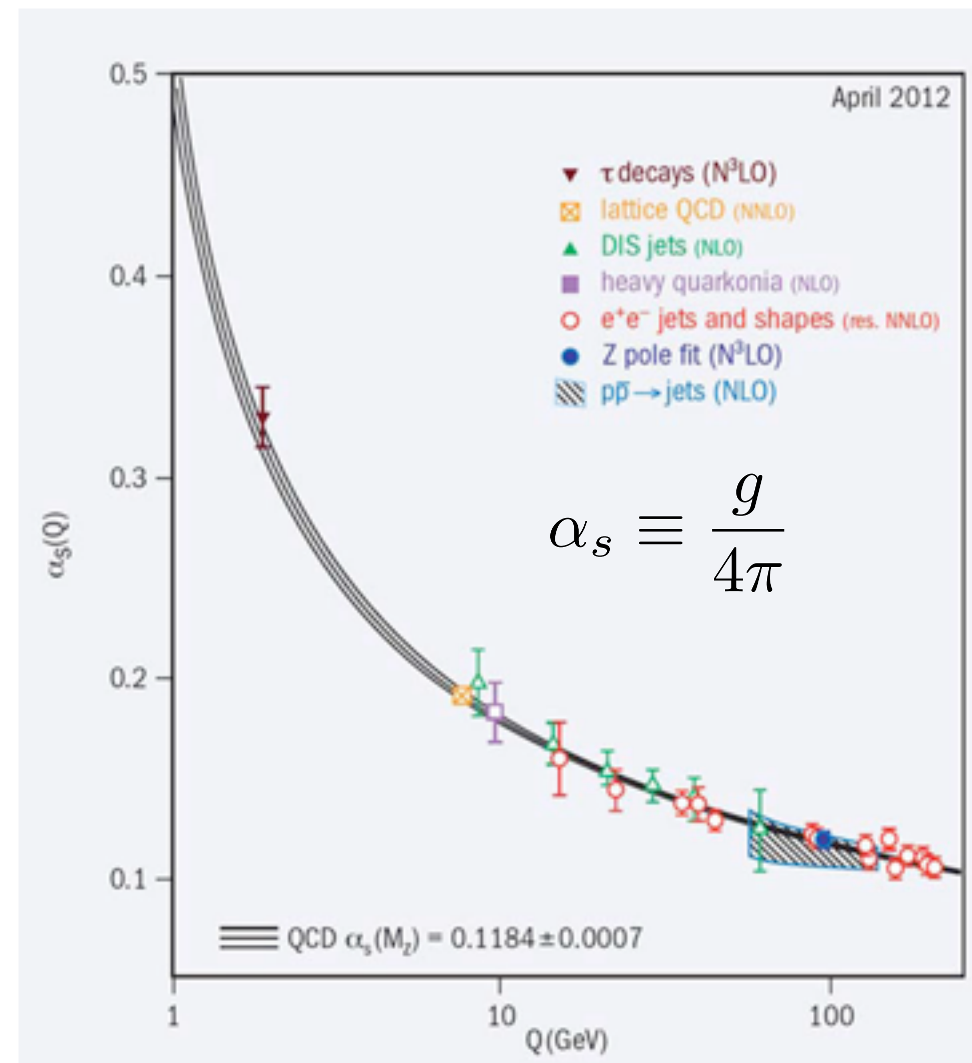
A Massive Pulsar in a Compact Relativistic Binary

John Antoniadis,* Paulo C. C. Freire, Norbert Wex, Thomas M. Tauris, Ryan S. Lynch, Marten H. van Kerkwijk, Michael Kramer, Cees Bassa, Vik S. Dhillon, Thomas Driebe, Jason W. T. Hessels, Victoria M. Kaspi, Vladislav I. Kondratiev, Norbert Langer, Thomas R. Marsh, Maura A. McLaughlin, Timothy T. Pennucci, Scott M. Ransom, Ingrid H. Stairs, Joeri van Leeuwen, Joris P. W. Verbiest, David G. Whelan

Theory of the strong interaction: Quantum chromodynamics

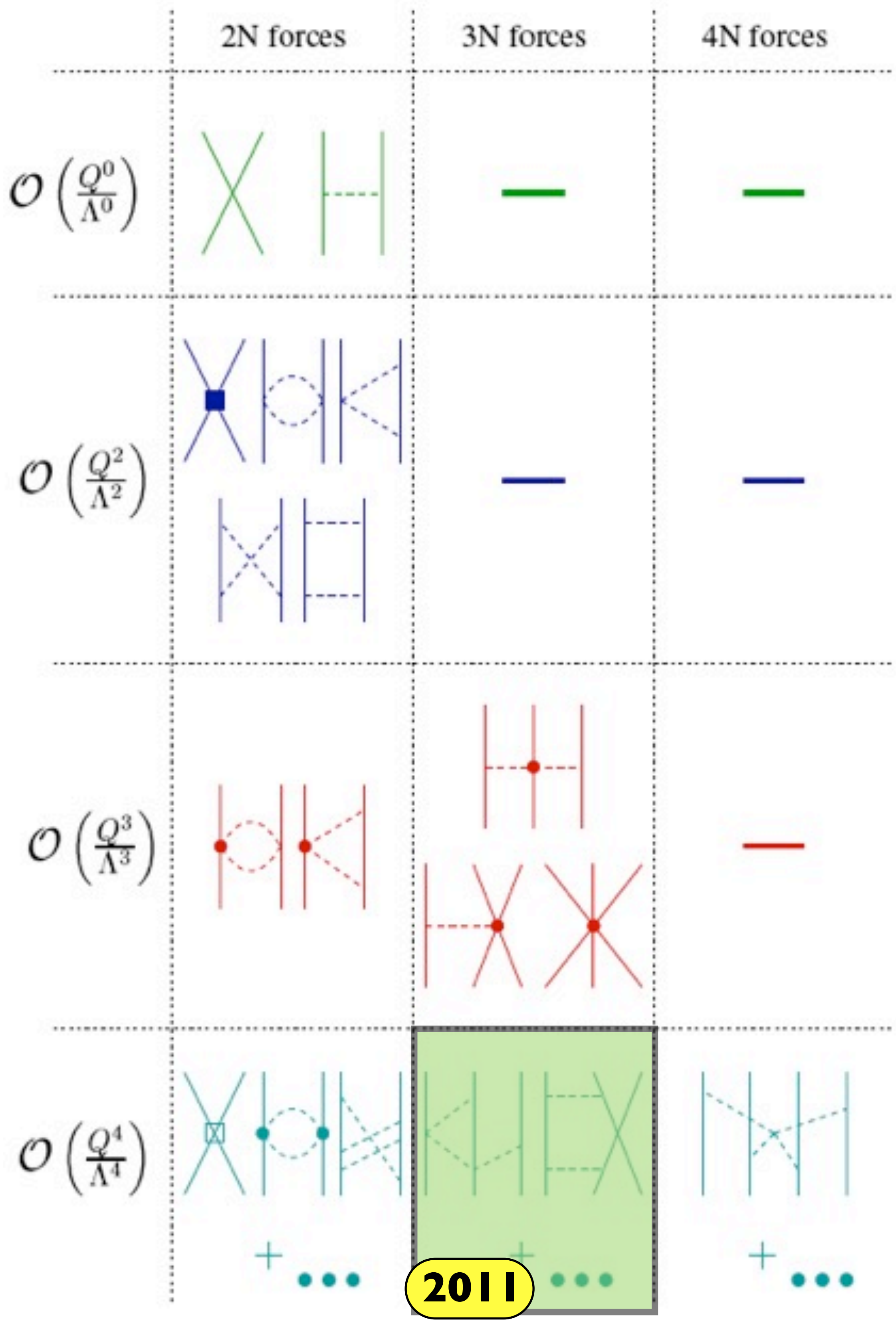
$$\mathcal{L}_{QCD} = -\frac{1}{4} F^a_{\mu\nu} F^{a\mu\nu} + \overline{q}(i\gamma^\mu \partial_\mu - m)q + g\overline{q}\gamma^\mu T_a q A^a_\mu$$

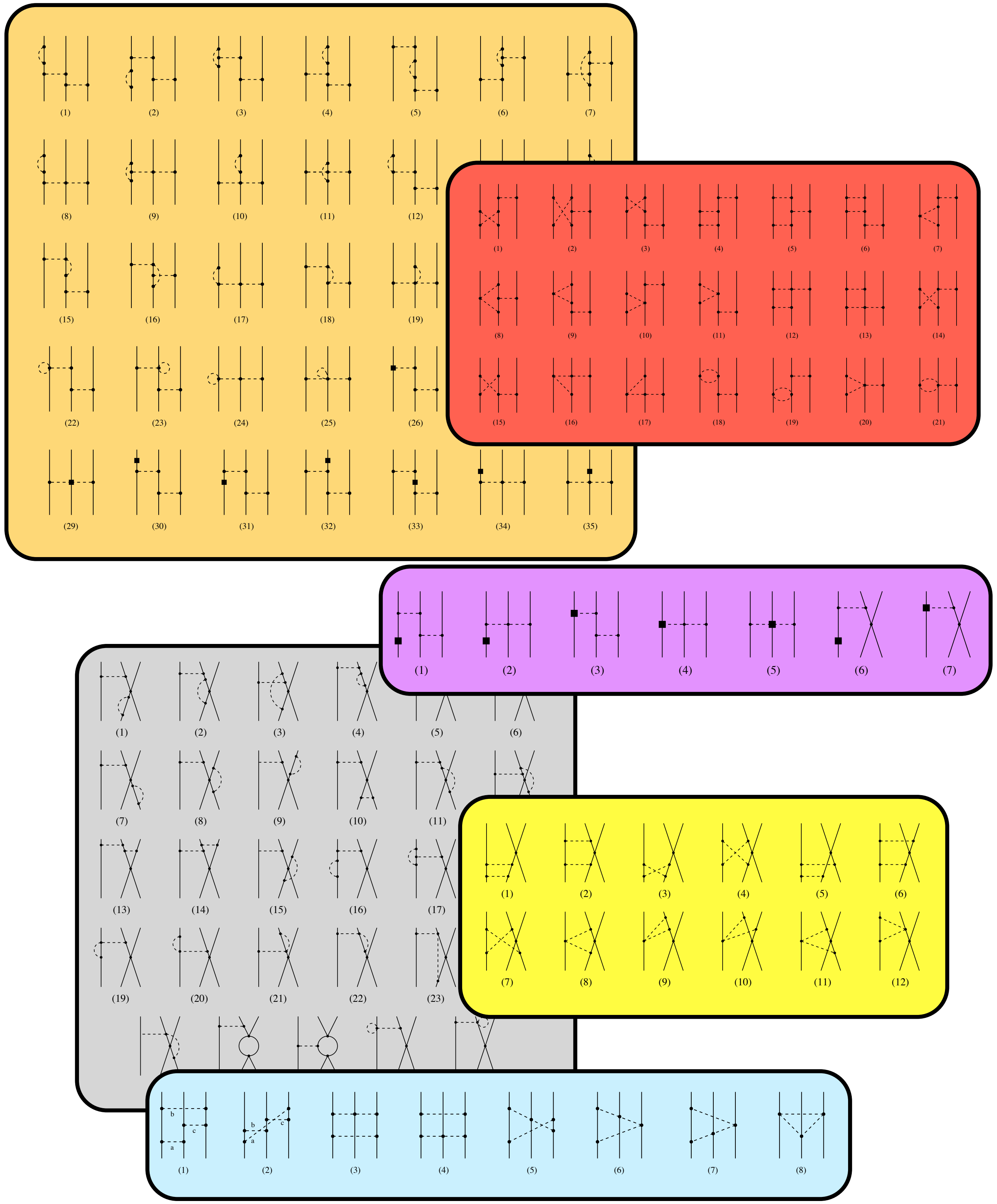




- theory perturbative at high energies
- highly non-perturbative at low energies

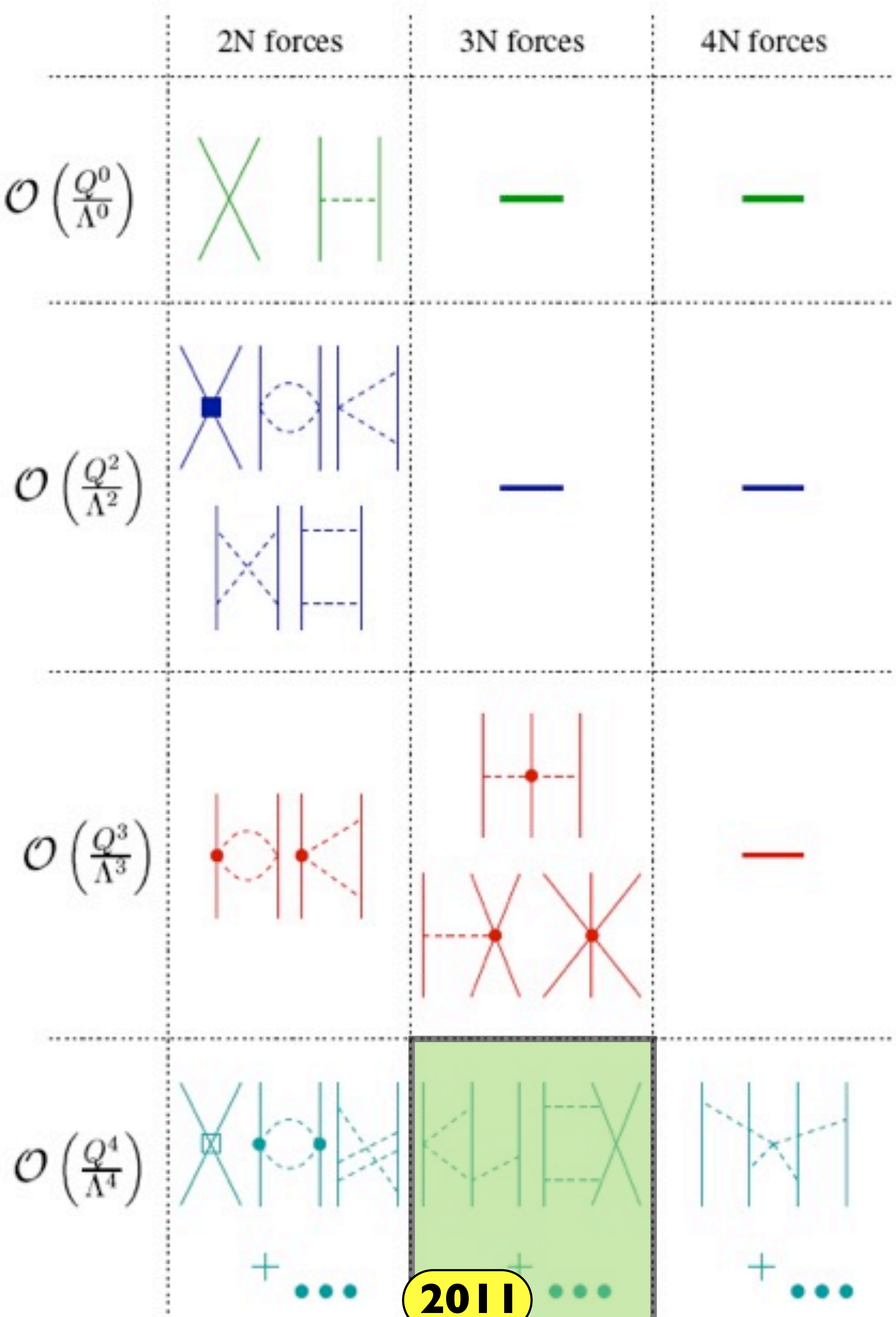
Chiral 3N forces at subleading order (N^3LO)

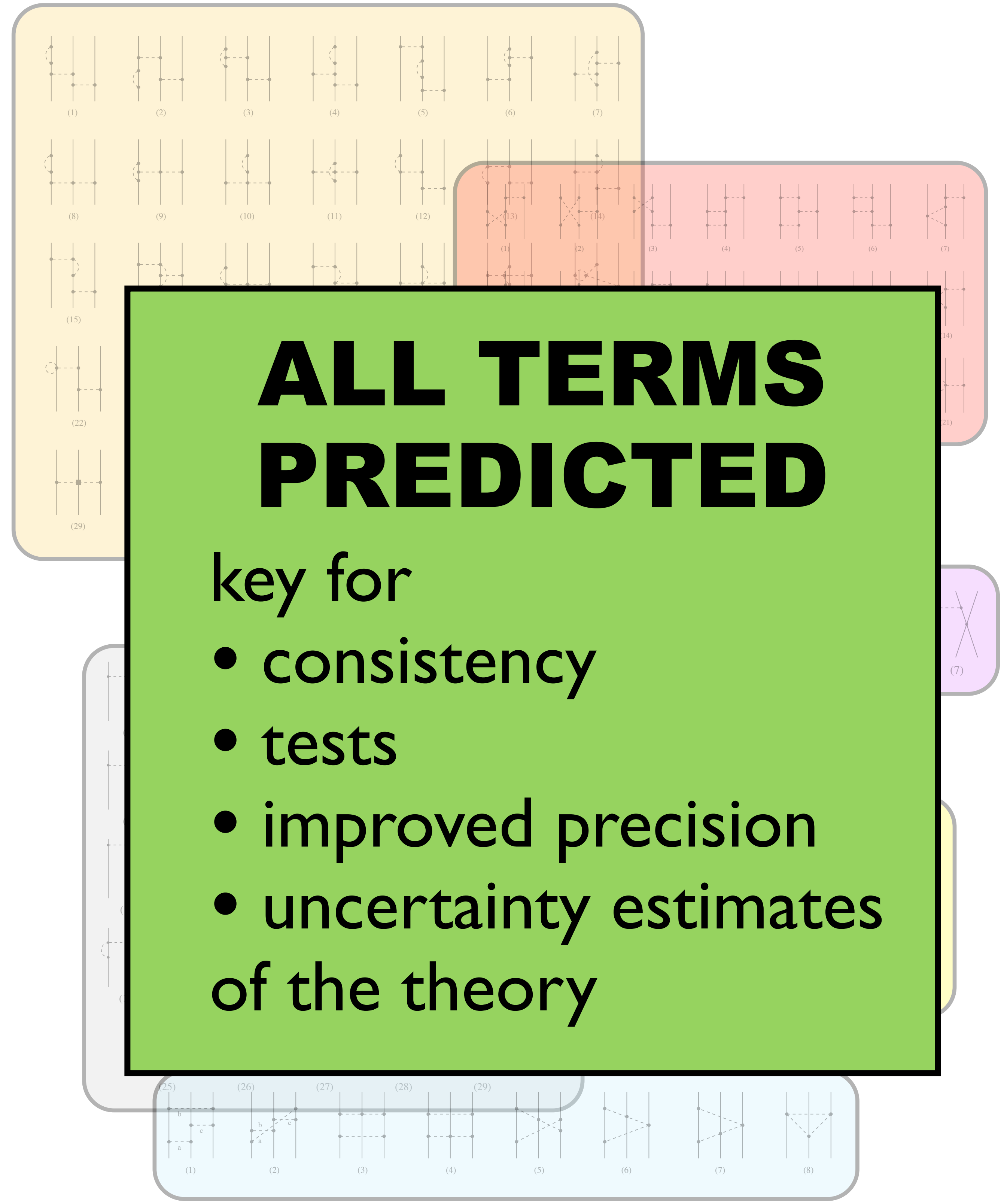




Bernard et al., PRC 77, 064004 (2008)
Bernard et al., PRC 84, 054001 (2011)

Chiral 3N forces at subleading order (N^3LO)





ALL TERMS PREDICTED

key for

- consistency
- tests
- improved precision
- uncertainty estimates

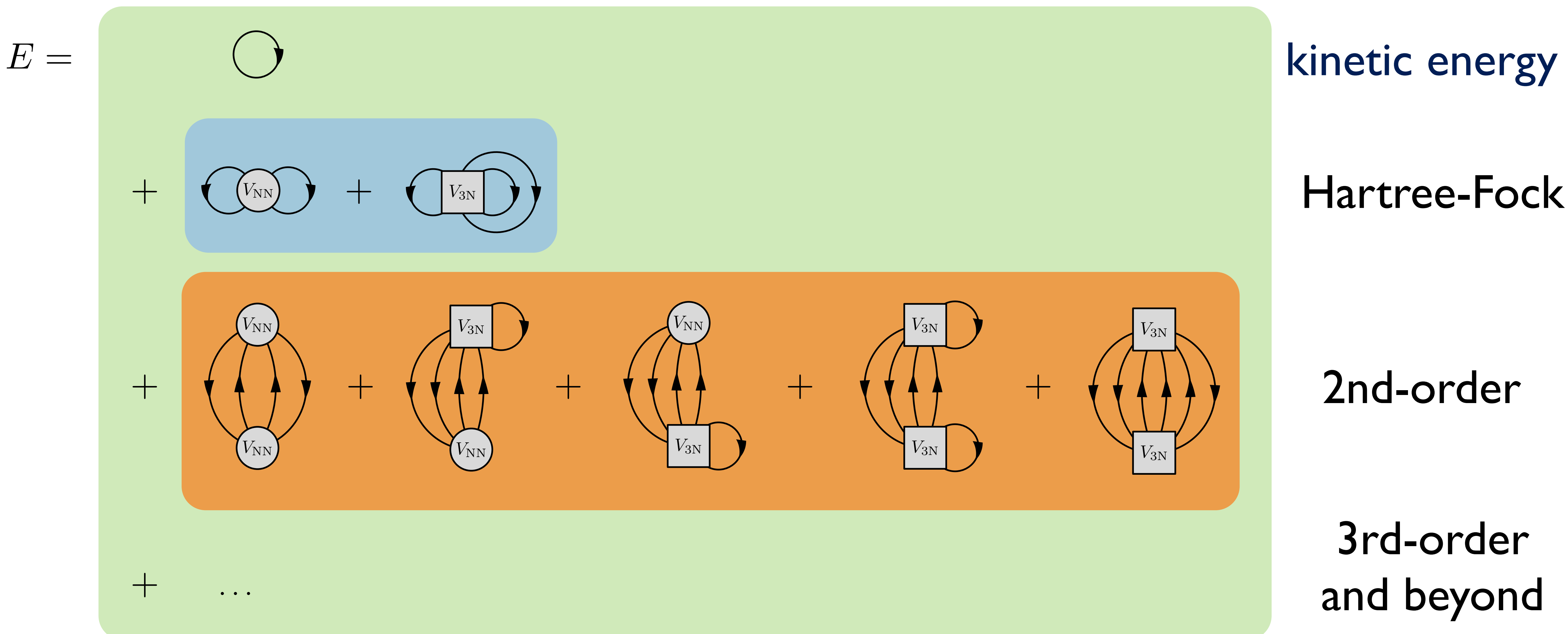
of the theory

Bernard et al., PRC 77, 064004 (2008)
Bernard et al., PRC 84, 054001 (2011)

Equation of state: Many-body perturbation theory

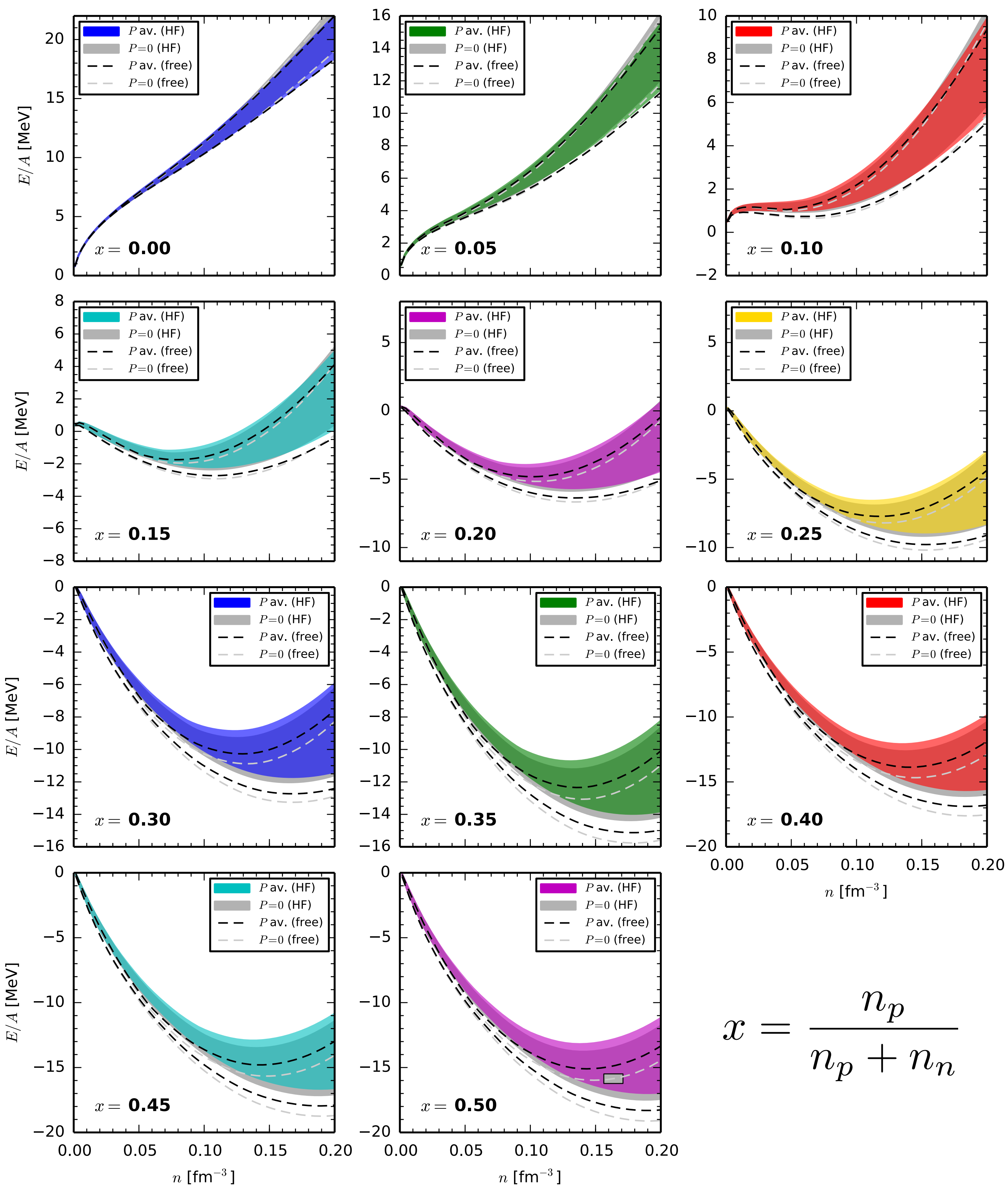
central quantity of interest: energy per particle E/N

$$H(\lambda) = T + V_{\mathrm{NN}}(\lambda) + V_{3\mathrm{N}}(\lambda) + \ldots$$



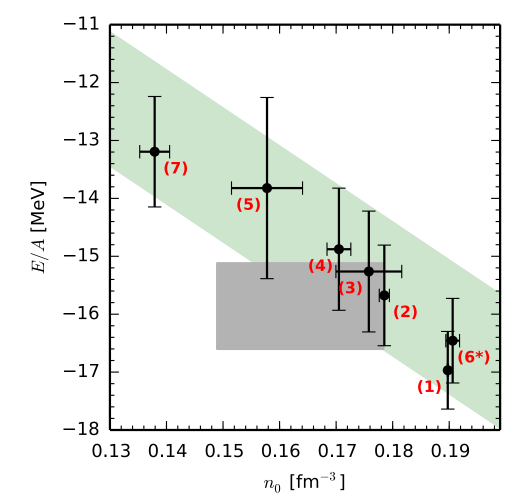
- "hard" interactions require non-perturbative summation of diagrams
- with low-momentum interactions much more perturbative
- inclusion of 3N interaction contributions crucial!

First application to isospin asymmetric nuclear matter



$$x = \frac{n_p}{n_p + n_n}$$

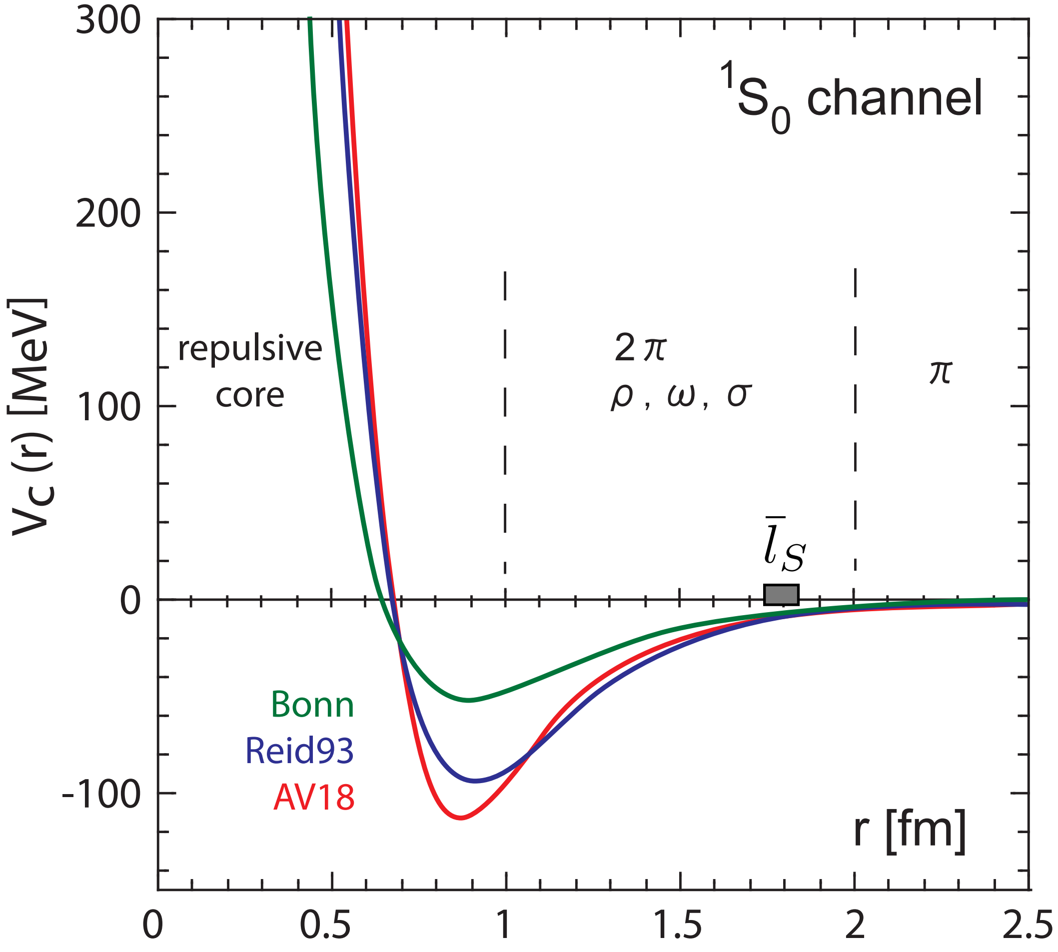
- uncertainty bands determined by set of 7 Hamitonians

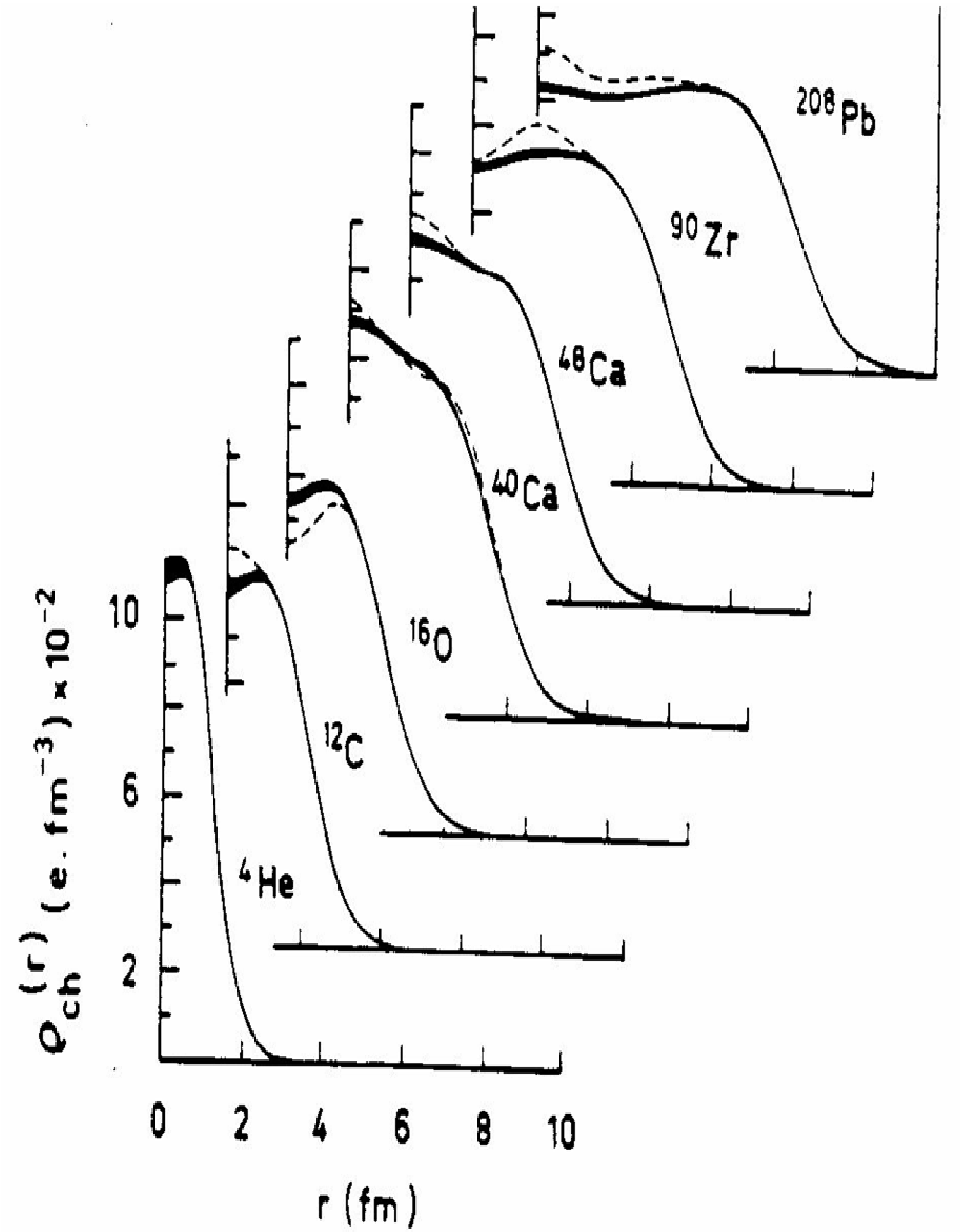


- many-body framework allows treatment of any decomposed 3N interaction

Drischler, KH, Schwenk,
in preparation

Equation of state of symmetric nuclear matter, nuclear saturation

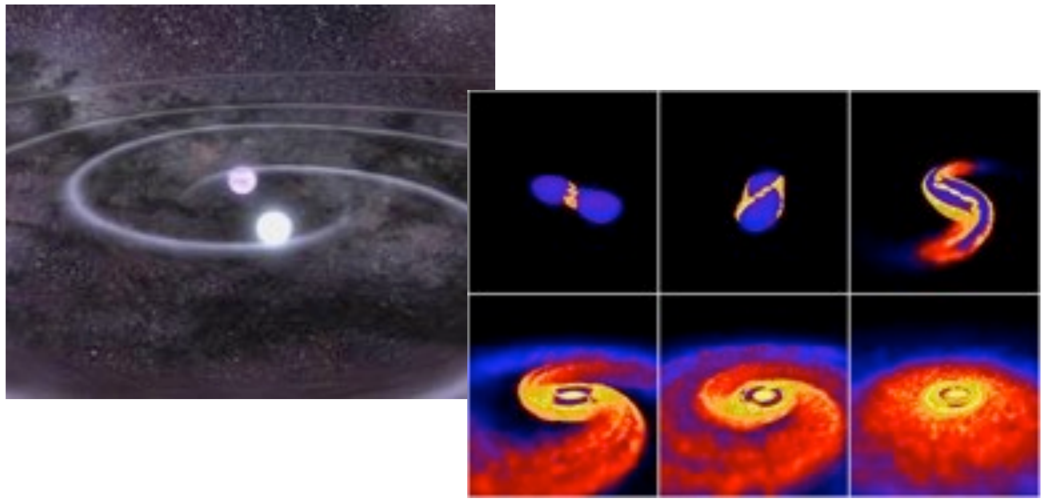


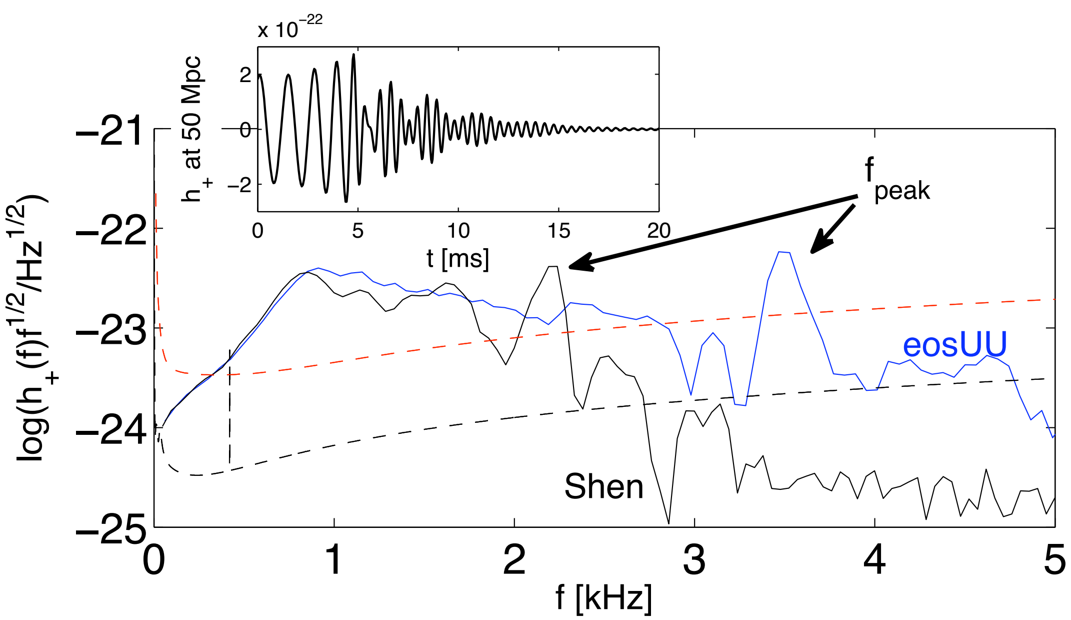


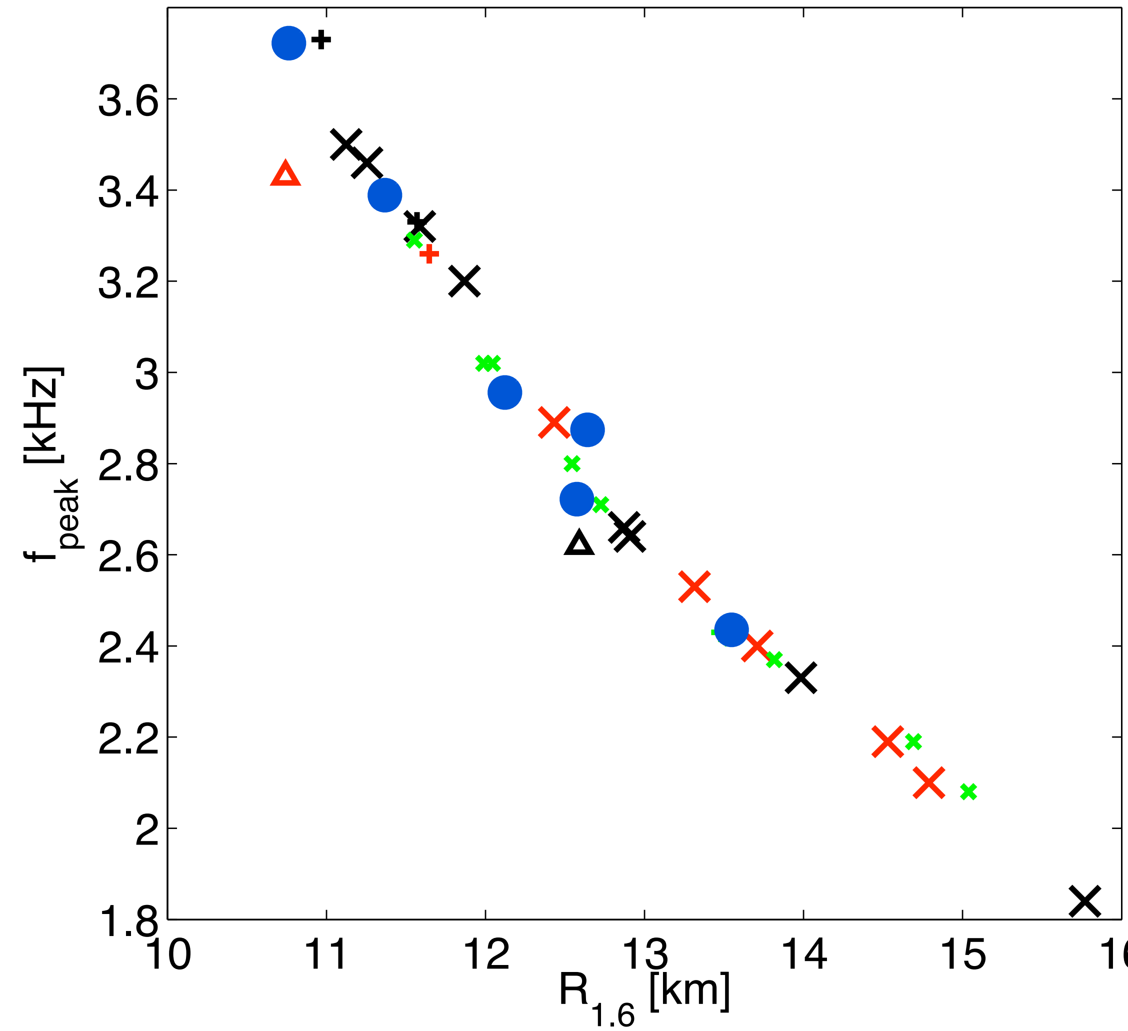
"Very soft potentials must be excluded because they do not give saturation; they give too much binding and too high density. In particular, a substantial tensor force is required."

Hans Bethe (1971)

Gravitational wave signals from neutron star binary mergers







Bauswein and Janka, PRL 108, 011101 (2012),
Bauswein, Janka, KH, Schwenk, PRD 86, 063001 (2012)

- simulations of NS binary mergers show strong correlation between between $f_{\rm peak}$ of the GW spectrum and the radius of a NS
- measuring $f_{\rm peak}$ is key step for constraining EOS systematically at large ρ